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February 1987

The Fe XIV Solar Corona at 5303 Angstroms:

An Atlas of Synoptic Charts from the Sacramento Peak Coronal Photometer May 1973-December 1985

JUN 1 1 1987

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February 1987

The Fe XIV Solar Corona at 5303 Angstroms:

An Atlas of Synoptic Charts from the Sacramento Peak Coronal Photometer May 1973-December 1985

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PREFACE

Synoptic photoelectric observations of the solar corona in the emission from the green line (5303Å) of Fe XIV were begun at the Sacramento Peak Observatory in June 1973 and continue to the present. The observations made during this program provide a record of the distribution and brightness of the hotter regions of the corona (1.8 x 10⁶ K) over an entire solar cycle. As such, the green line data may provide a link between the active regions of the chromosphere and the large scale structure of the corona. In order to allow comparison of these data with other observations of the sun and corona, we present here an atlas of the green line measurements in the same format as that used to display the white light coronal structure recorded by the High Altitude Observatory's K-coronameters on Mauna Loa, in previously published atlases. It is our intent that this atlas be brought up to date periodically as the observations permit.

R. C. Altrock, L. B. Gilliam

National Solar Observatory,

D. G. Sime, R. R. Fisher

High Altitude Observatory

September 1986.

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I. INTRODUCTION

The purpose of this Note is to describe and present synoptic observations of the solar corona, made over the last solar cycle, in the green Fe XIV line at 5303 Å. This is attempted in a way which displays both individual regions of bright emission and the global properties of the emission. The long string of observations, 1973 to 1984, is divided into and displayed in individual Carrington rotations, allowing the study of the global coronal emission structure at any given time as well as permitting the long term variability in the coronal emission to be seen. Further, the format permits easy comparison with other data sets so that the scientific potential of the observations can be achieved more easily.

Since the green line emission arises in coronal material with temperatures of about 1.8 x 10⁶ K it is widely thought to trace out those parts of the corona heated by chromospheric active regions. Solar coronal features of a variety of scales, including coronal holes (Fisher and Musman, 1975), have also been detected in this line, suggesting that the record of green line observations can be viewed in a sense as a synthesis of medium scale chromospheric and semi-global coronal properties. However, the overall relationship of regions which are bright in 5303 Å to other solar coronal structures has never fully been examined. With this Note, we provide a tool to explore this relationship more completely.

The large scale structure of the solar corona as seen in white light has been observed for the last 20 years with a series of K-coronameters operated at the HAO Mauna Loa, Hawaii, site. The data from the last of these instruments (Fisher et al., 1981) have also been displayed in a series of atlases (Fisher et al., 1985) which show the distribution of the coronal white light polarized brightness with latitude and longitude. These data, from which coronal electron density can quite directly be inferred, have been presented in maps displaying one Carrington rotation at a time. Here we present, in the same format, the record of the distribution of the green line radiance in latitude and longitude.

II. DESCRIPTION OF THE INSTRUMENTATION

The data presented in this Note were gathered with the National Solar Observatory/Sacramento Peak photoelectric coronal photometer operating at the Coudé feed of the 40 cm coronagraph. A mica filter with full width at half maximum (FWHM) of 0.65 Å is presently used to isolate the Fe XIV line at 5303 Å, but up until 1983 January 4, a birefringent filter of FWHM 0.58 Å was used. A piezo-electric modulator is used to chop between the central bandpass and the nearby continuum in order to remove the contribution of the continuum at this wavelength.

Data are collected by means of a scan in position angle at one of a selection of heights above the limb. These heights are usually $0.15~R_{\odot}$ and $0.35~R_{\odot}$ with an additional scan at 0.25, 0.45 or $0.55~R_{\odot}$ as available. Data are gathered at points separated by 3° in position angle and a complete scan in position angle takes approximately 5 minutes. In this Note, we present only the data taken at $1.15~R_{\odot}$. The data are collected on magnetic tape for later analysis. Control of the instrument and of the data collection was handled initially by a DEC PDP 11-10 computer, and subsequently (after 1982 August 31) by a PE 3220. The instrument and its operation are described more fully by Fisher (1973) and by Smartt (1982). An example of the data collected is shown in Fisher and Musman (1975).

III. DATA REDUCTION AND PRODUCTS

Daily data scans have been examined for quality and freedom from artifacts and then collected on magnetic tape for further processing. The contents of these tapes, a record of the coronal brightness at 5303Å as a function of position angle at a given height above the limb, have then been reformatted to appear similar to the data collected by the High Altitude Observatory's (HAO's) Mk-III K-coronameter. This allows easy application of much of the software written to manipulate and analyze the white light data. In particular, for this Note the data have been reduced to the synoptic map format by the algorithms originally developed for the HAO synoptic white light coronal data collected at the Observatory's Mauna Loa, Hawaii, observing station.

The format of the display is similar to that used in the recent atlases of the white light data (Fisher et al., 1985). Each day's position angle scan from 0 to 180 degrees is assigned the longitude of the East limb of the sun for the time of the observation. The data are then inserted at the appropriate position in a rectangular grid representing solar latitude and longitude, organized into Carrington rotations. Similarly, the other part of the scan, from 180 to 360 degrees, is used to construct a map based on observations of the west limb.

The data are shown in the following section as rectangular maps for each Carrington rotation from May 1973 to December 1985. The data thus span Carrington rotations 1601 to 1770, except for rotations 1612 and 1629, which are missing. East limb data only are available for the interval covered by rotations 1601 to 1611, and while the data are rather sparse for other periods in 1974, 1975, and 1976, maps have been included in order to give a sense of the continuity, slow evolution and low intensity in the emission line corona. Special care should be taken in interpreting these data (see below).

The map based on East limb observations is shown at the top of the page, while that for West limb data appears at the bottom of the page. The day number of limb passage for each observation is indicated below the maps in the appropriate location, while days when no data are present are marked by a cross instead of a day number. For simplicity, only even day numbers are shown, although every day on which no data are available is indicated.

Note that where data are absent, the contour plotting routine has been allowed to interpolate across the gap in order to preserve some indication of the nearby large scale structure. An exception occurs when a data gap exists near the end of a calendar year. In such a case, no interpolation is carried out and a gap appears on the synoptic charts. These periods should be used only with great caution and, especially when comparisons are made to other solar data, care should be taken to see that data actually exist. However, in spite of these periods of sparse data coverage and the difficulties outlined below, we believe the maps give a reasonable indication of the true distribution of brightness in the corona. On each page, a legend is shown which gives the values of radiance in the observing bandpass (in millionths of the brightness of the solar disc at 5303 Å) for each of the shading and contour levels shown on that page. Readers who desire more specific information on the observations for any particular day are referred to Altrock and Gilliam (1985).

Notes on the Interpretation of NSO/SP Coronal Photometer Data

General Comments

- (1) Occasional discontinuities occur near geographic South as a result of clouds' causing suspension of the observations. This occurs mainly in observations prior to 1976.
- (2) Although the system appears to be stable, early data (up through 1982) suffer from too infrequent calibration. Calibration was done only every few months. Daily transparency changes were ignored, and this causes intensities to be too low for scans taken early or late in the day and during episodes of volcanic dust pollution, etc. In addition, telescope polarization varies diurnally from 0 to 12%, and these variations may affect the data. Since 1982, calibrations have been done each day with the scans whenever possible and the data are corrected for changes in light level between the time of the scan and the calibration.
- (3) Under conditions of excessive scattering, due to contamination on the objective lens, or possibly sky conditions, the data suffer from an East-West asymmetry. The East limb data are biased to be too bright. Conversely, the West limb observations are too dark. The

practical effect on the West limb observations is more marked, because the A/D converter does not accept values less than a zero level set by the observer. Thus, West limb data in these situations appear to be clipped. Intervals when this was a particular problem are noted below.

Specific Comments

1973-1974: Calibration may be wrong by 50%; West limb scans are of doubtful value.

1975-1978: Radius errors through August 1978--a constant (incorrect) solar radius was used, resulting in an annual variation of the height at which data were recorded. For example, scans labelled 1.15 were actually recorded from 1.16 in January to 1.20 in July. This can cause intensities in July to be as much as 30° too low, based on an $r^{-7.5}$ variation.

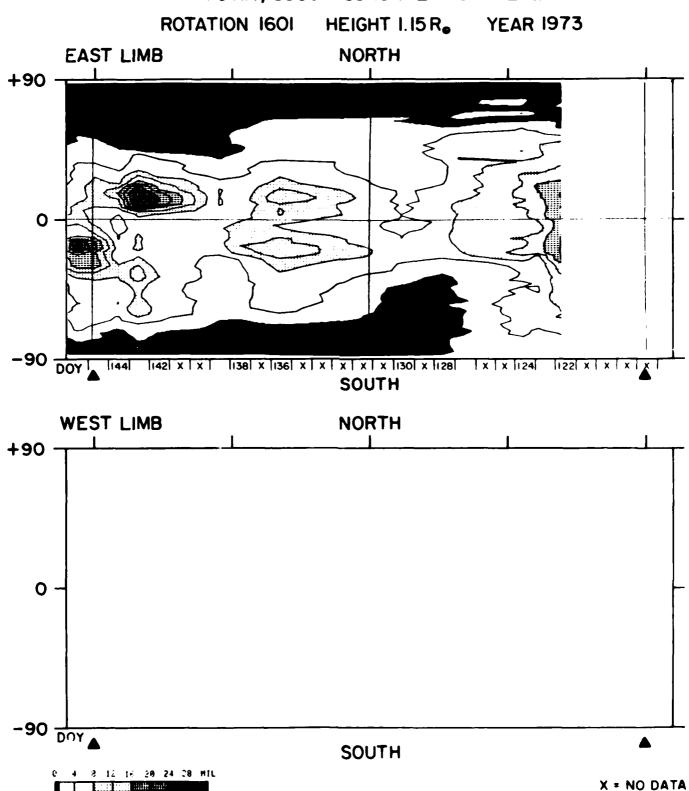
1977: Severe excessive scatter DOY 220 through 258.

1978: Severe excessive scatter DOY 178 through 180.

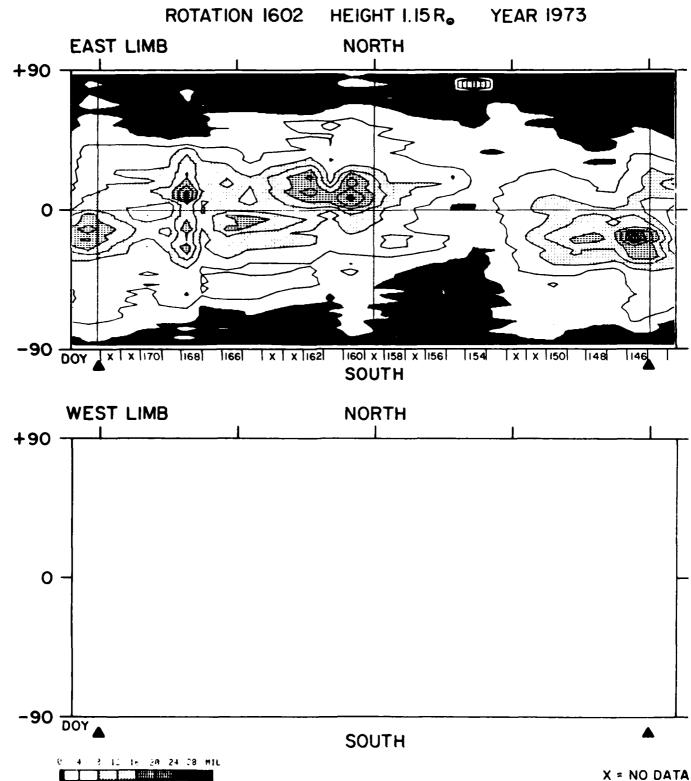
1980-1983: Error in ephemeris program-radius errors of order $< 0.01~R_{\odot}$ will affect intensities by a few percent due to incorrect year being entered in ephemeris call.

IV. SYNOPTIC MAPS OF THE GREEN LINE CORONA

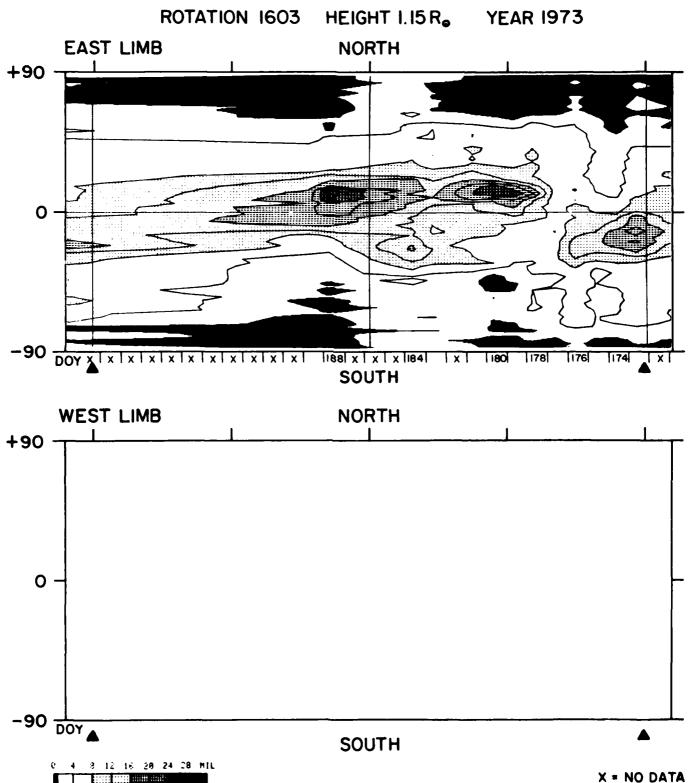
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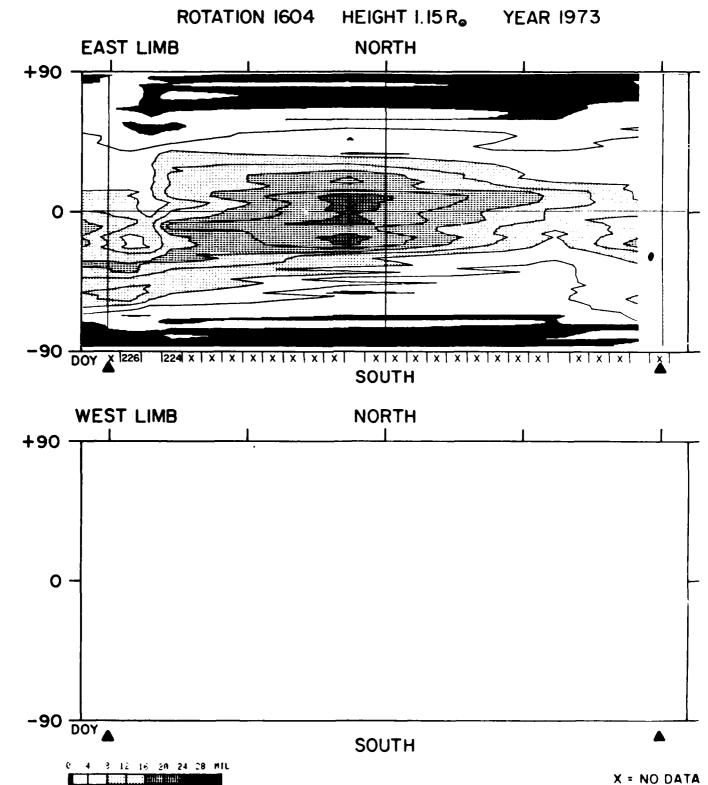
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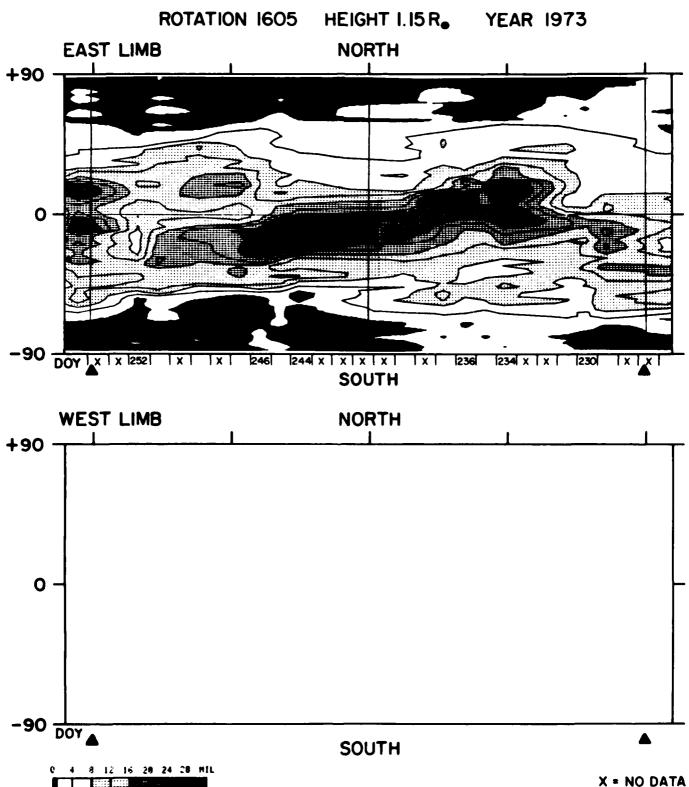
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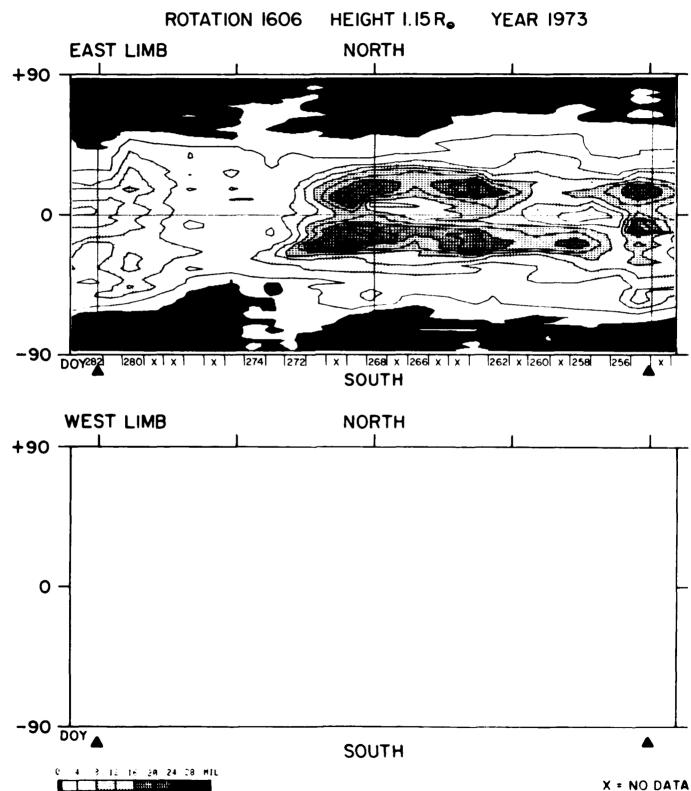
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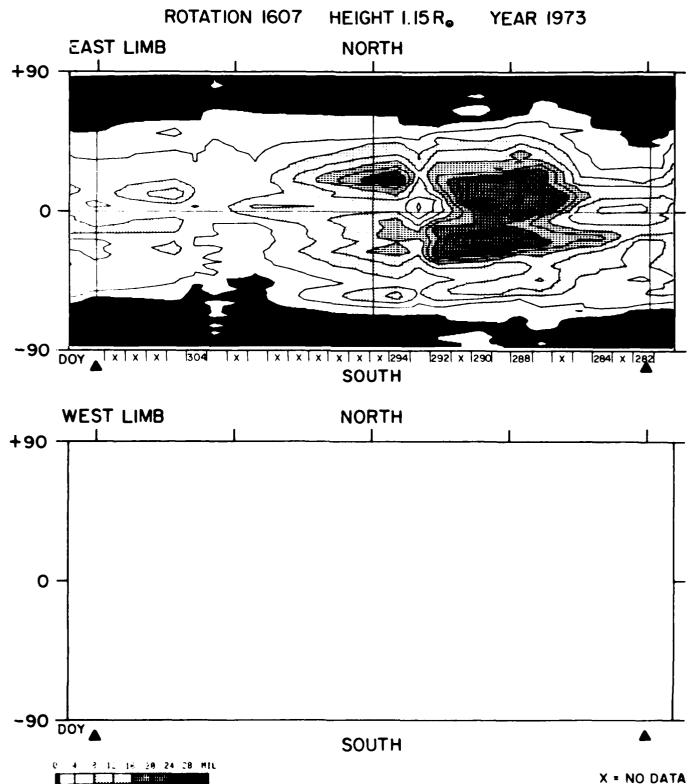
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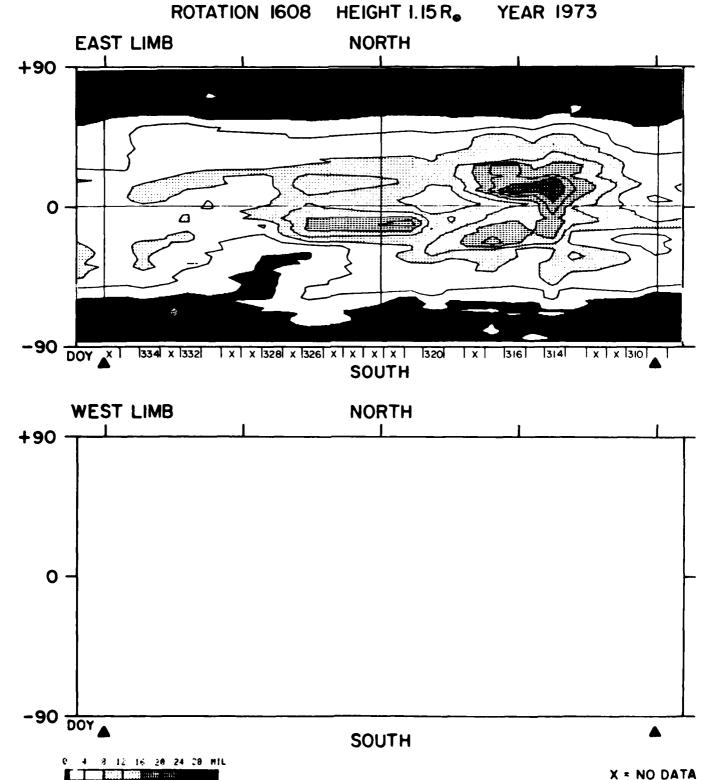
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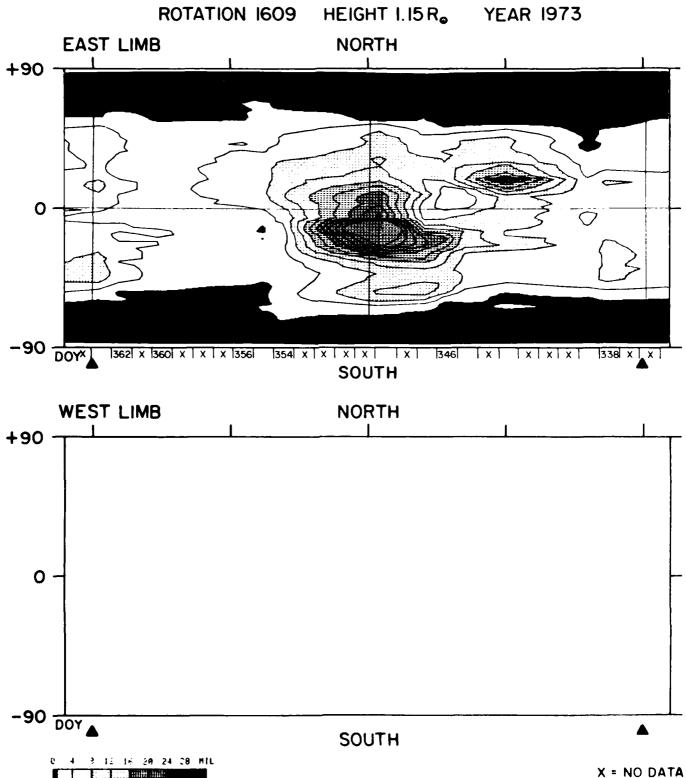
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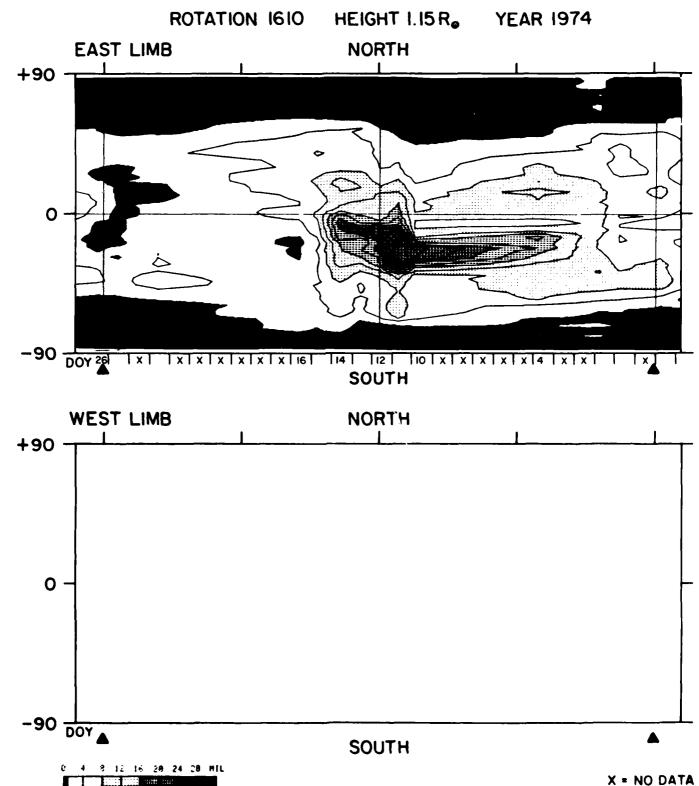
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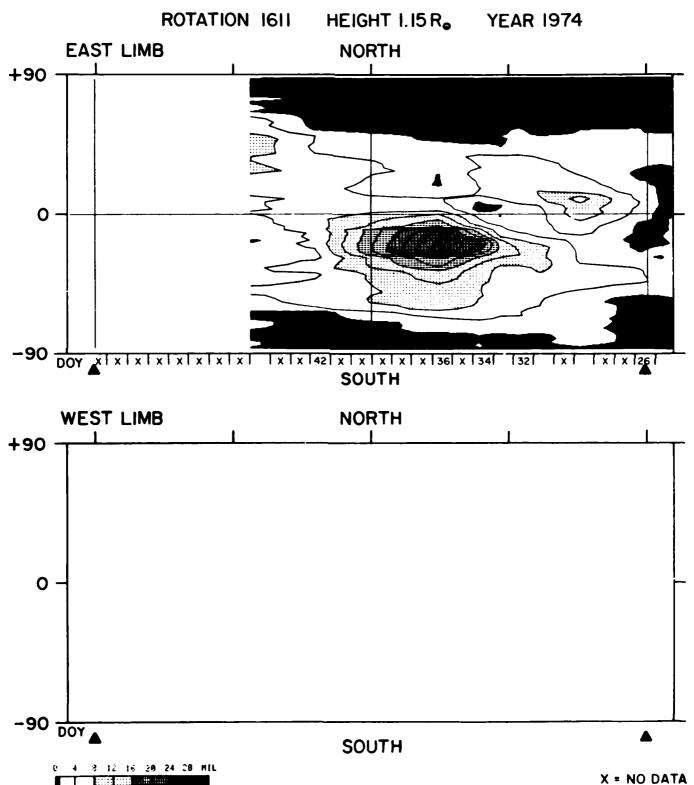
Fe XIV, 5303 & CORONAL PHOTOMETER



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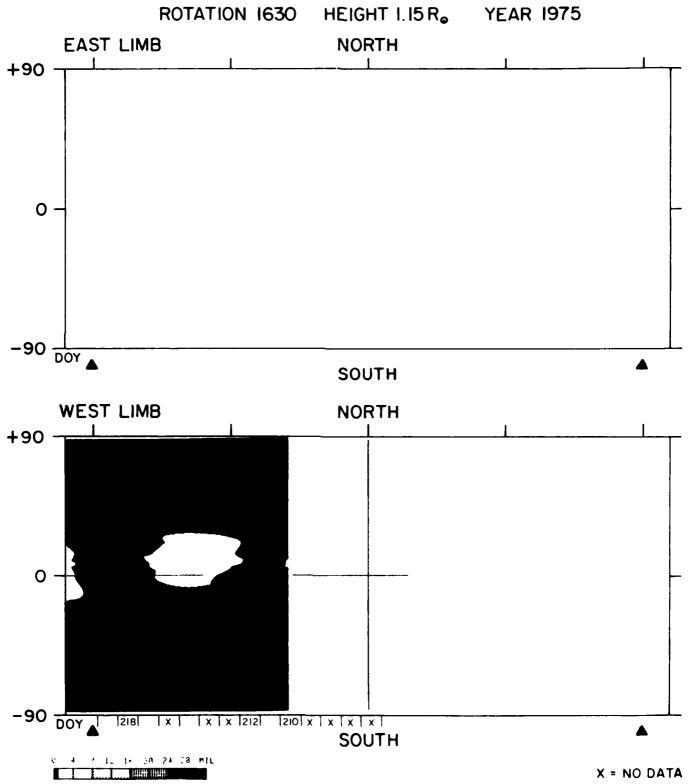


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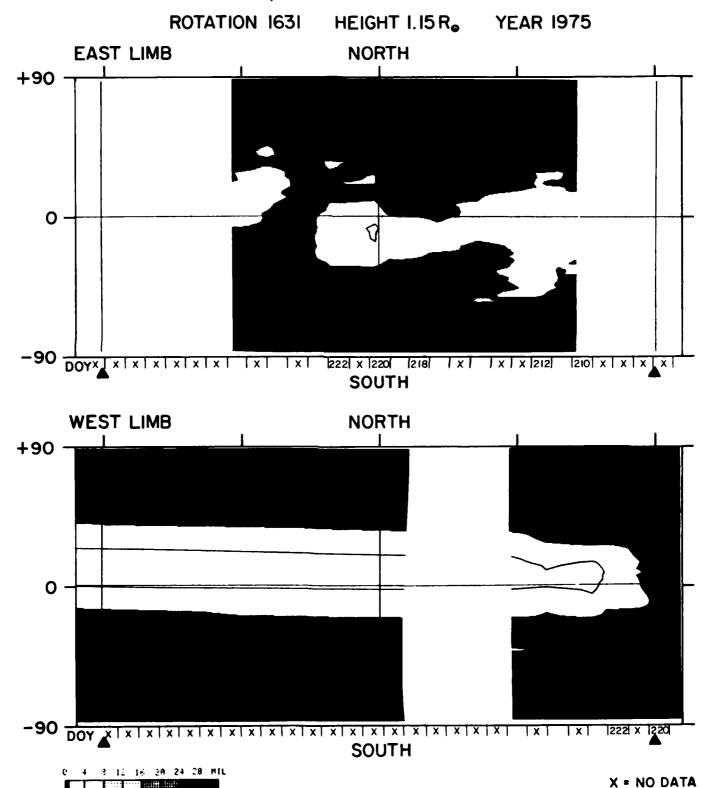


No data for Carrington rotations 1612-1629

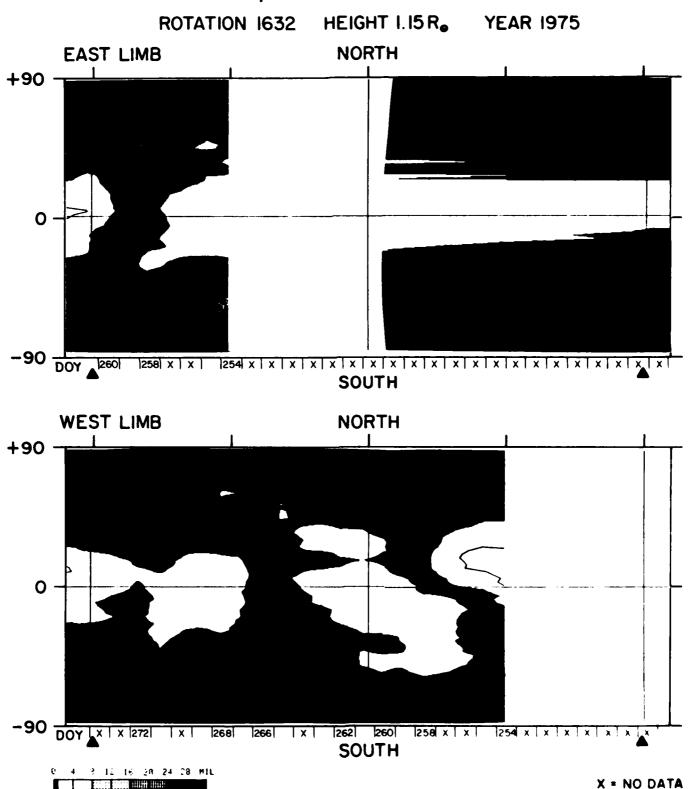
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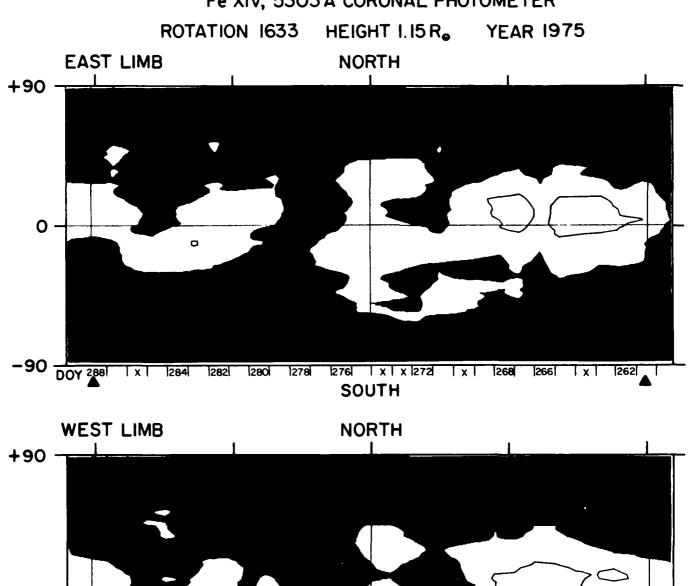
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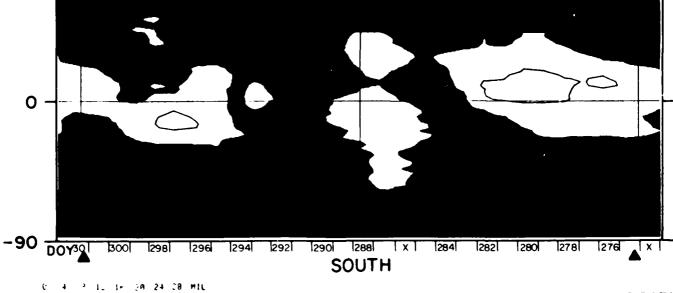


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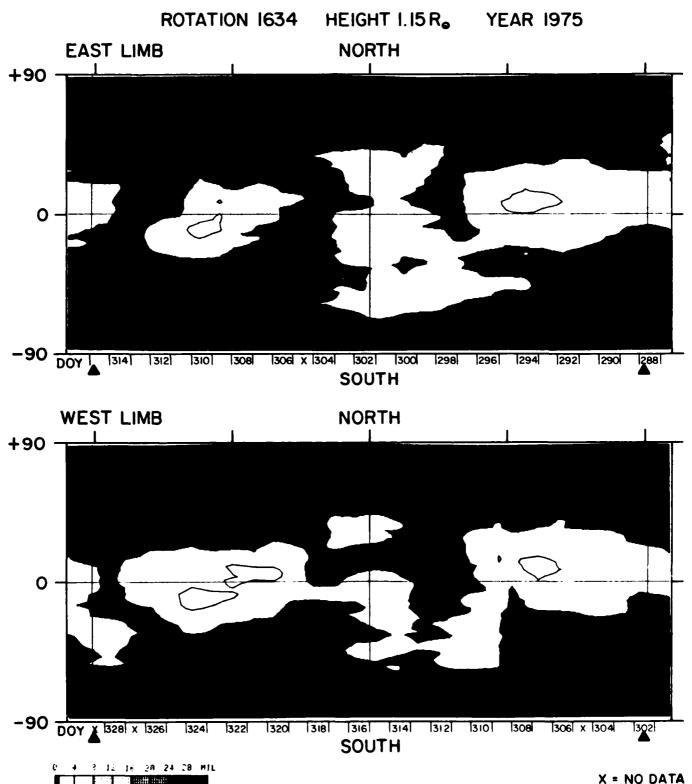
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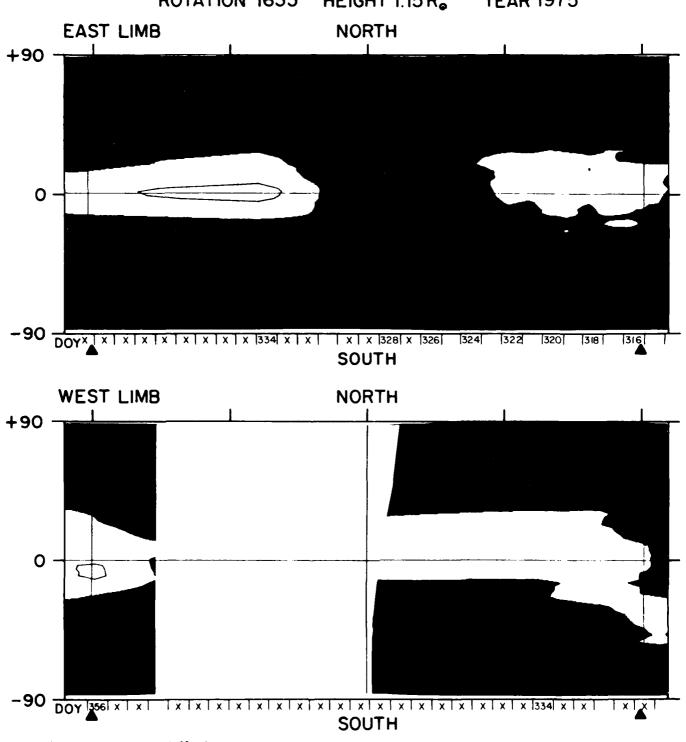


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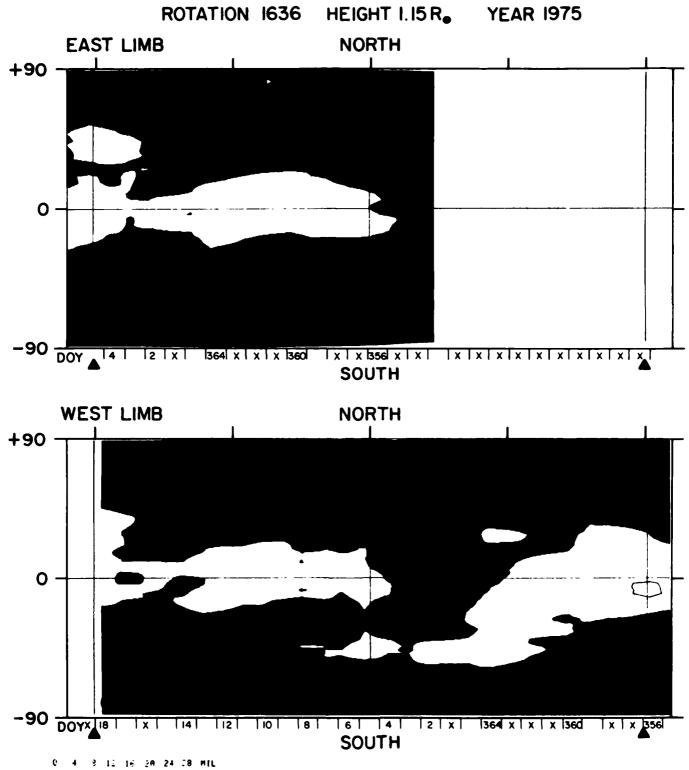


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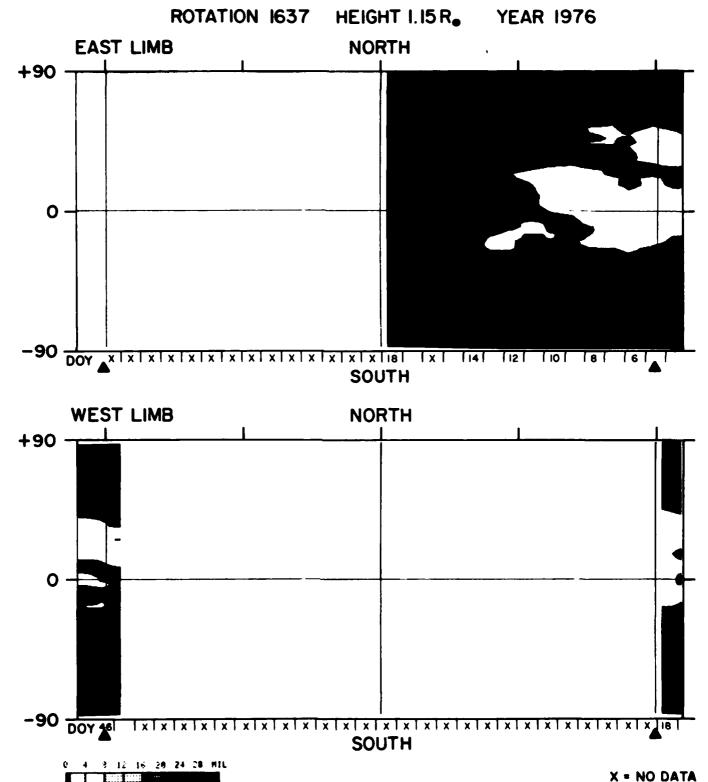
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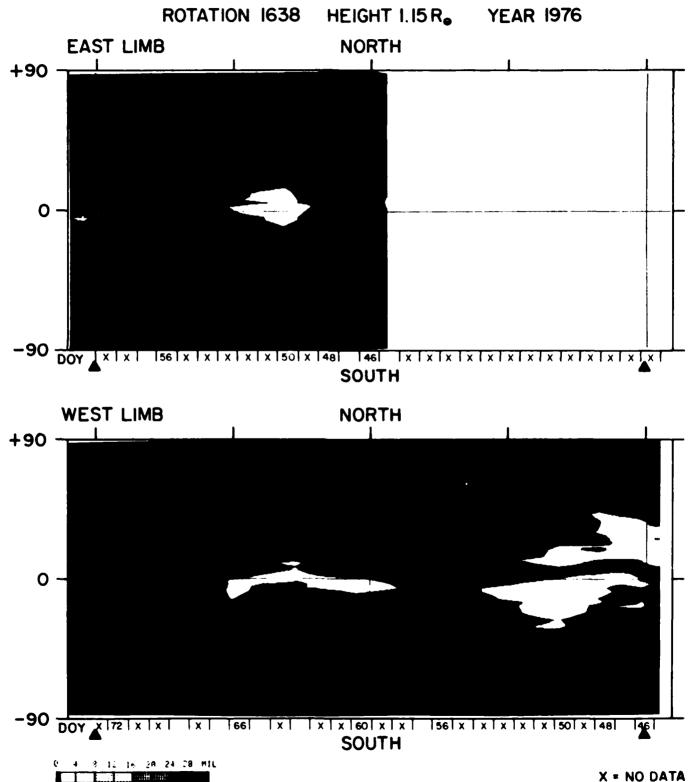


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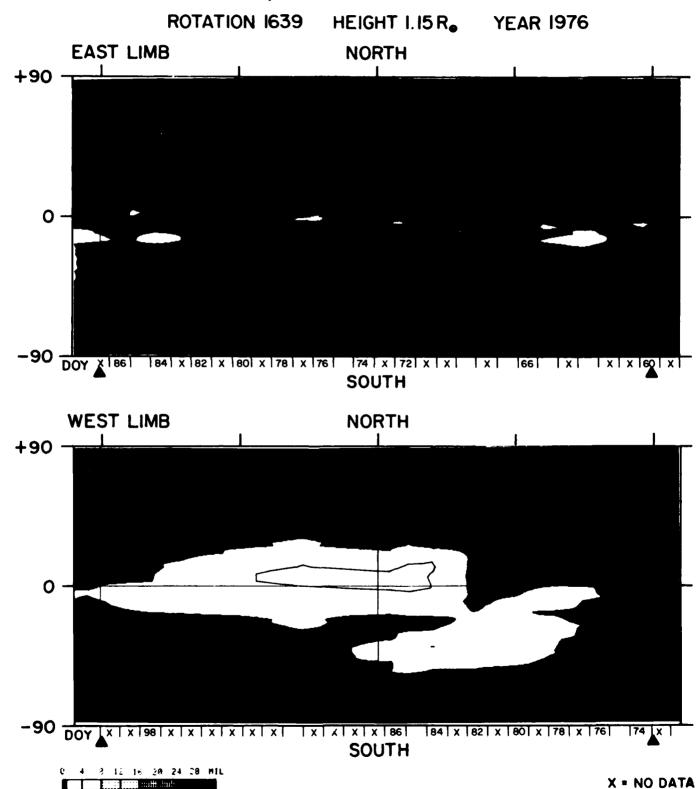
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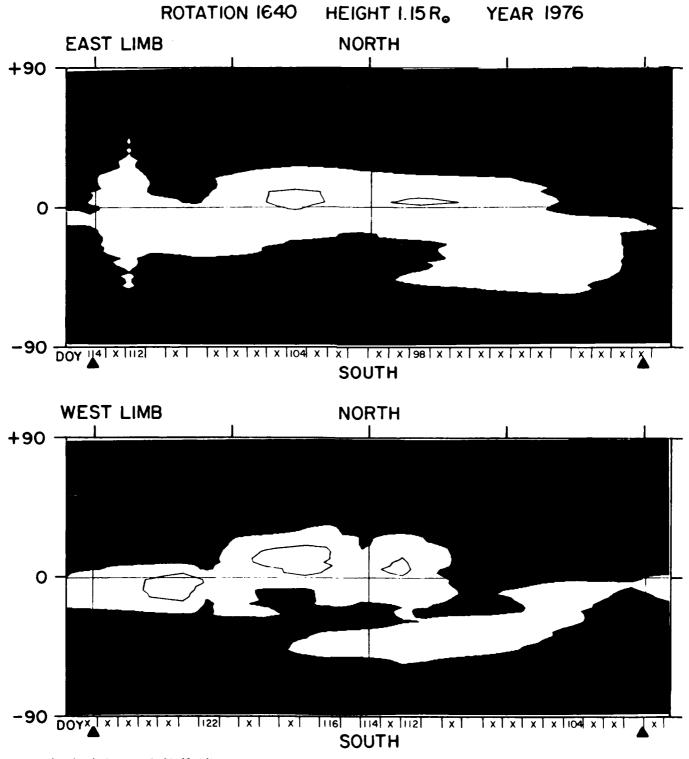
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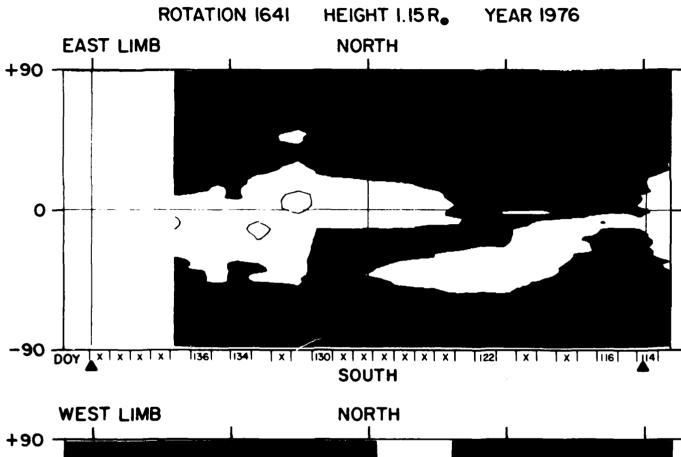


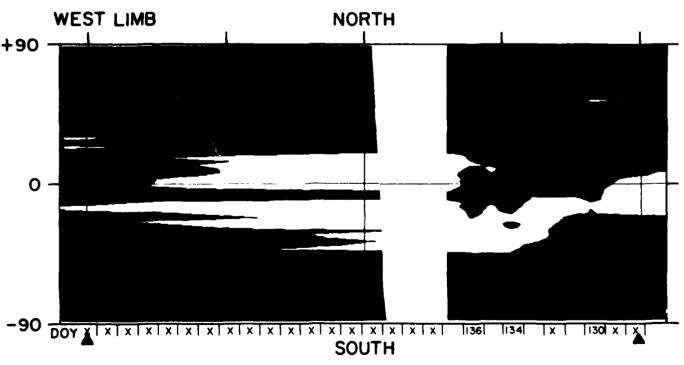
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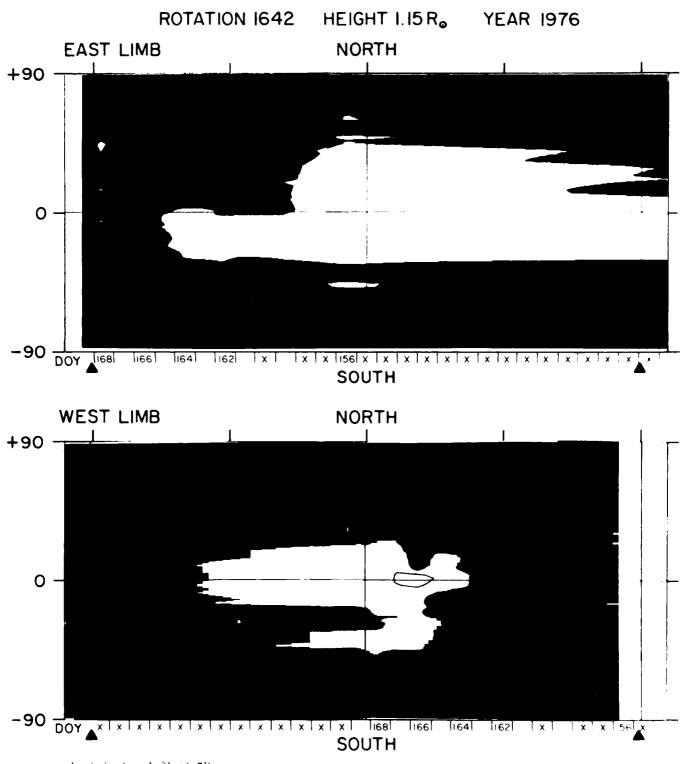
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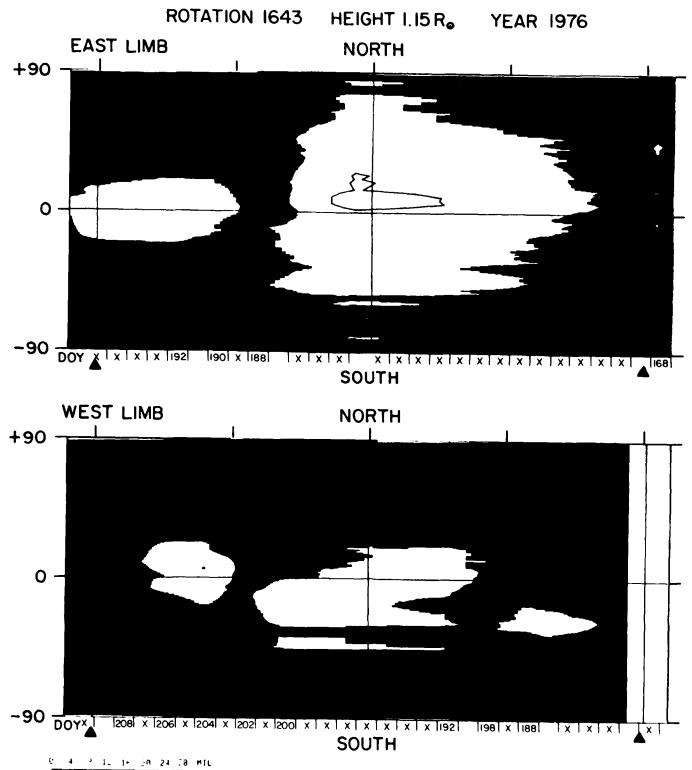


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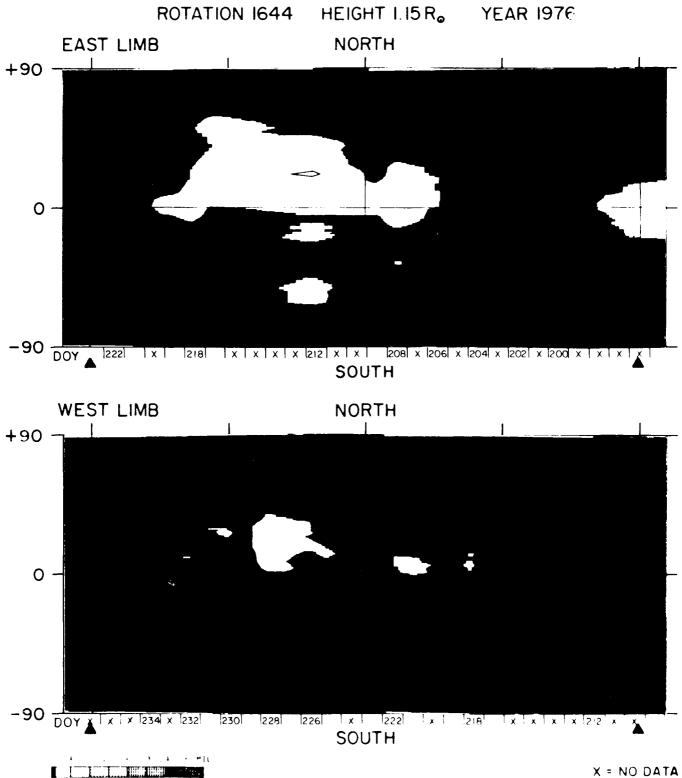
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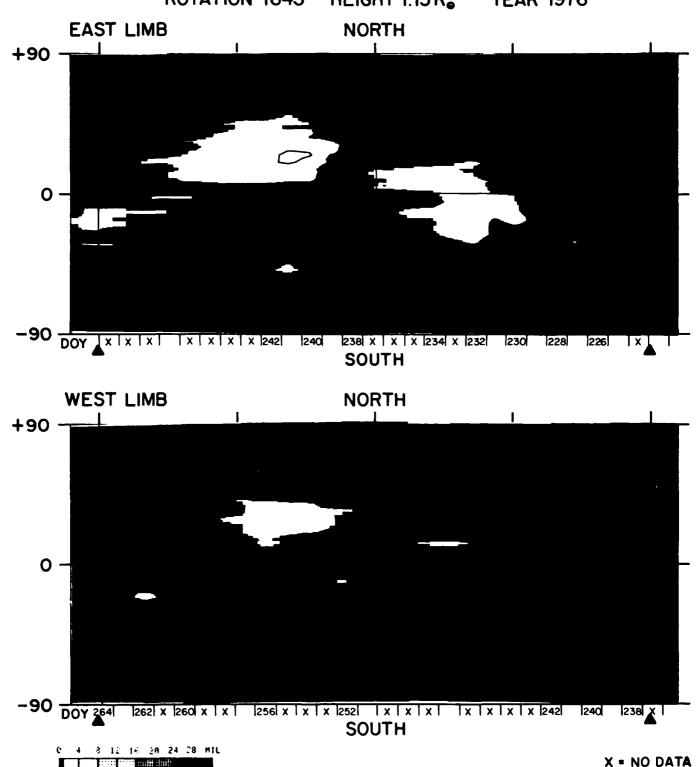
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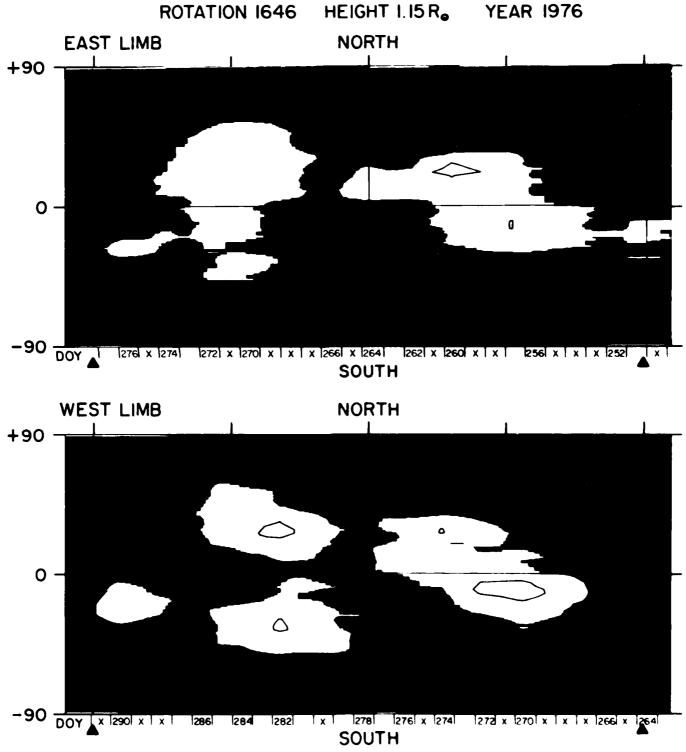
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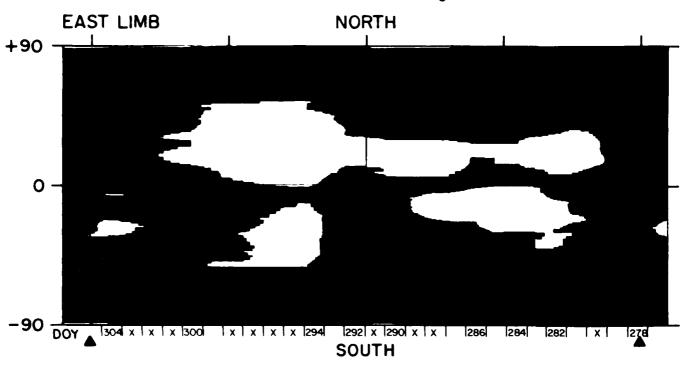
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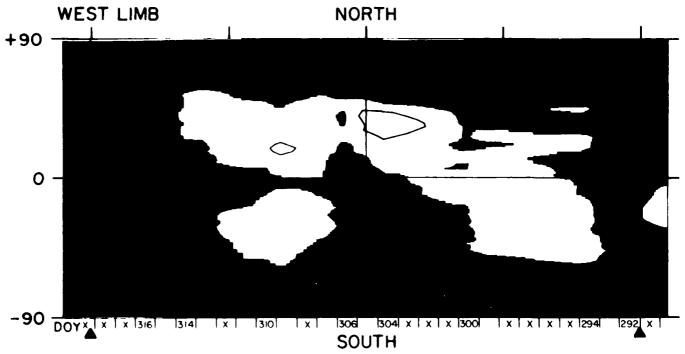


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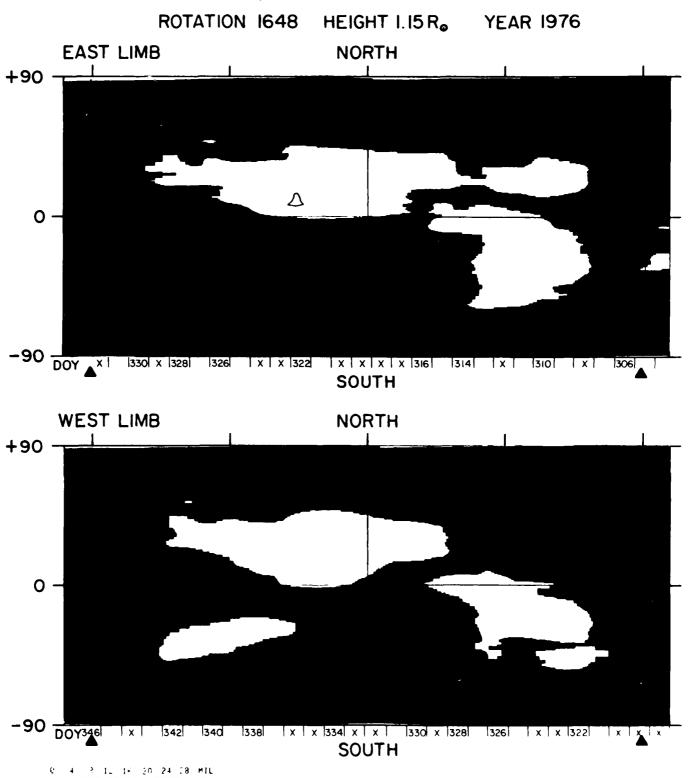


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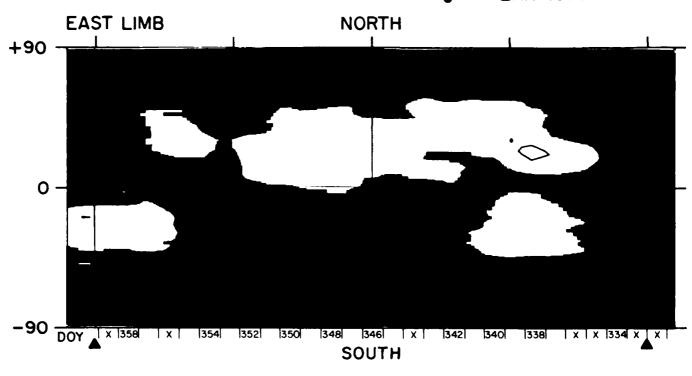


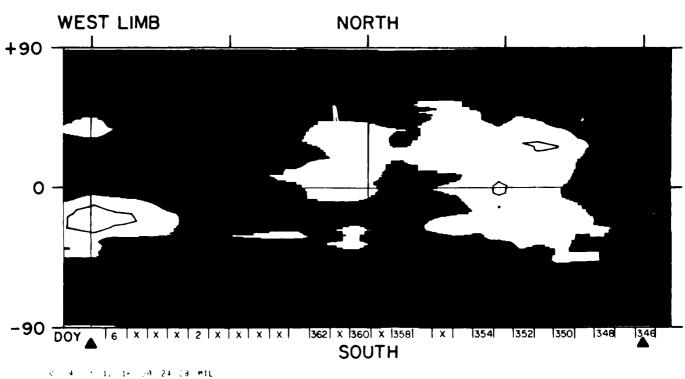


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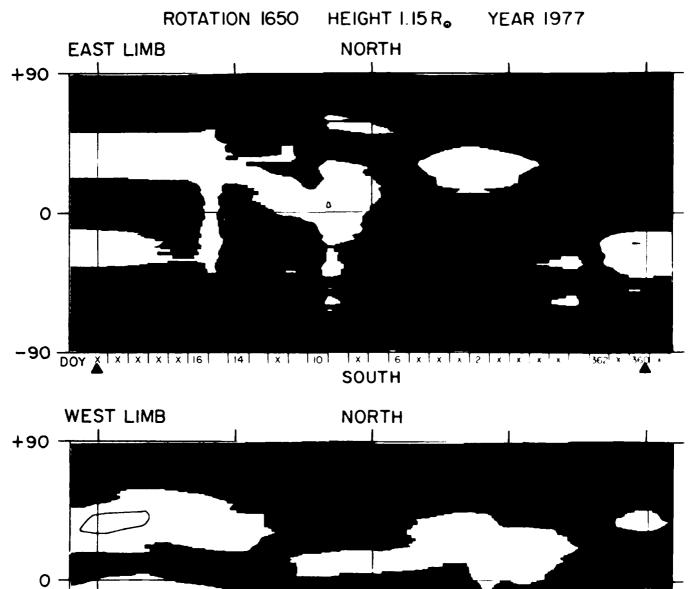


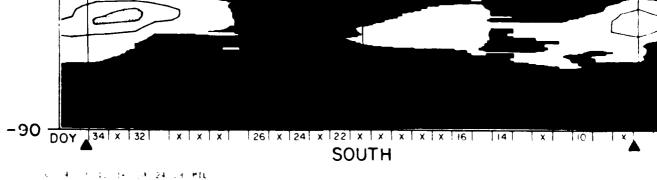
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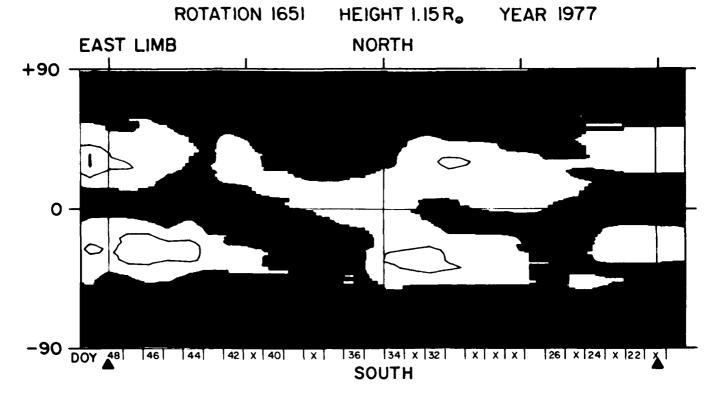
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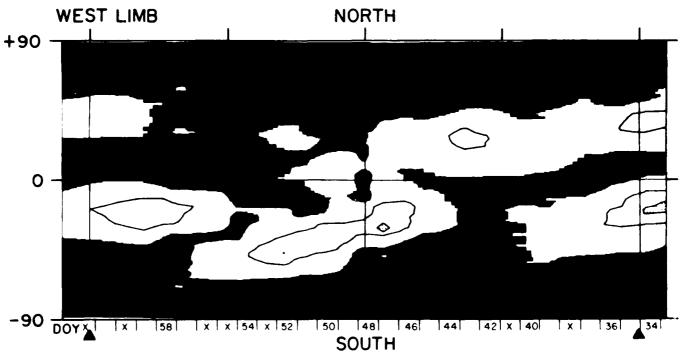




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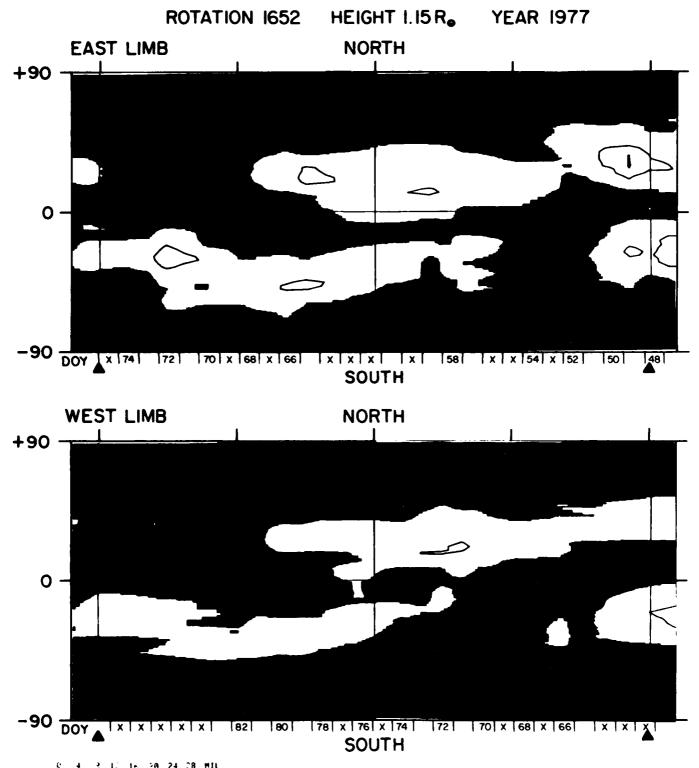
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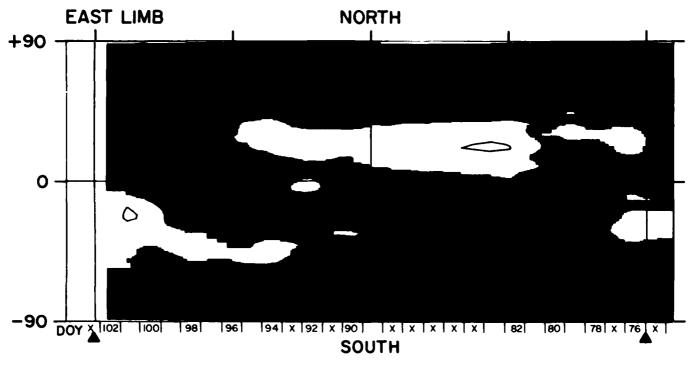


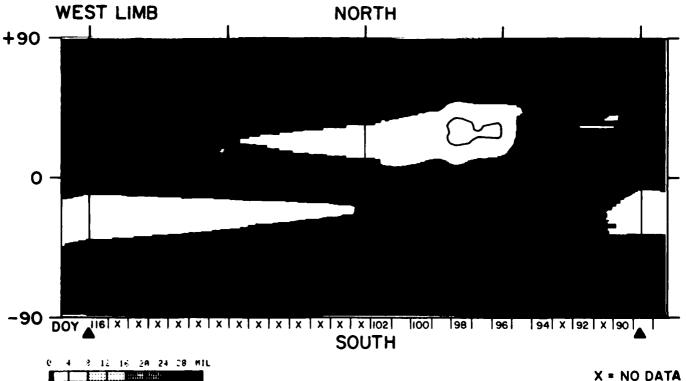
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Fe XIV, 5303 & CORONAL PHOTOMETER

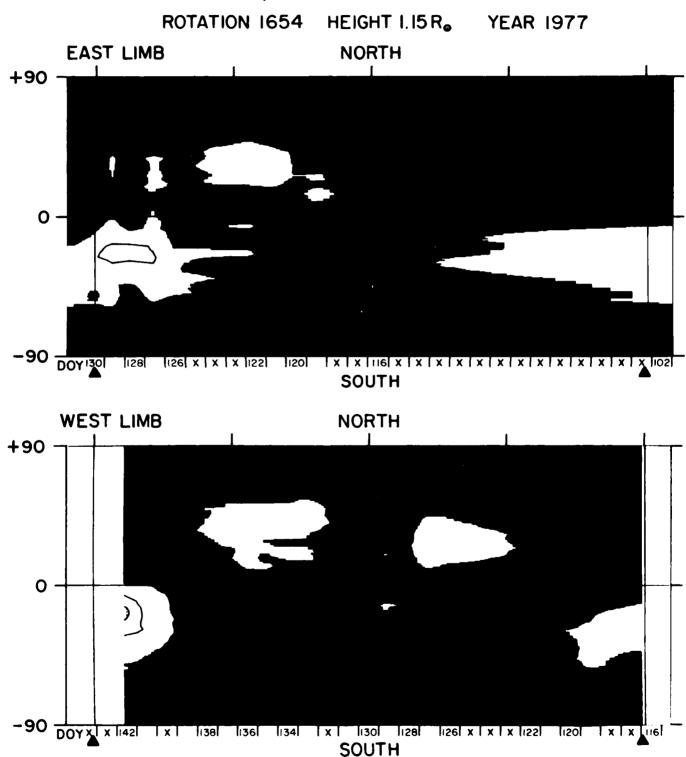


Fe XIV, 5303 & CORONAL PHOTOMETER **ROTATION 1653** HEIGHT 1.15 R. **YEAR 1977**

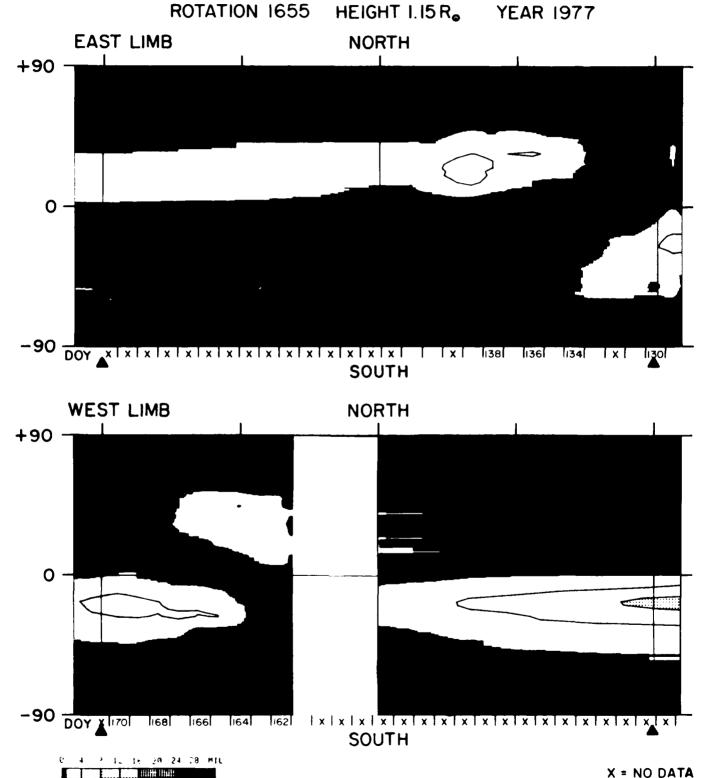




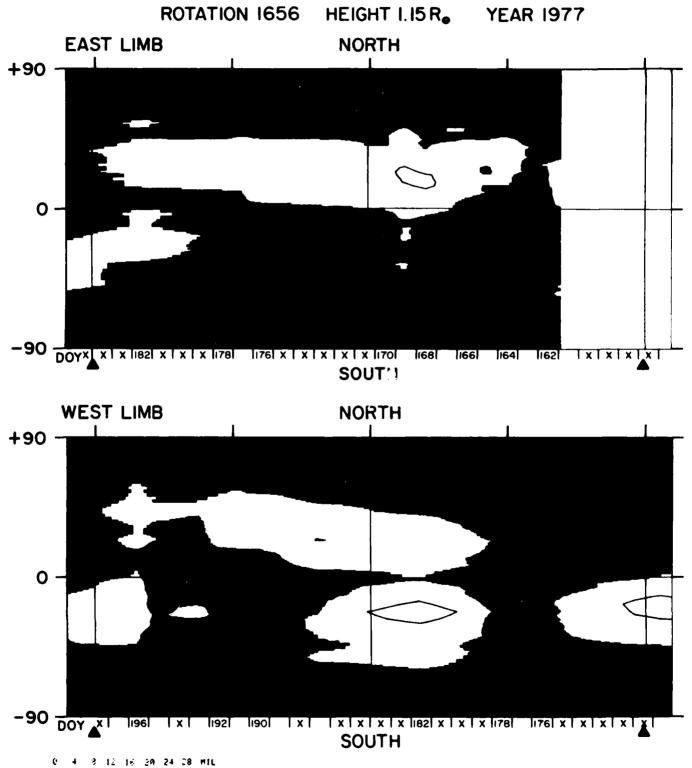
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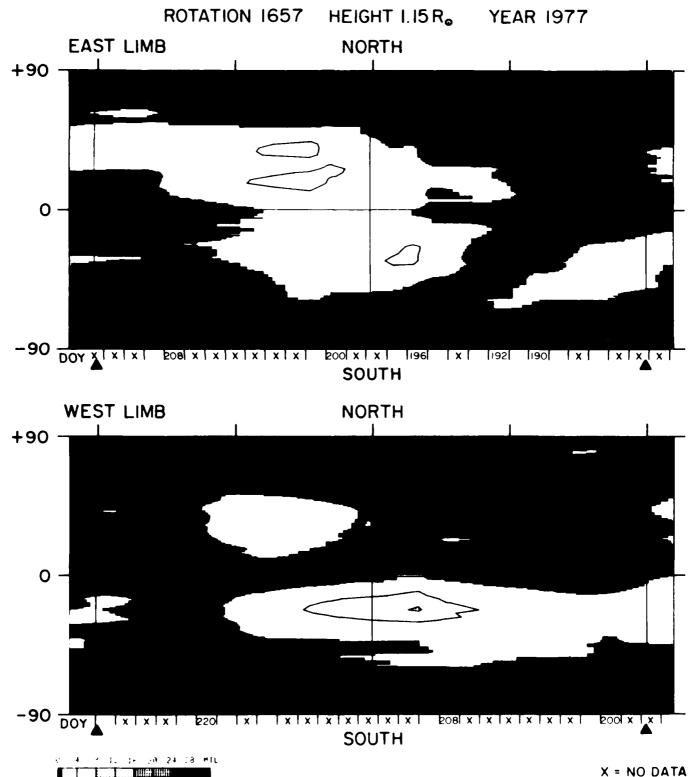
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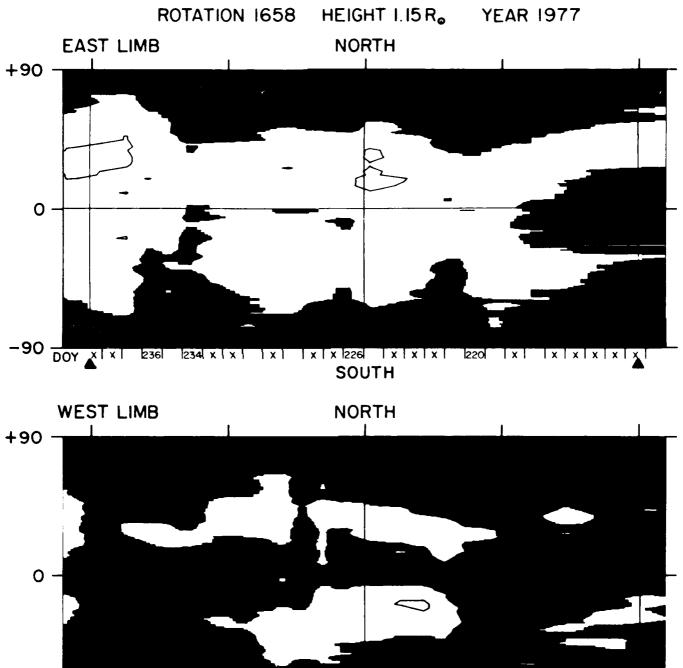
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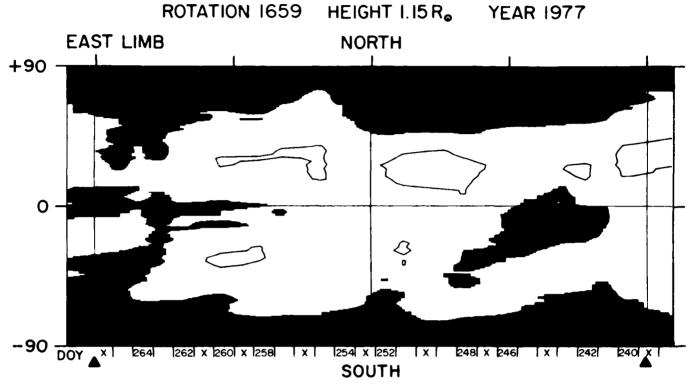
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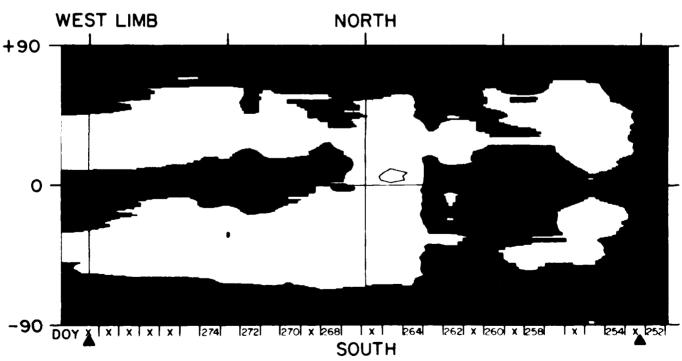


SOUTH

U 4 3 12 1+ 2A 24 18 MIL

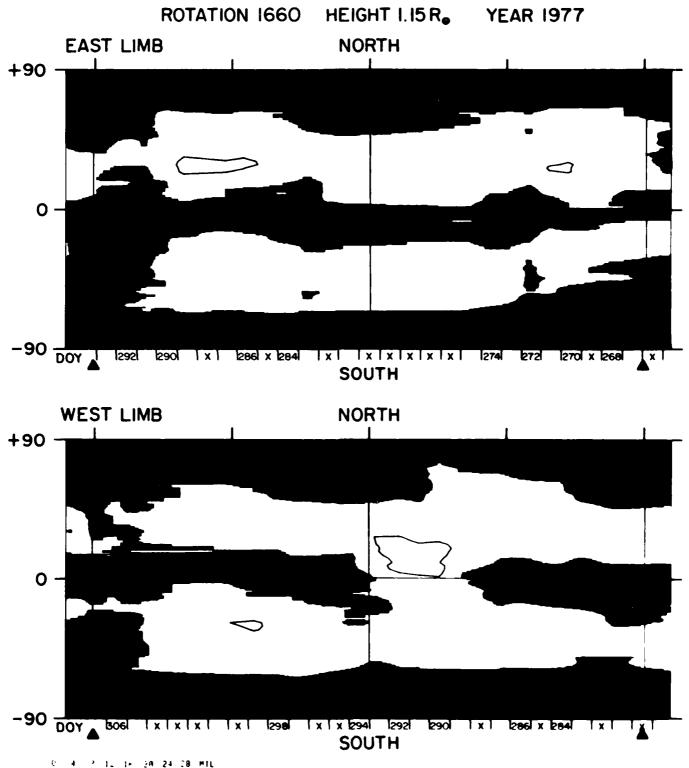
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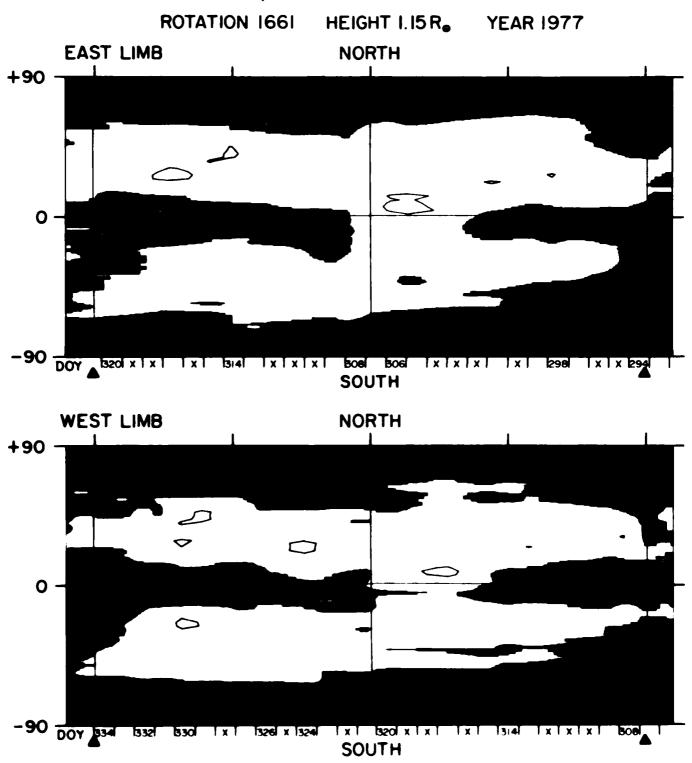


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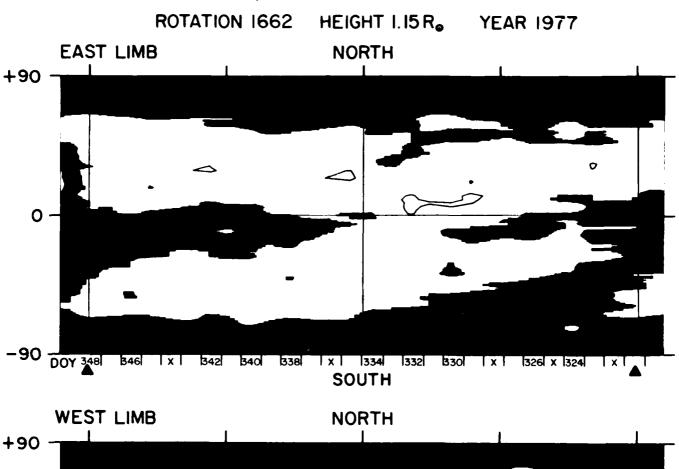
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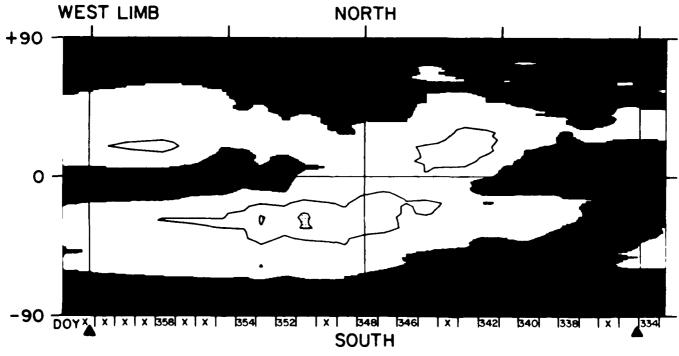


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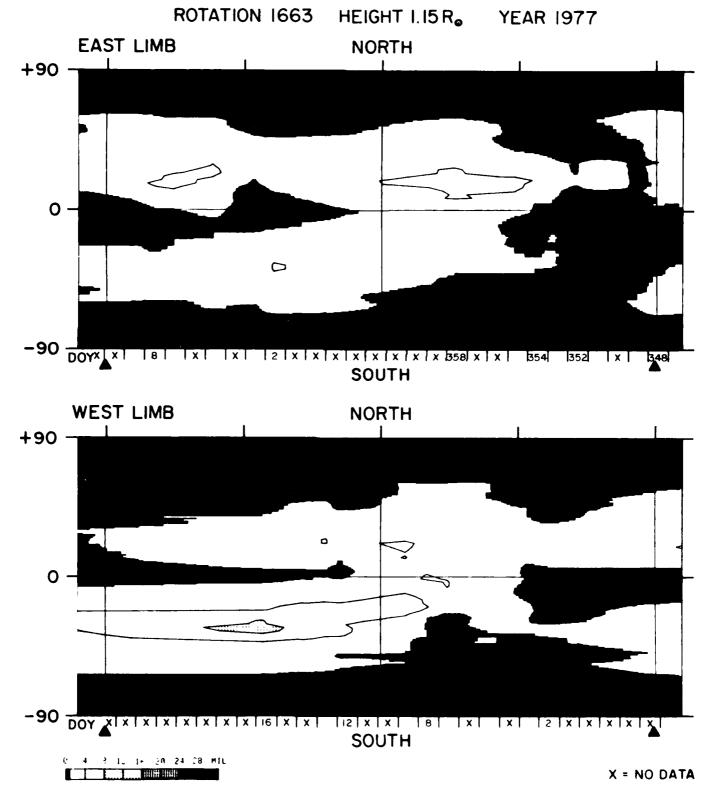
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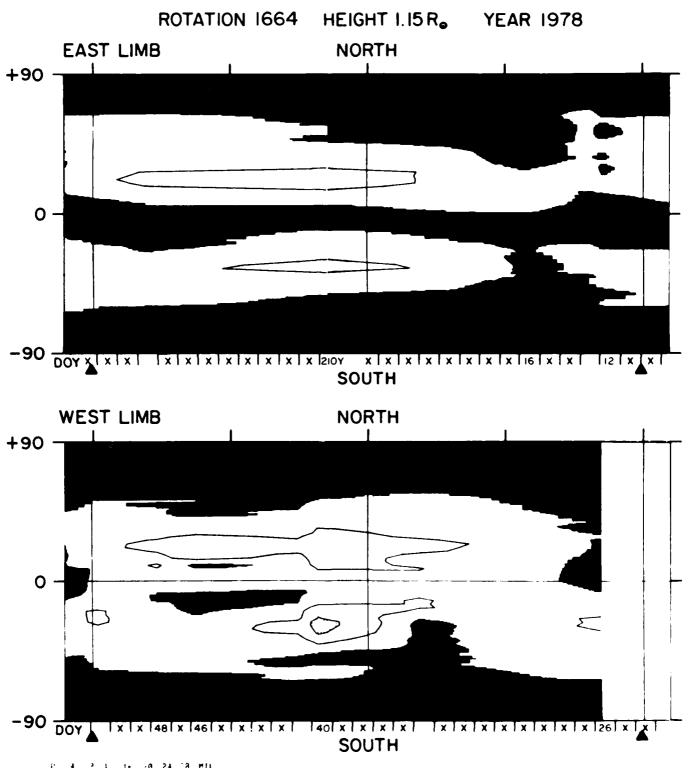


√ 4 2 12 1× 20 24 28 MIL

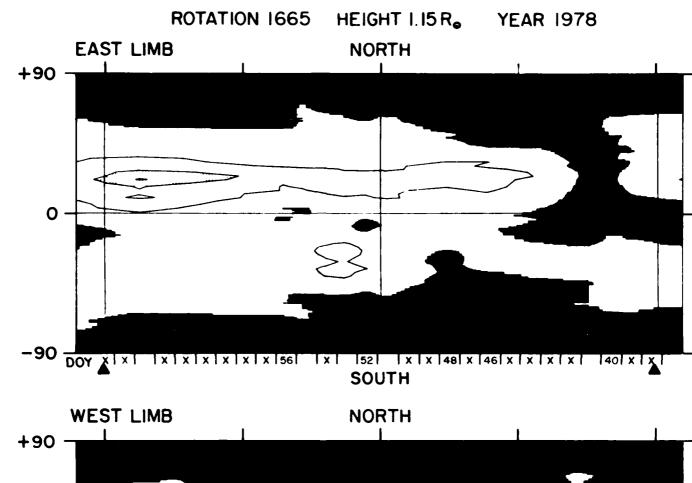
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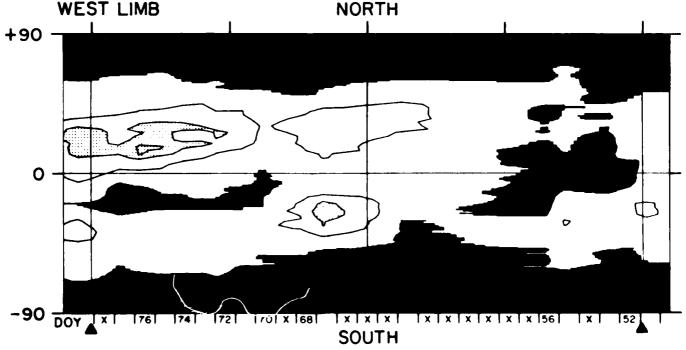


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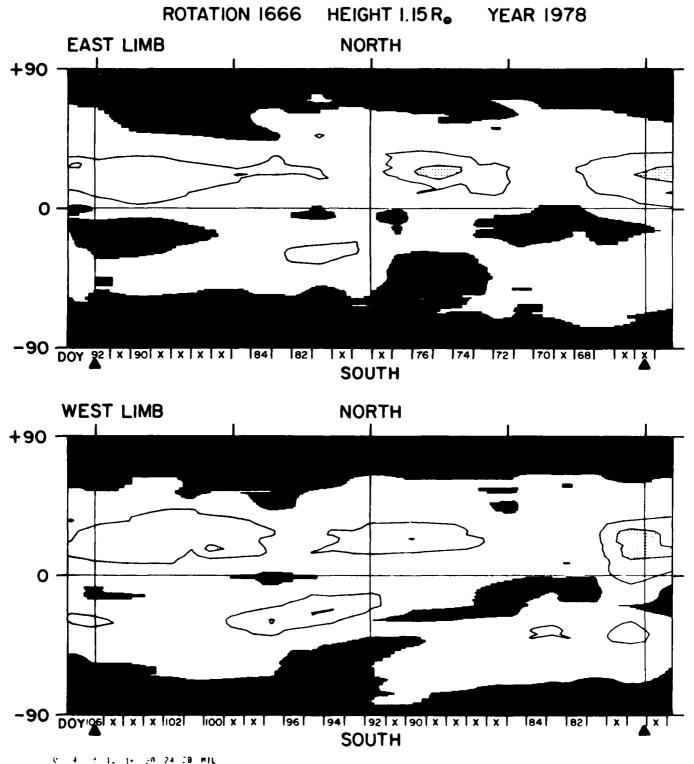


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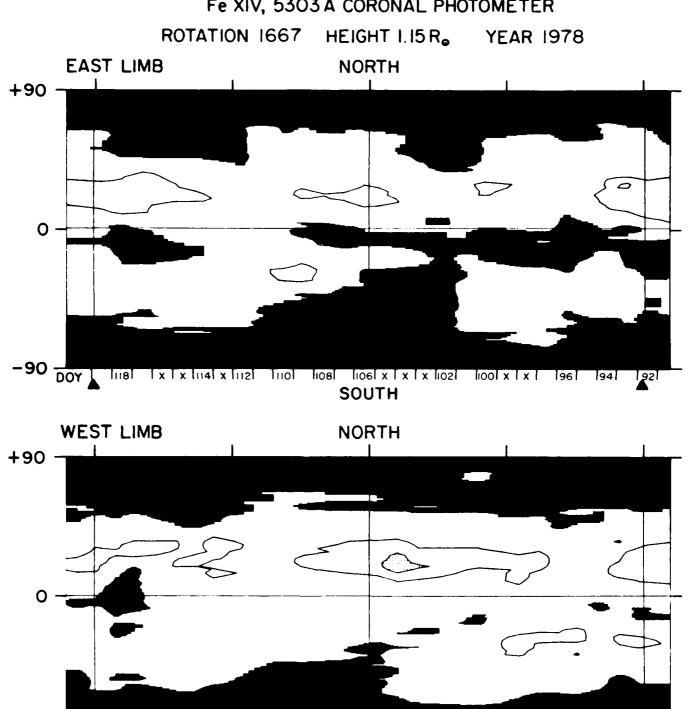




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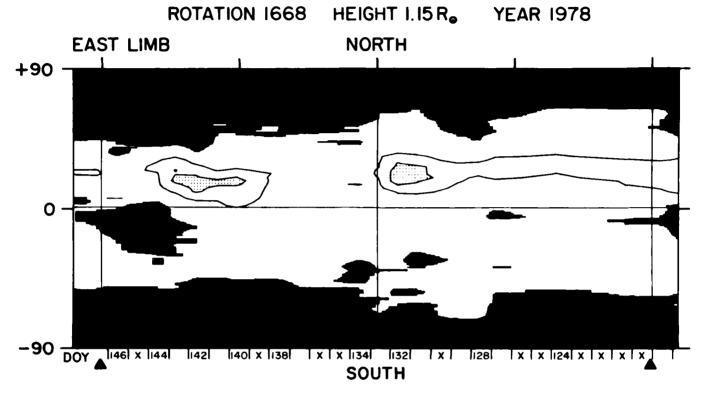


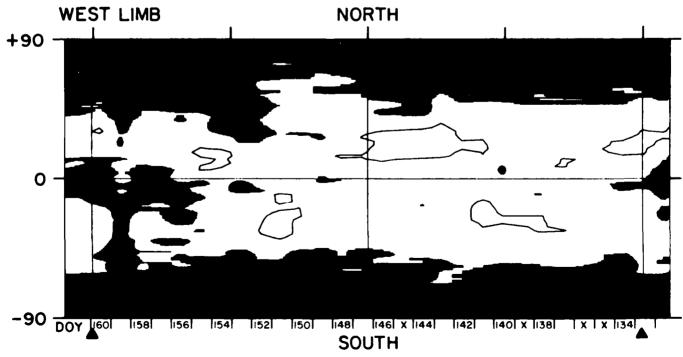
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SOUTH

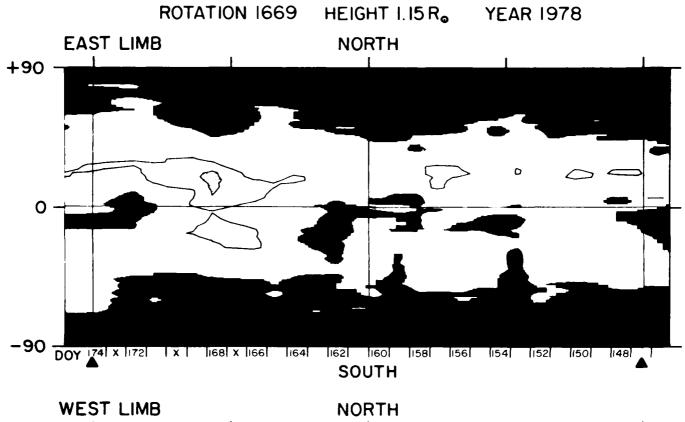
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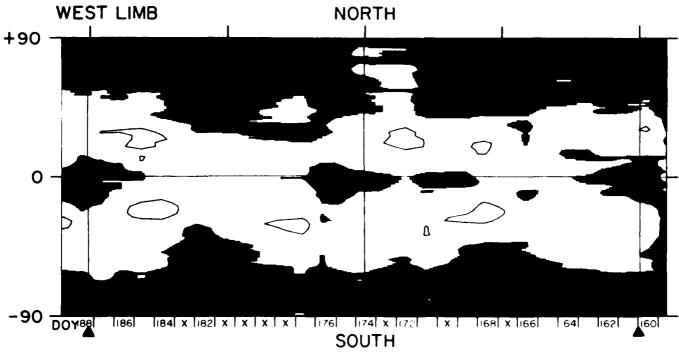




C 4 3 12 16 20 24 28 MIL

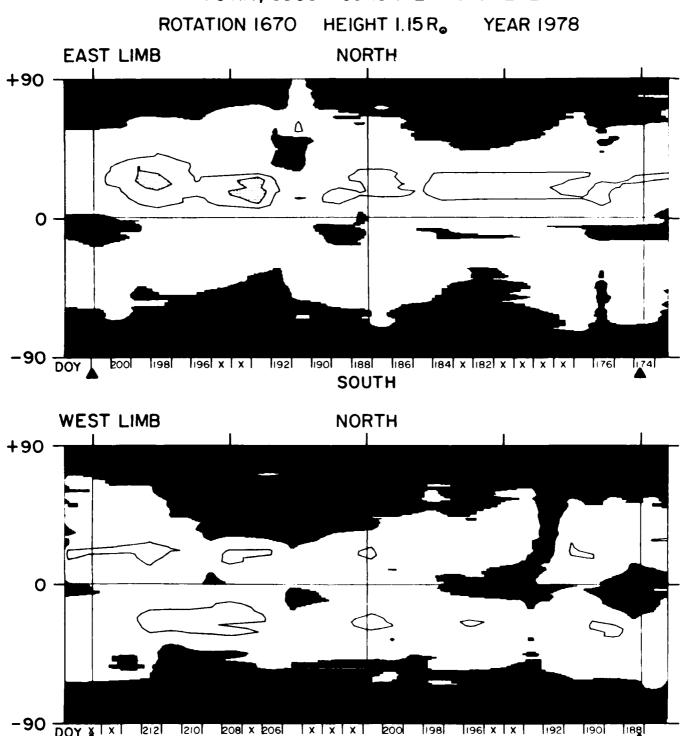
Fe XIV, 5303 Å CORONAL PHOTOMETER





€ 4 2 12 1× 20 24 28 MI

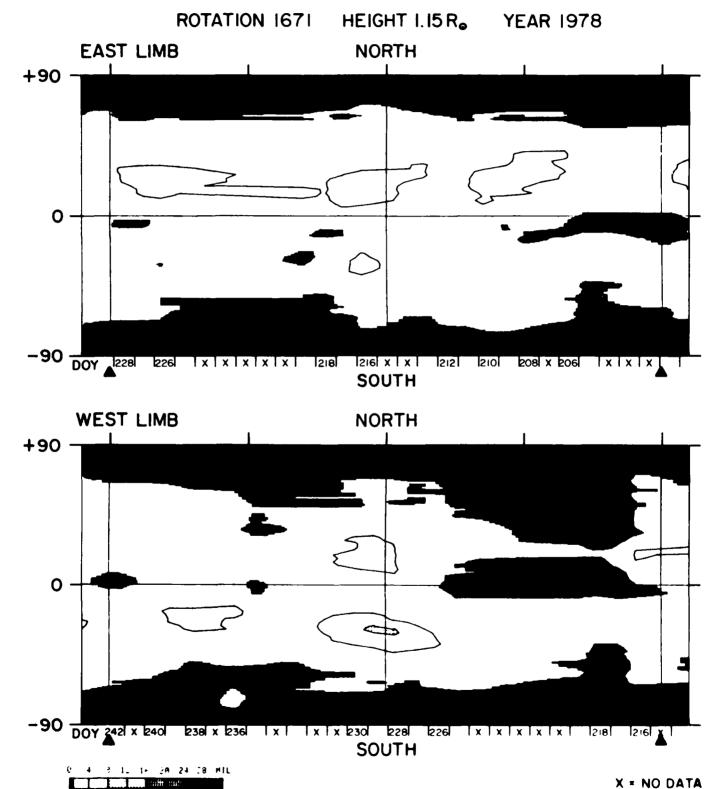
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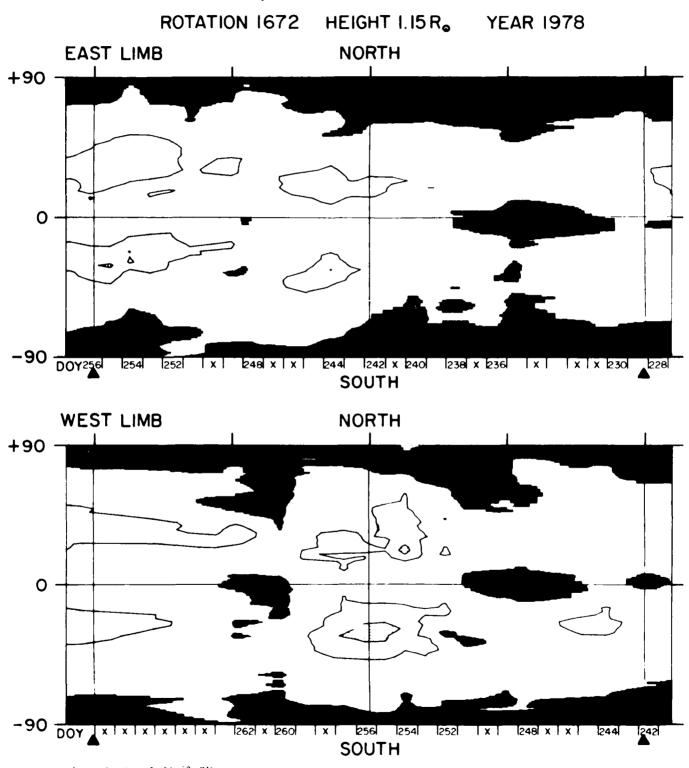
SOUTH

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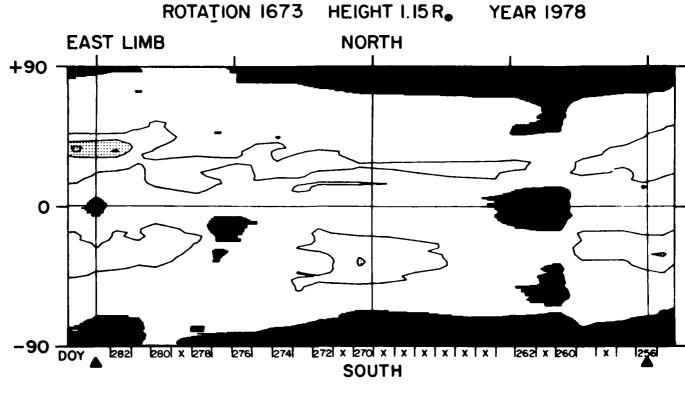
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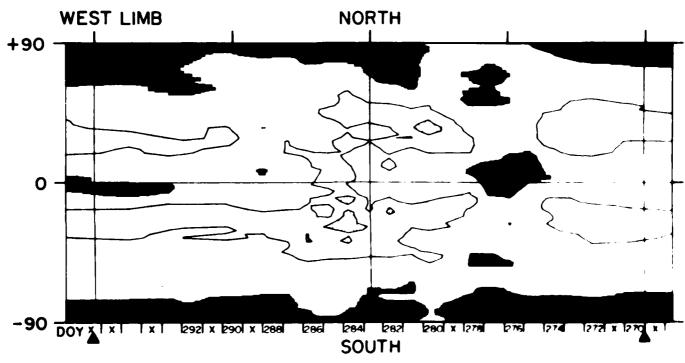


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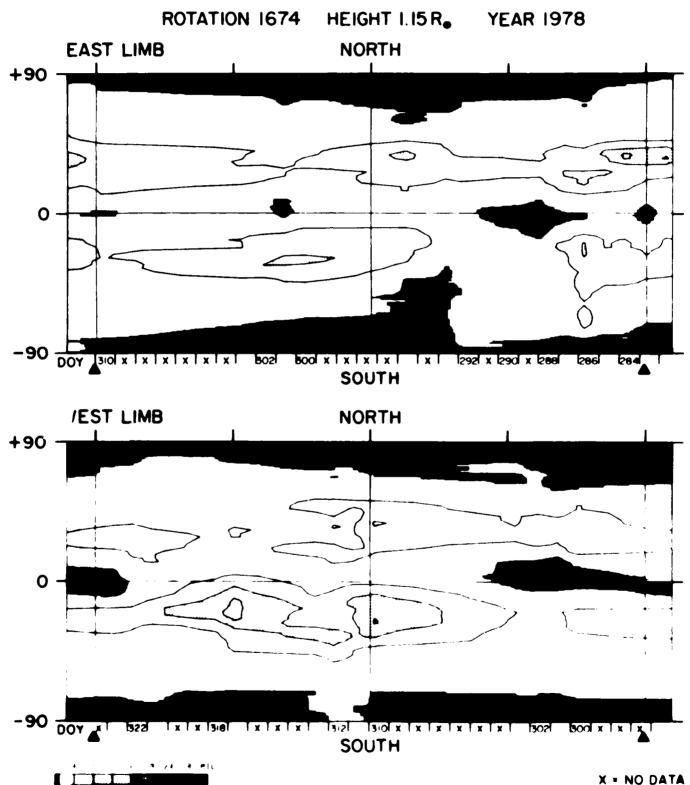


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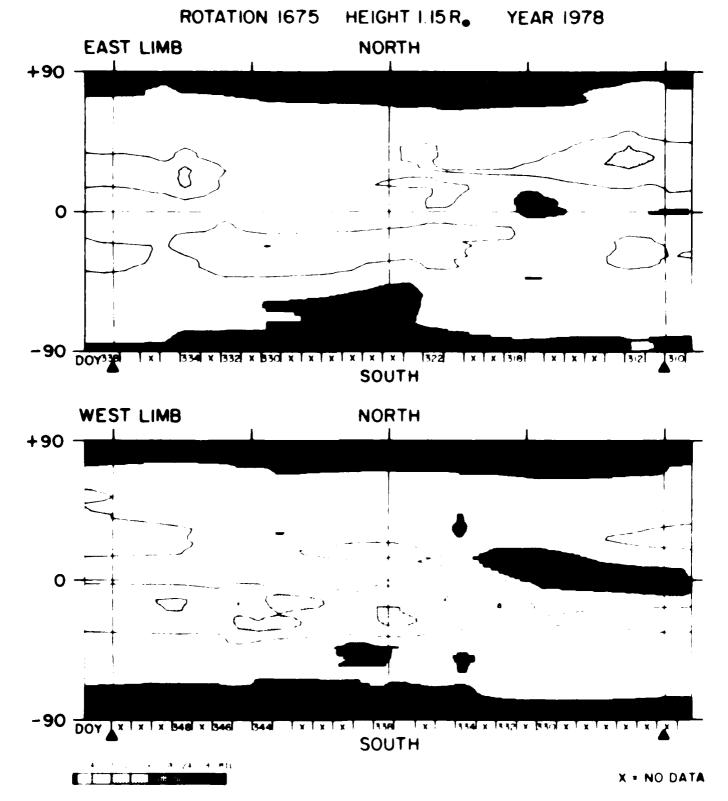




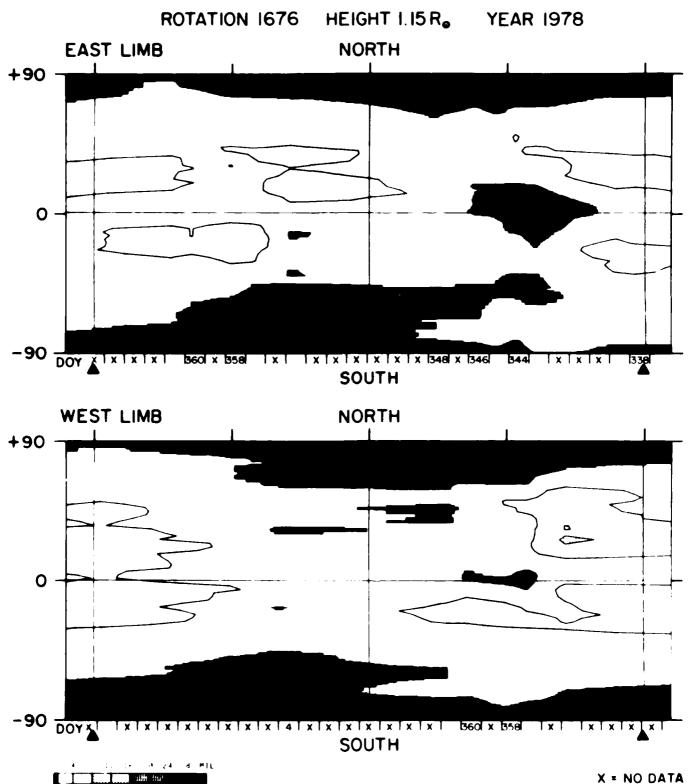
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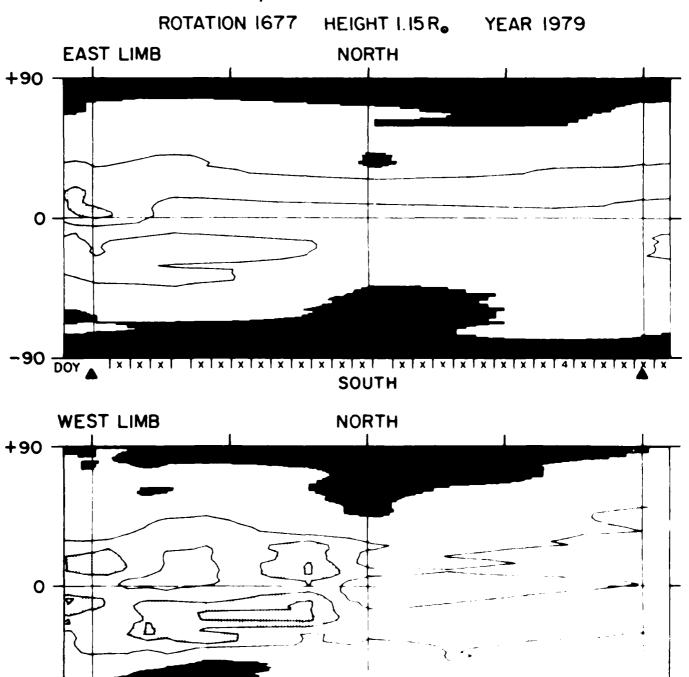
Fe XIV, 5303 A CORONAL PHOTOMETER



Fe XIV, 5303 & CORONAL PHOTOMETER

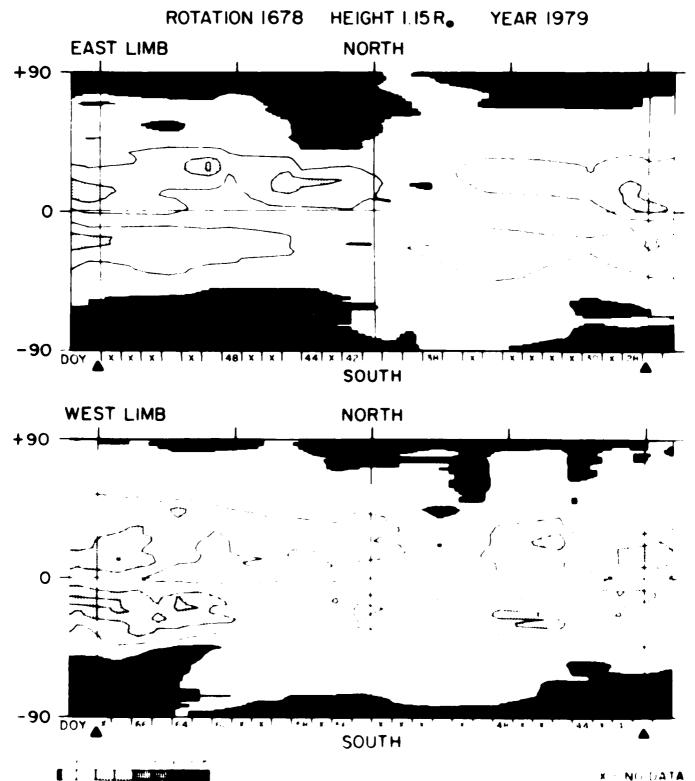


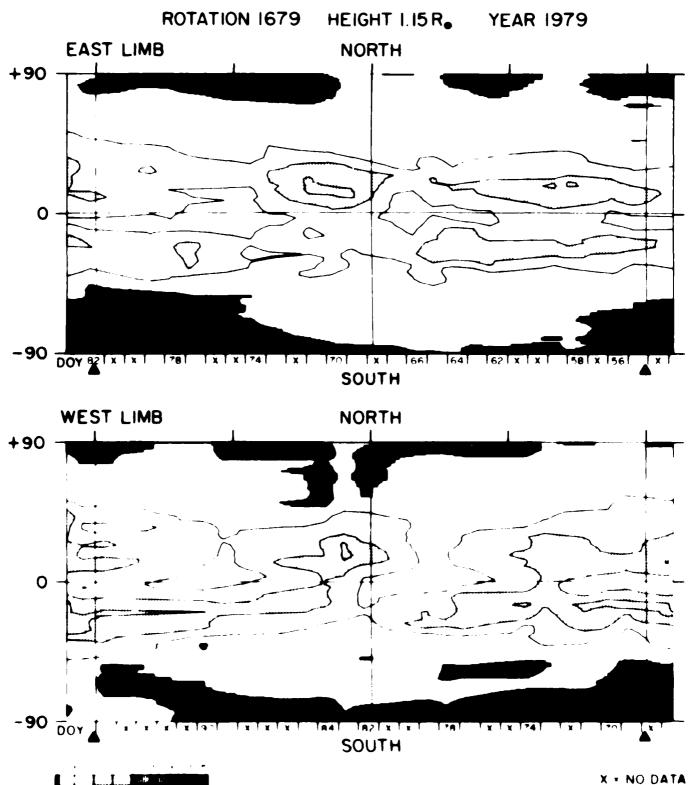
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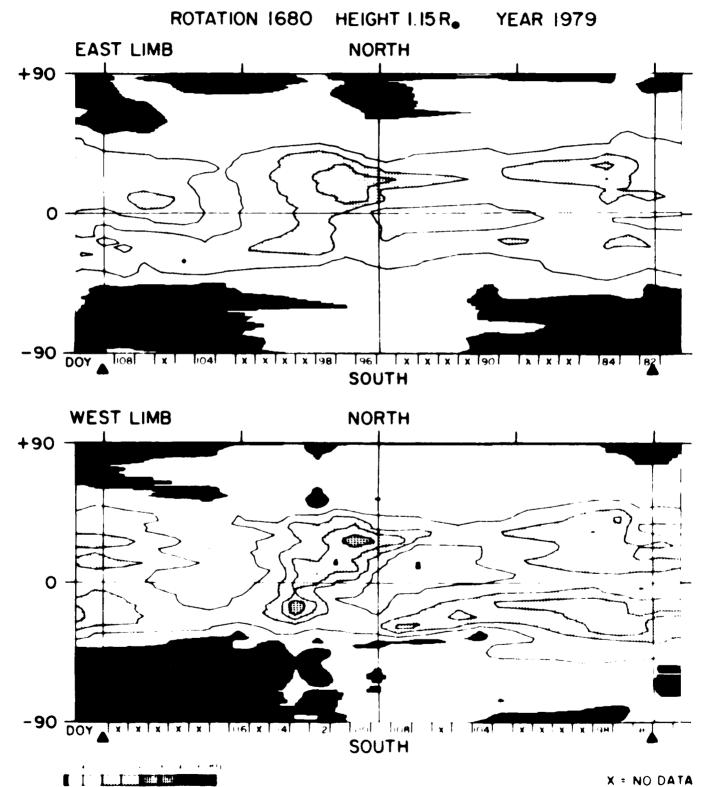


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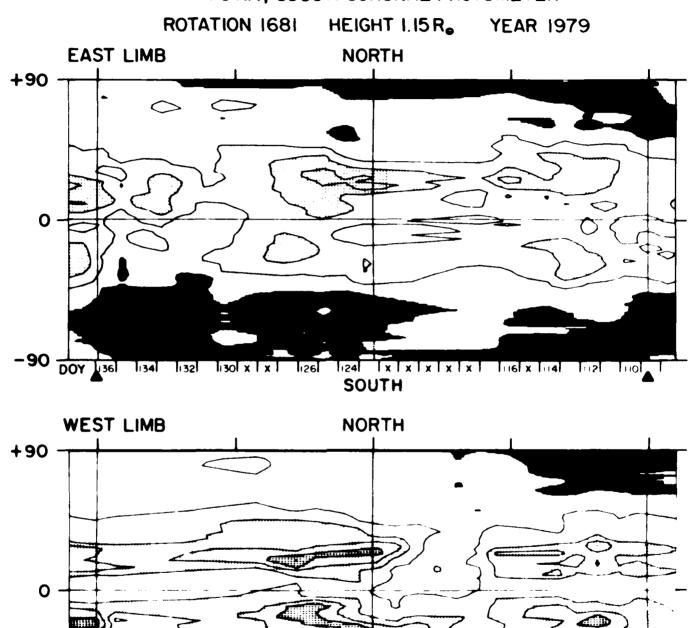
er de la de la viva de Mila.







Fe XIV, 5303 A CORONAL PHOTOMETER

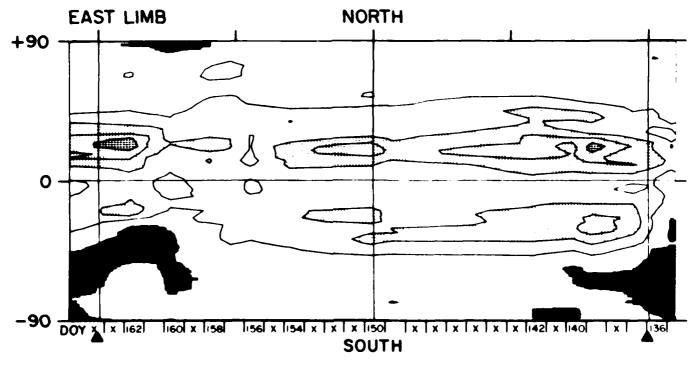


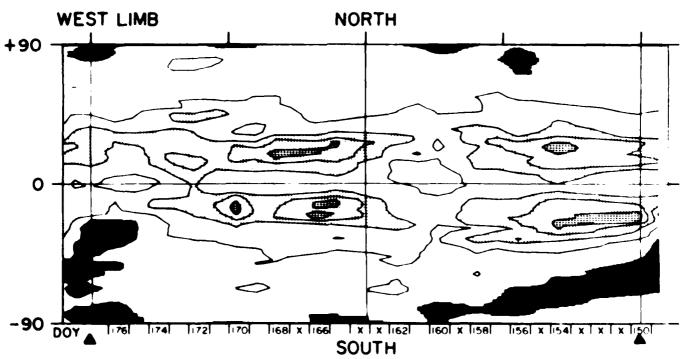
SOUTH

A 24 d MIC

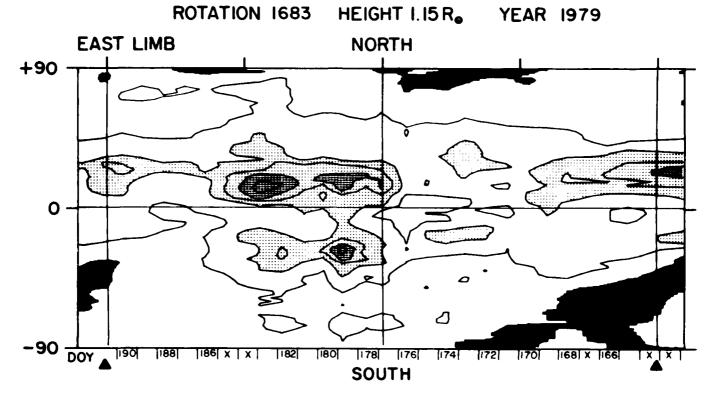
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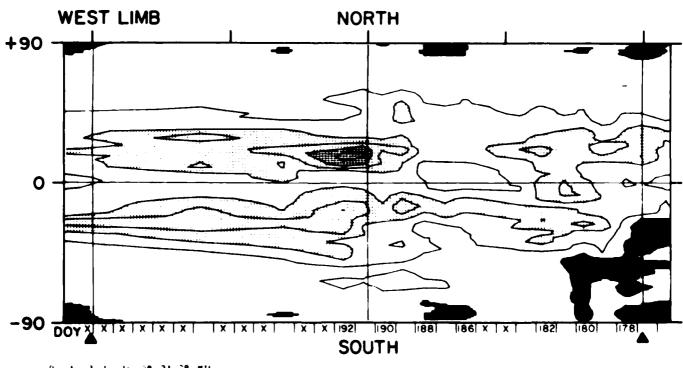
ROTATION 1682 HEIGHT 1.15 R. YEAR 1979



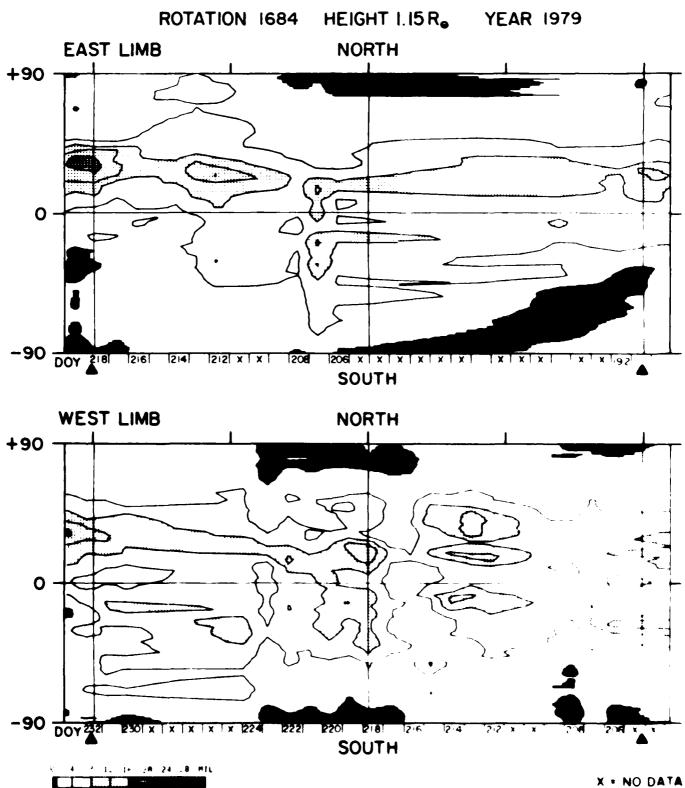


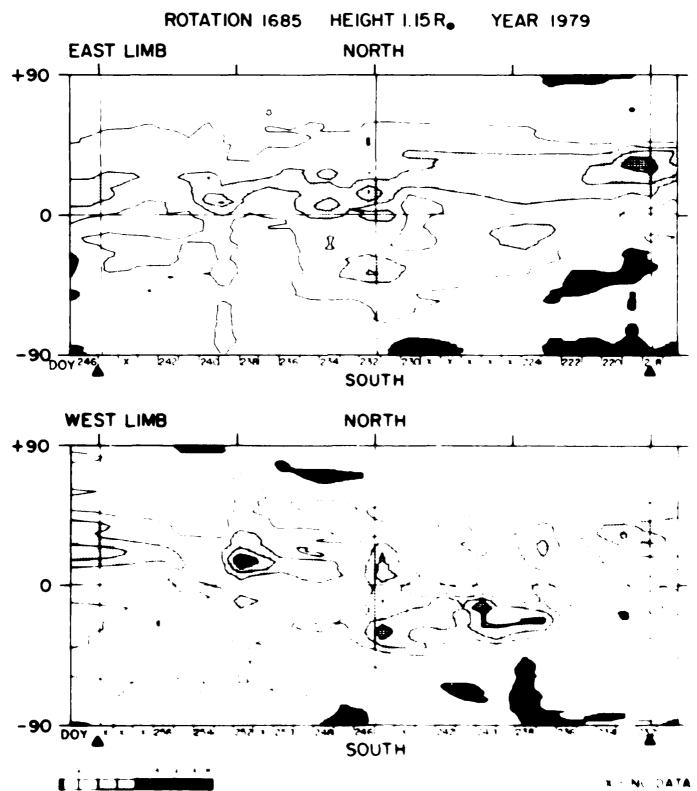
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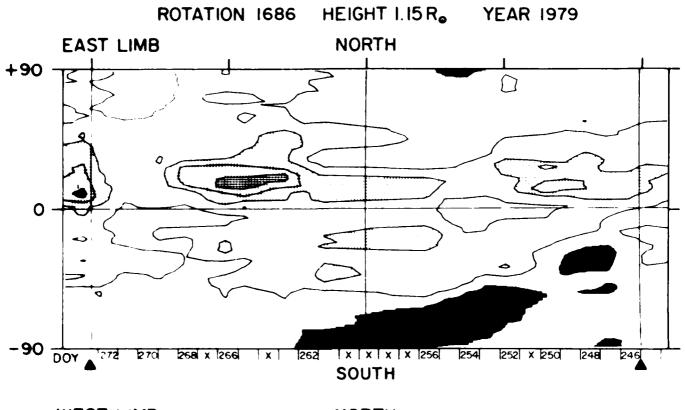


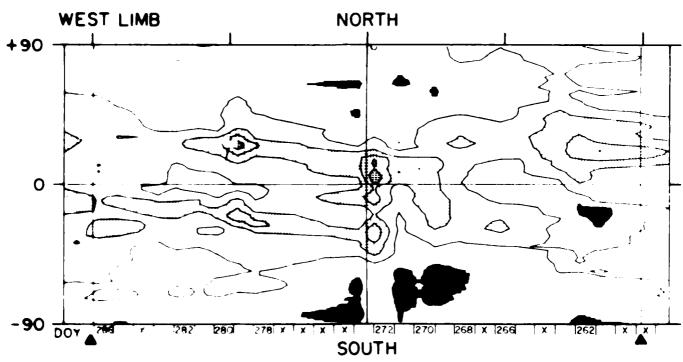
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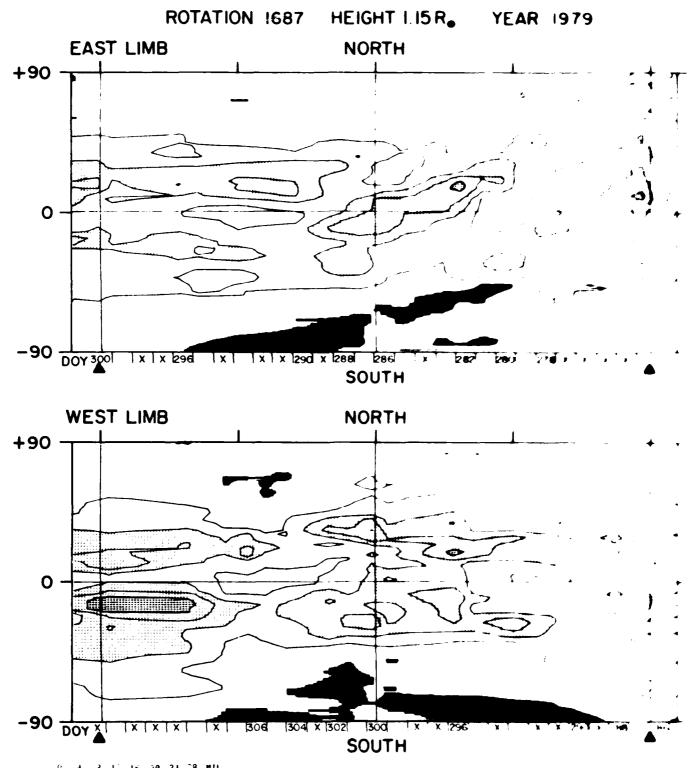




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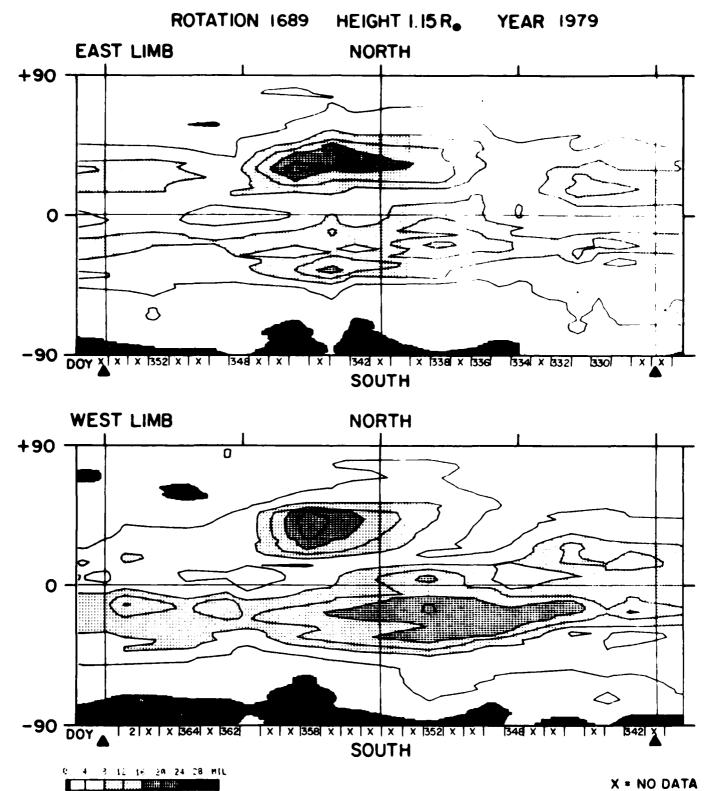




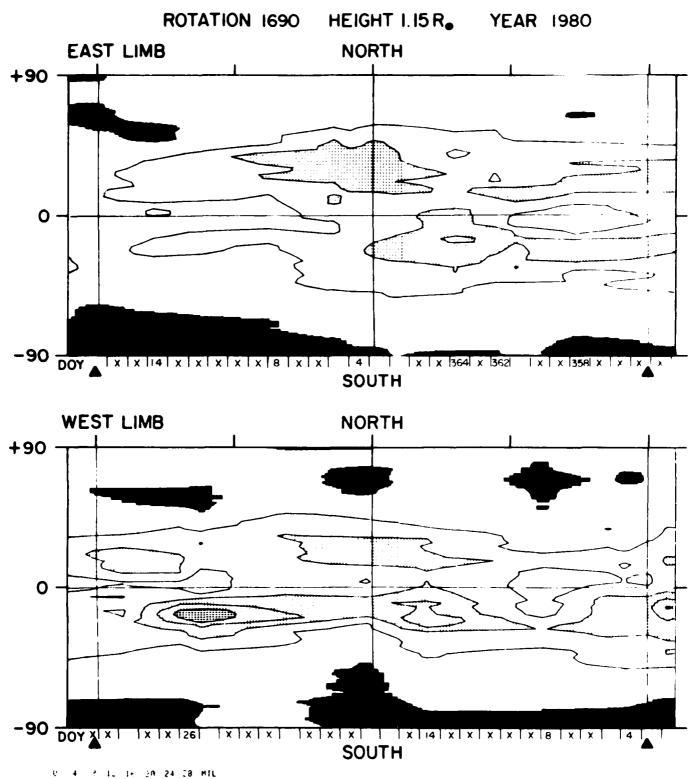
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HEIGHT | ISR. YEAR 1979 ROTATION 1688 NORTH EAST LIMB NORTH SOUTH X . NO DATA

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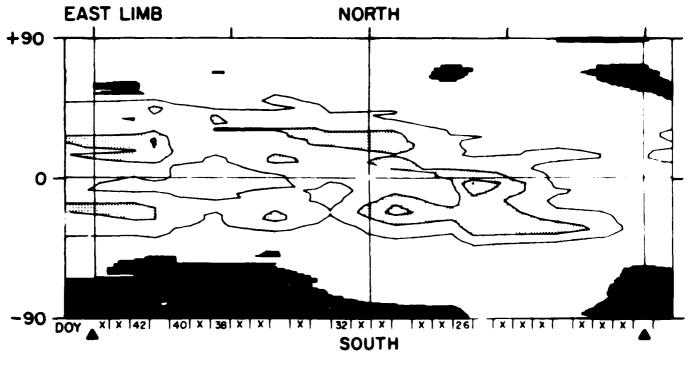


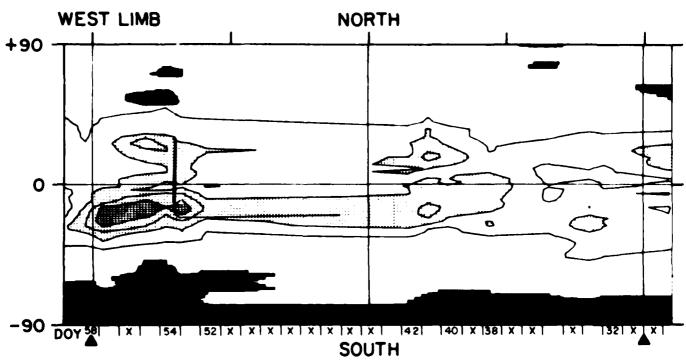
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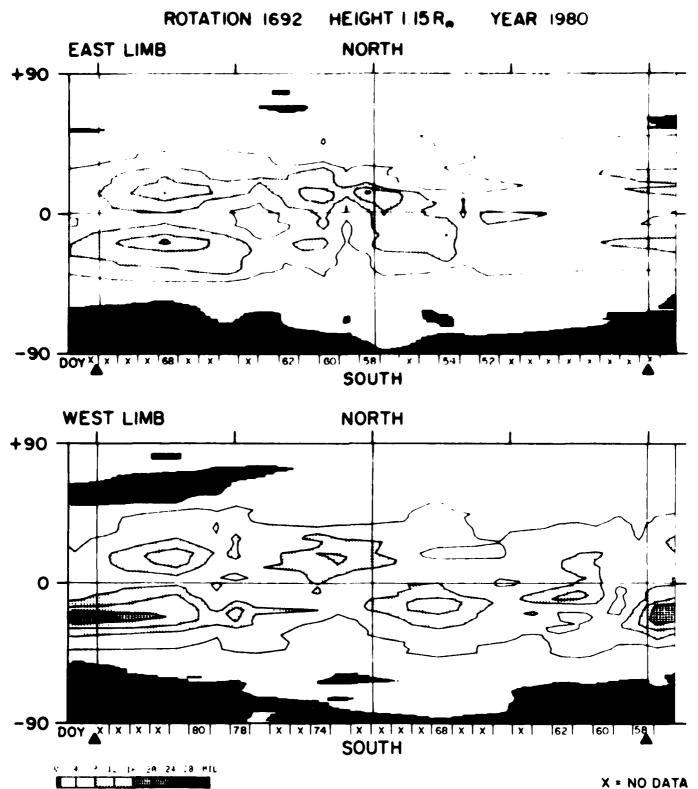
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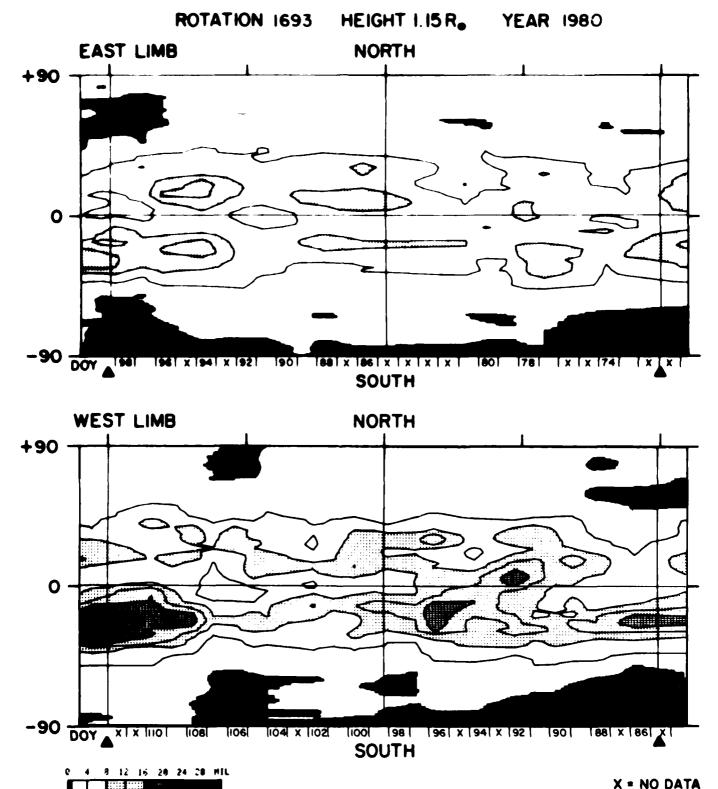
ROTATION 1691 HEIGHT 1.15 R. YEAR 1980



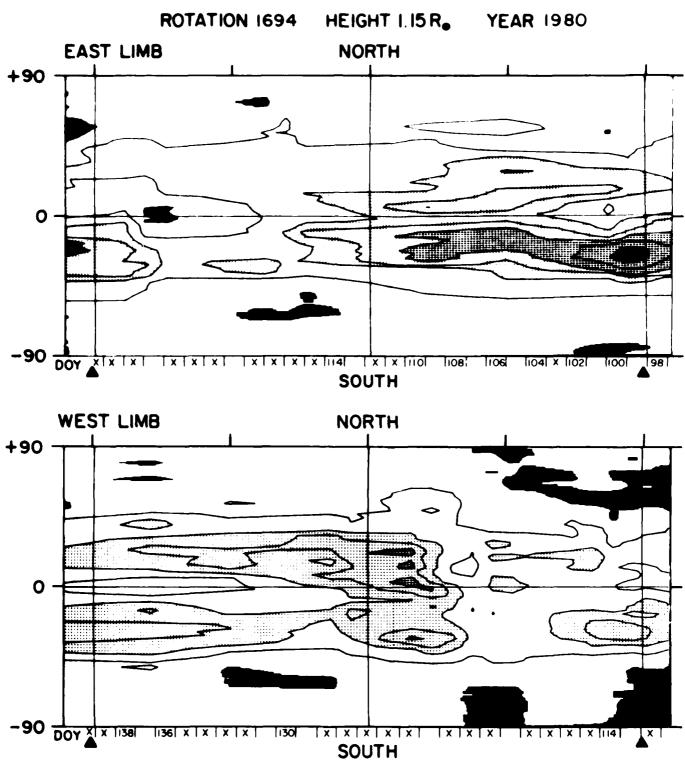


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Fe XIV, 5303 A CORONAL PHOTOMETER

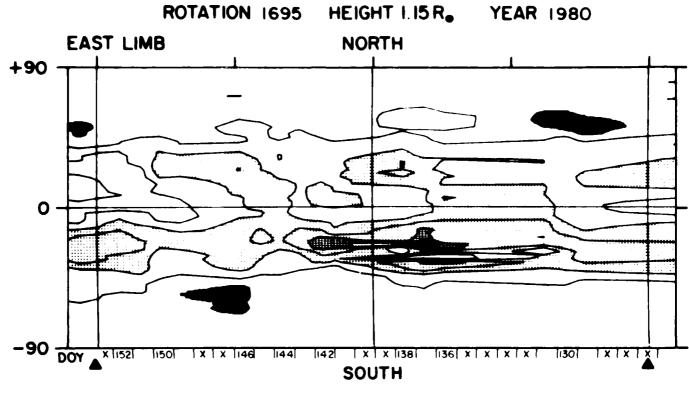


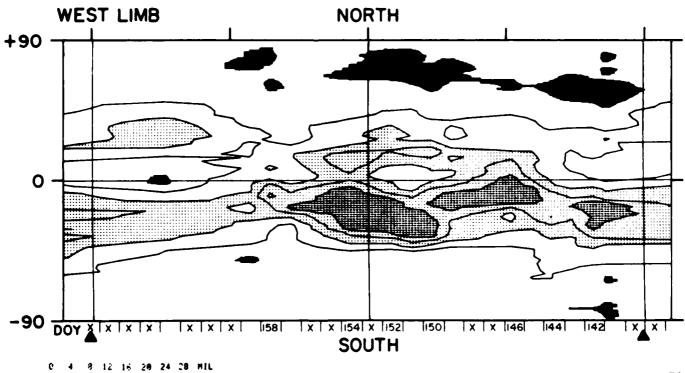
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X = NO DATA

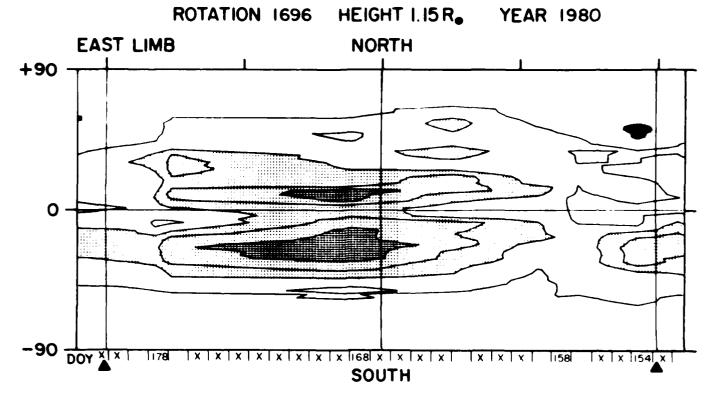
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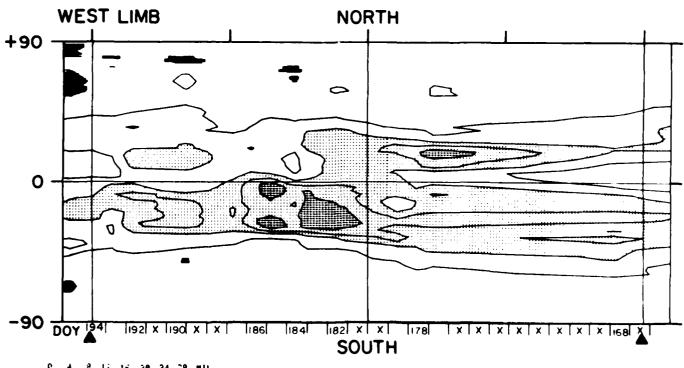
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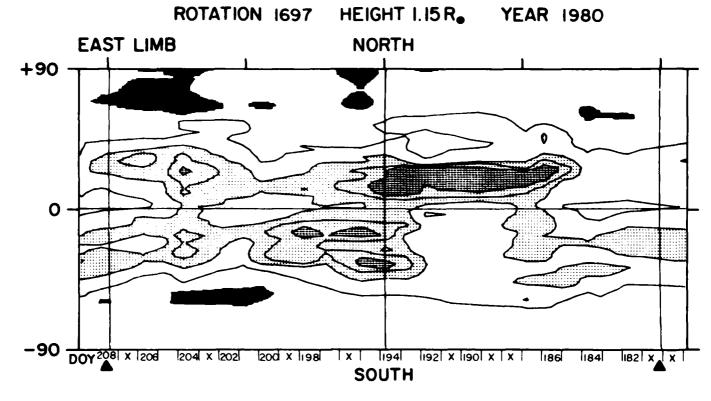
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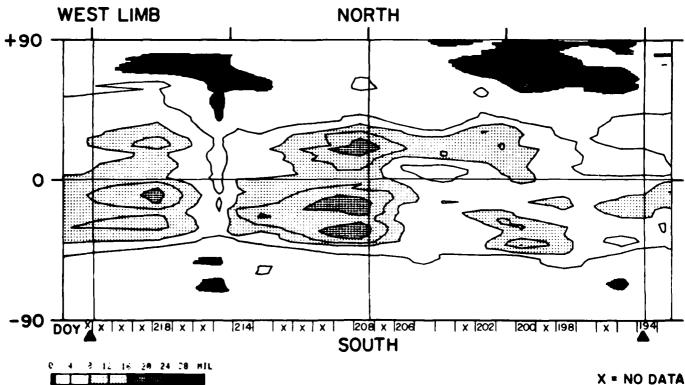


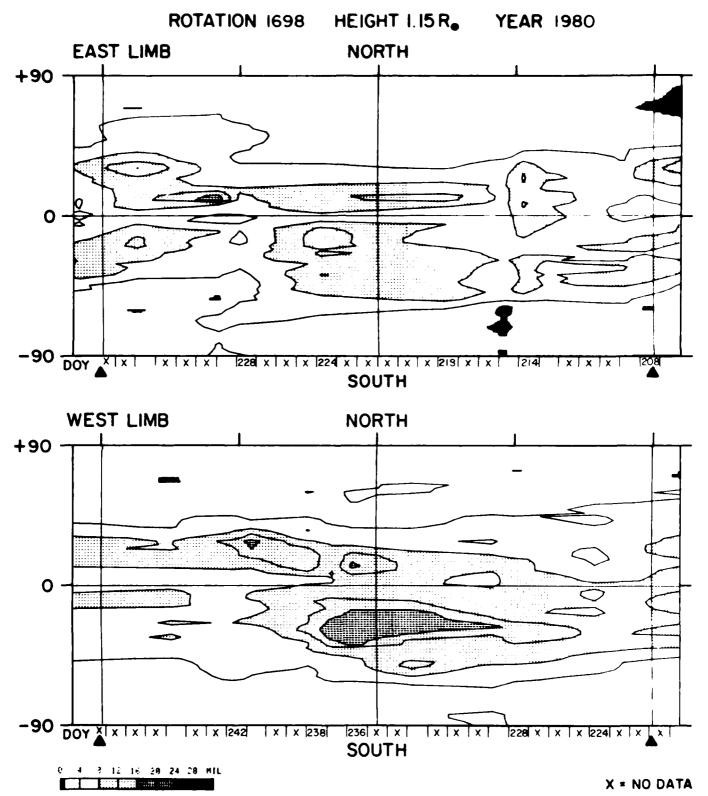


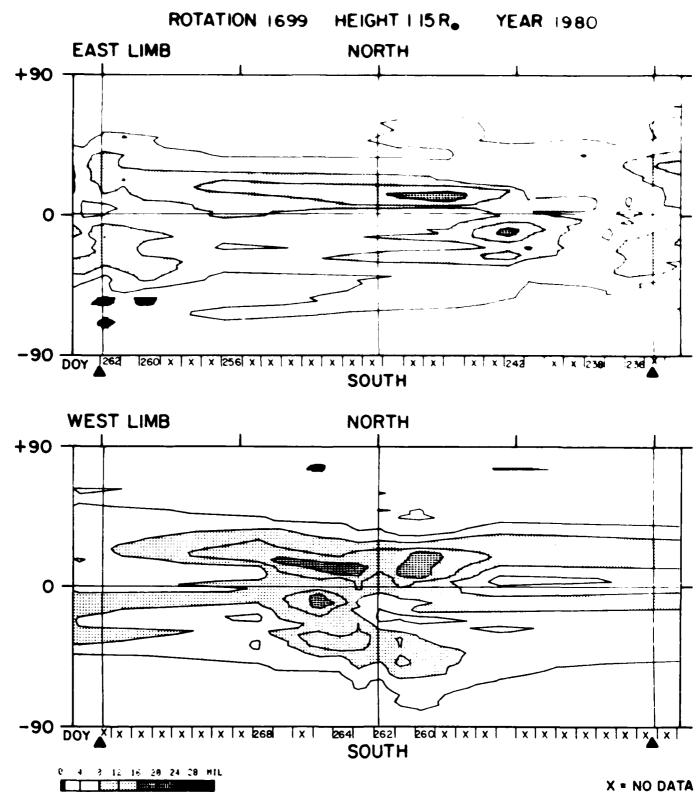
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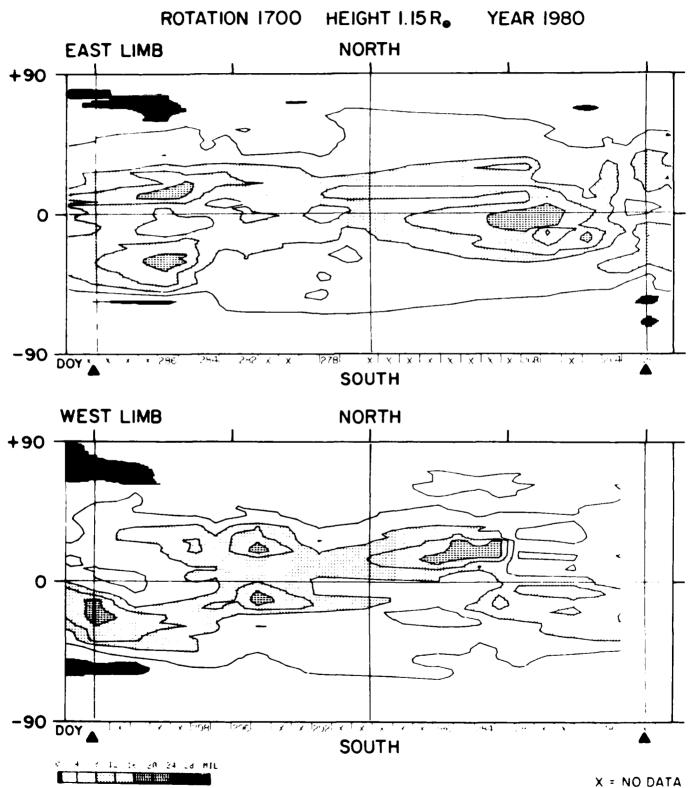
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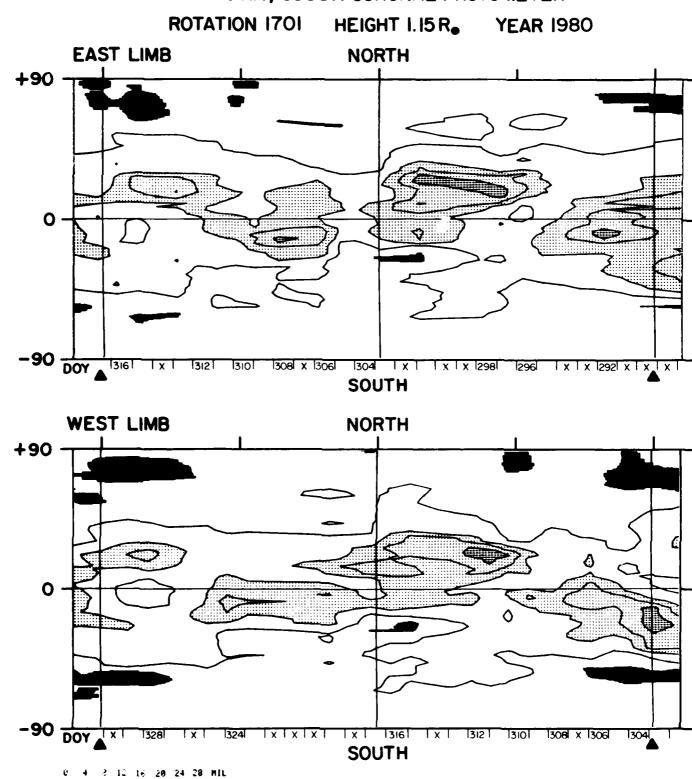


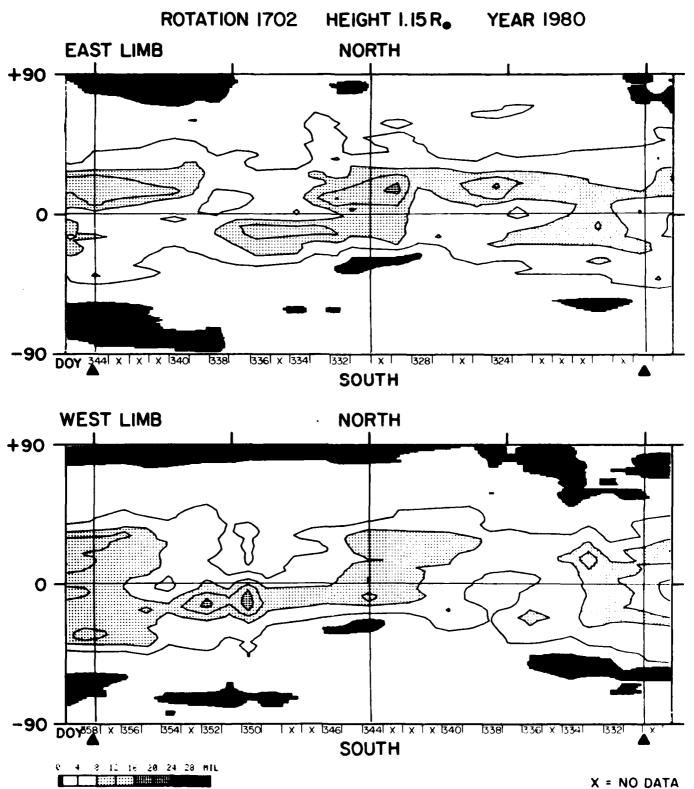


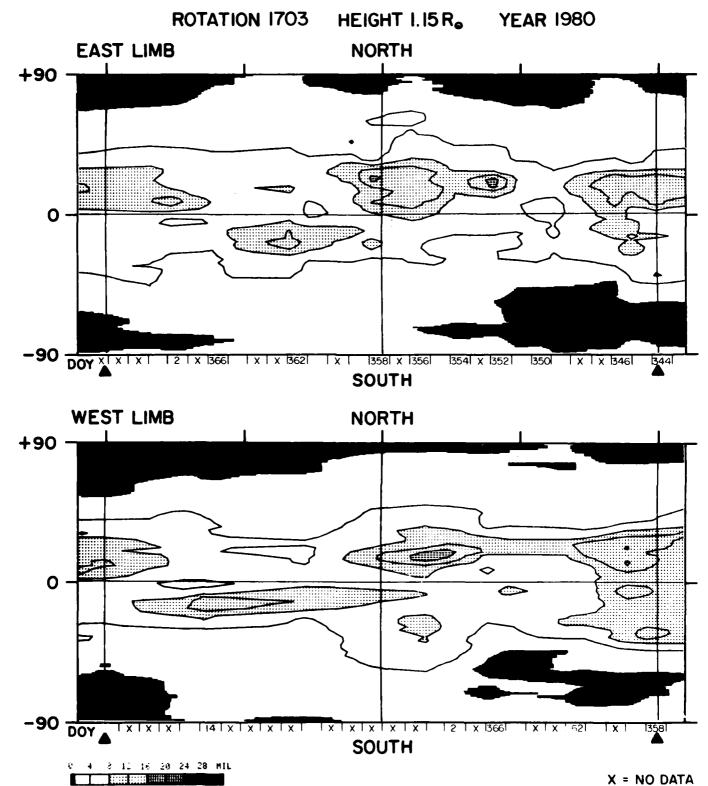


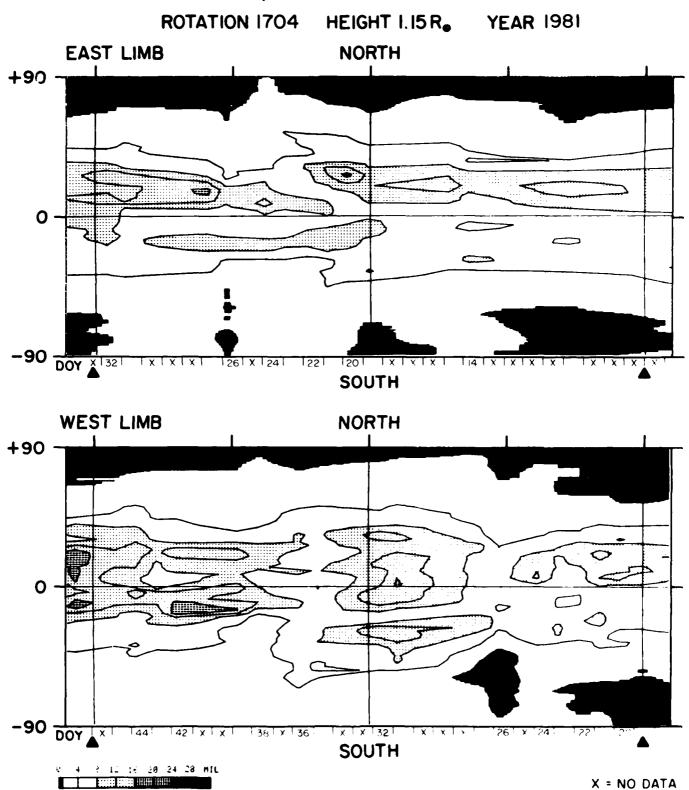


Fe XIV, 5303 & CORONAL PHOTOMETER

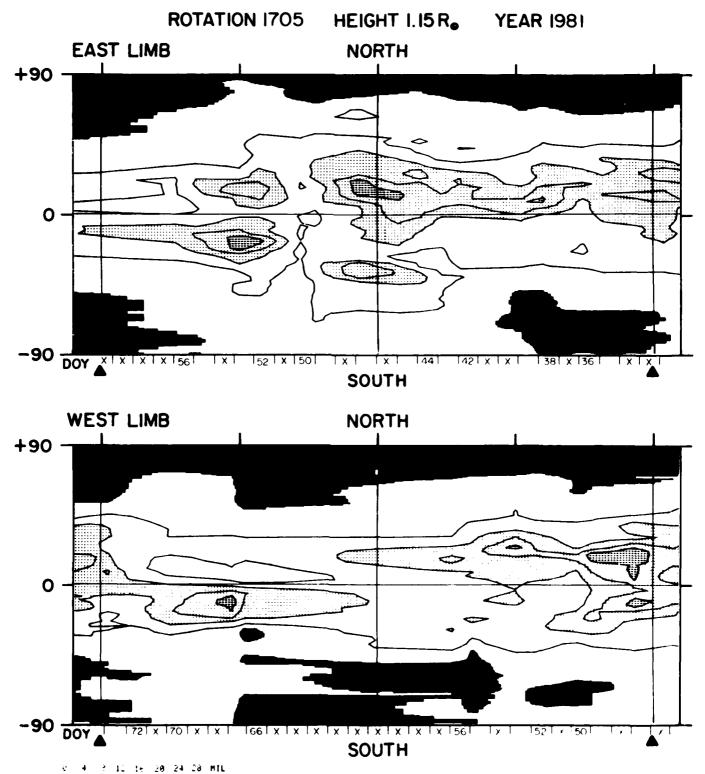


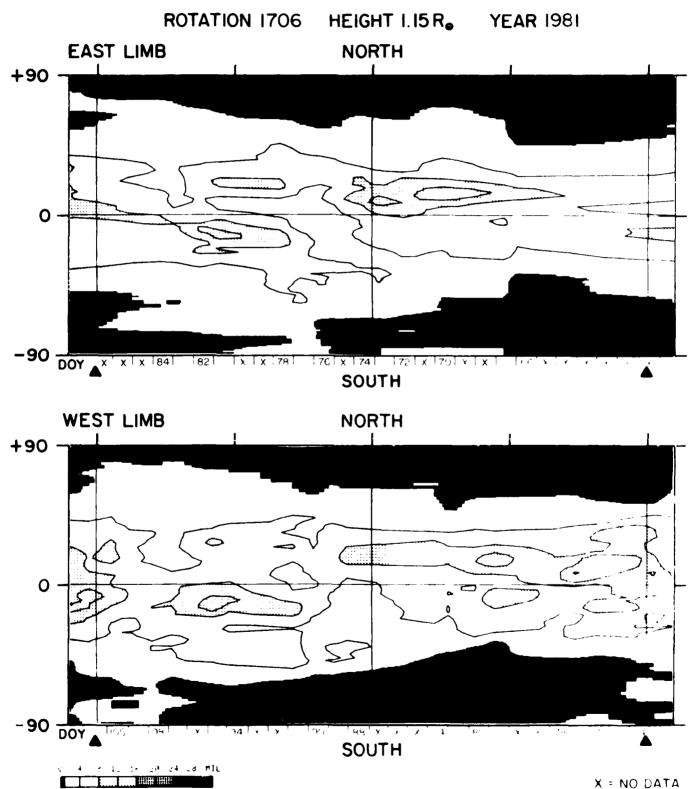




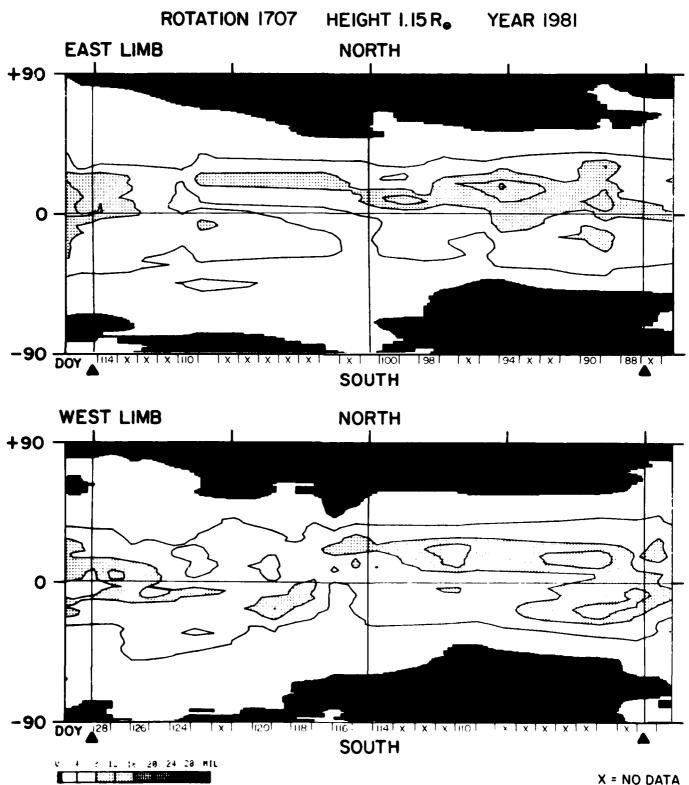


Fe XIV, 5303 Å CORONAL PHOTOMETER

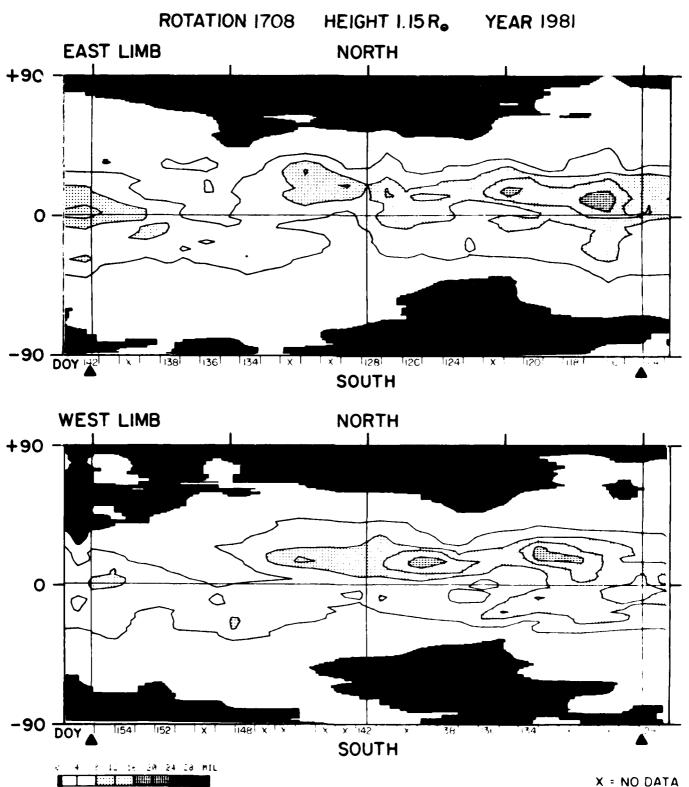




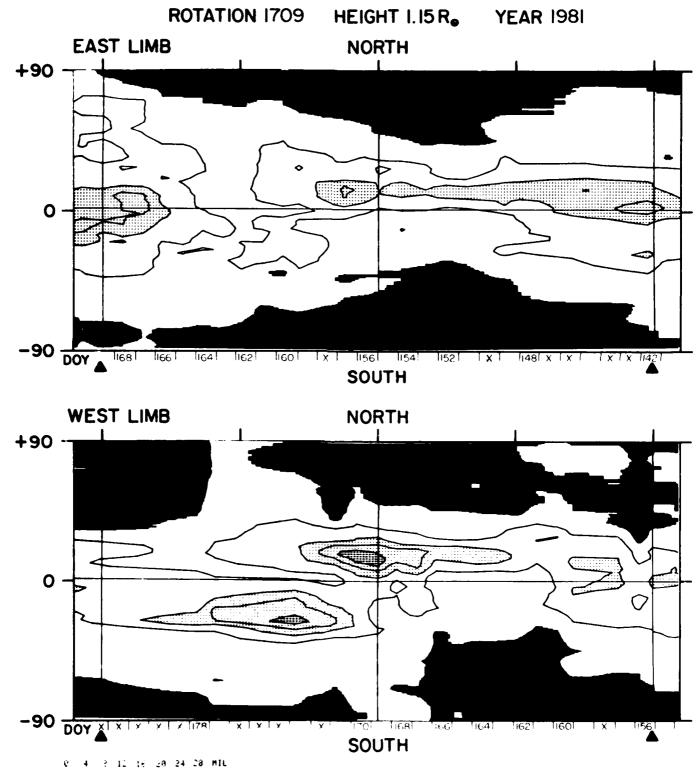
Fe XIV, 5303 $\mbox{\normal}$ CORONAL PHOTOMETER



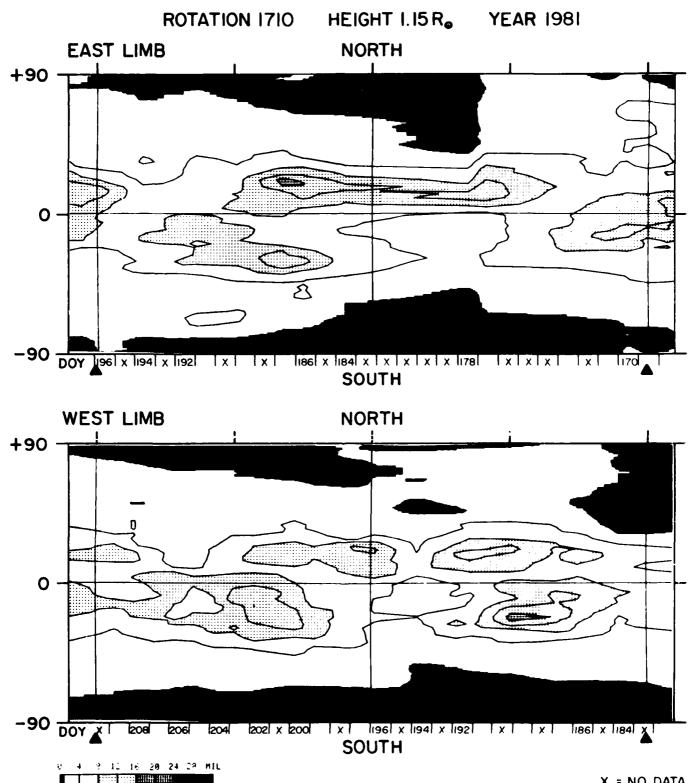
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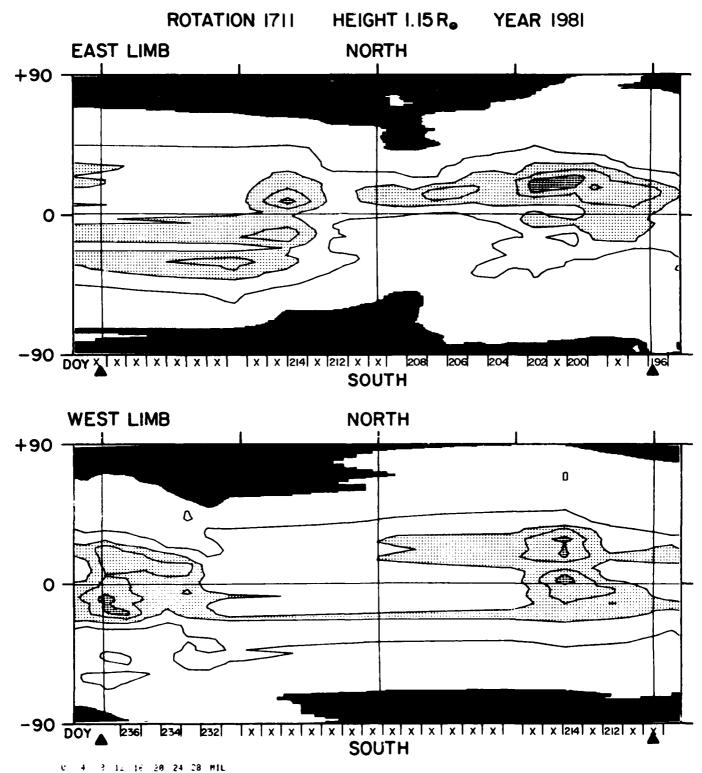
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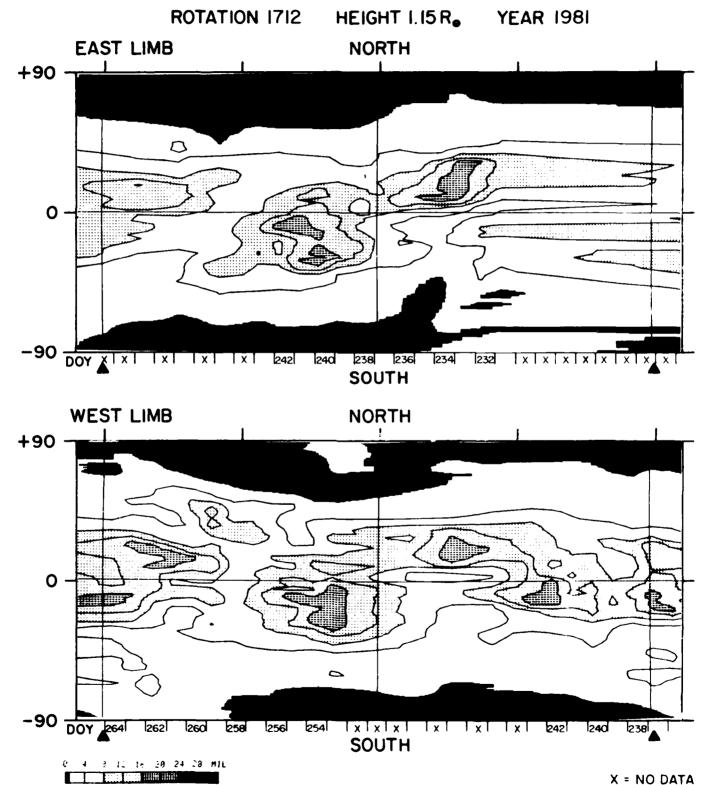


Fe XIV, 5303 & CORONAL PHOTOMETER

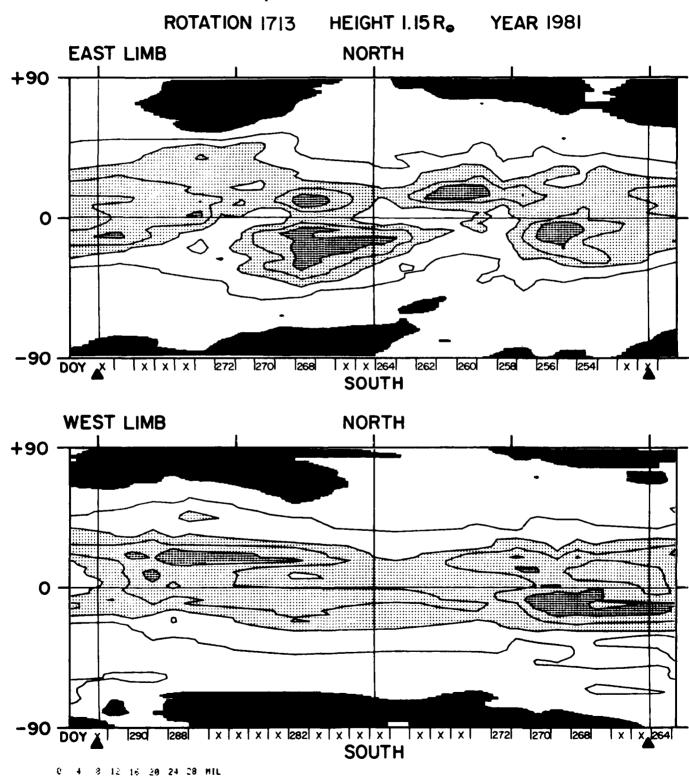


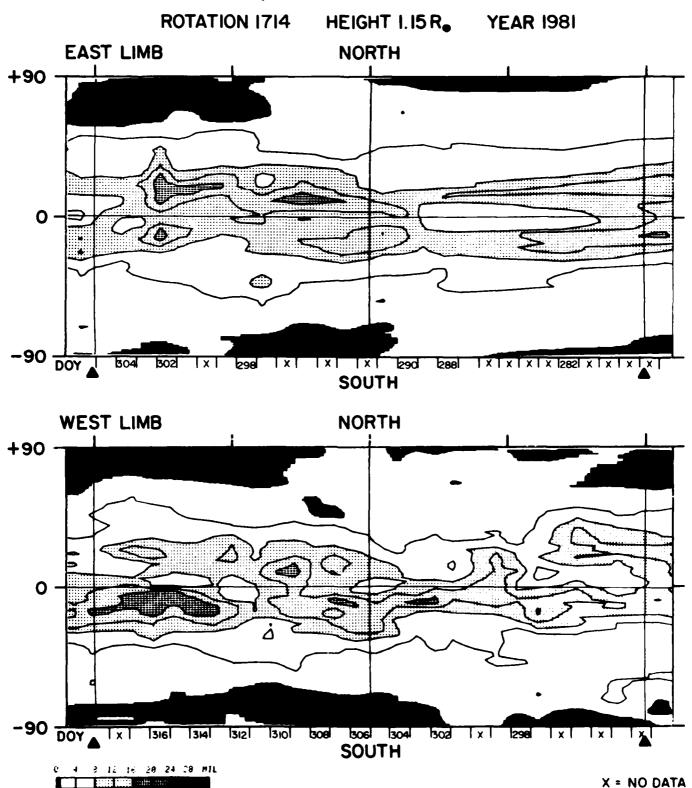
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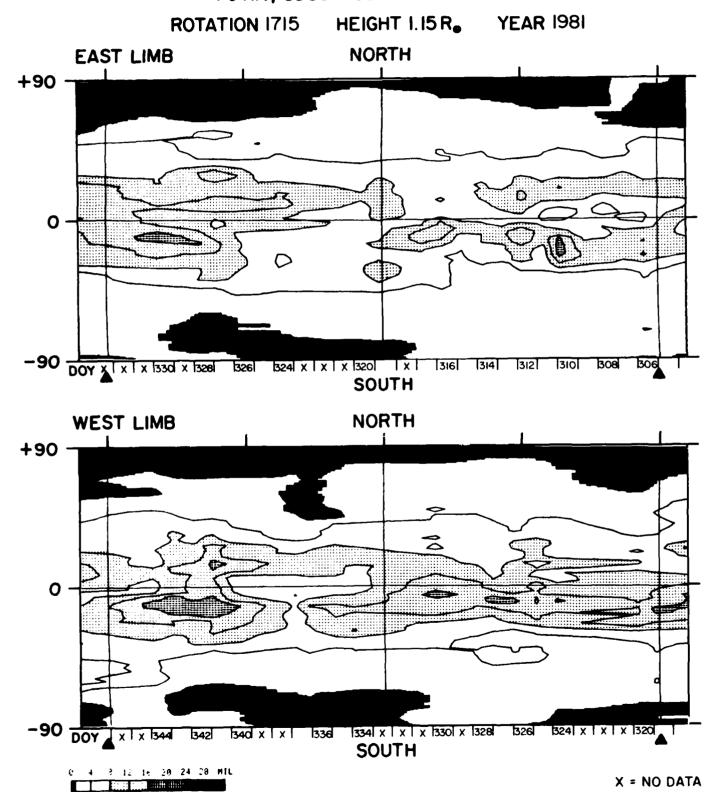




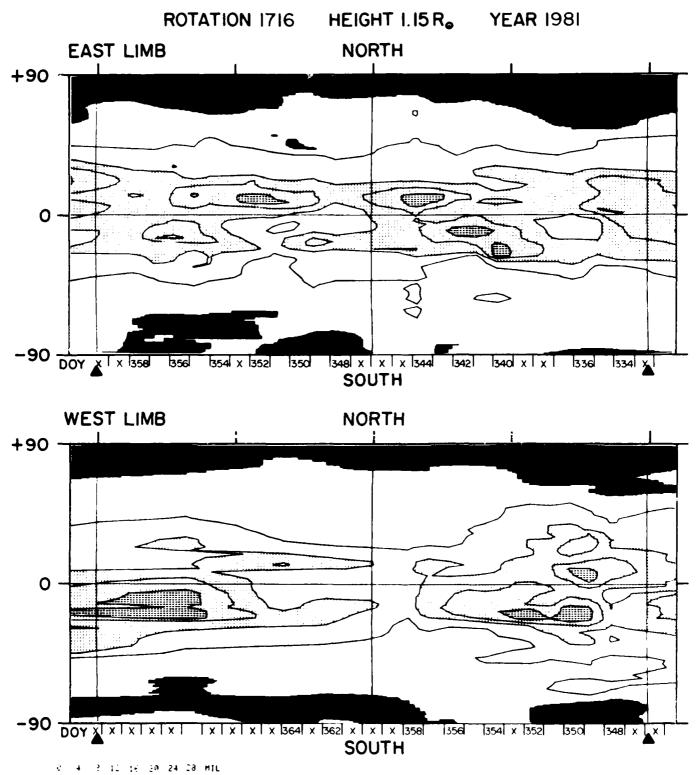
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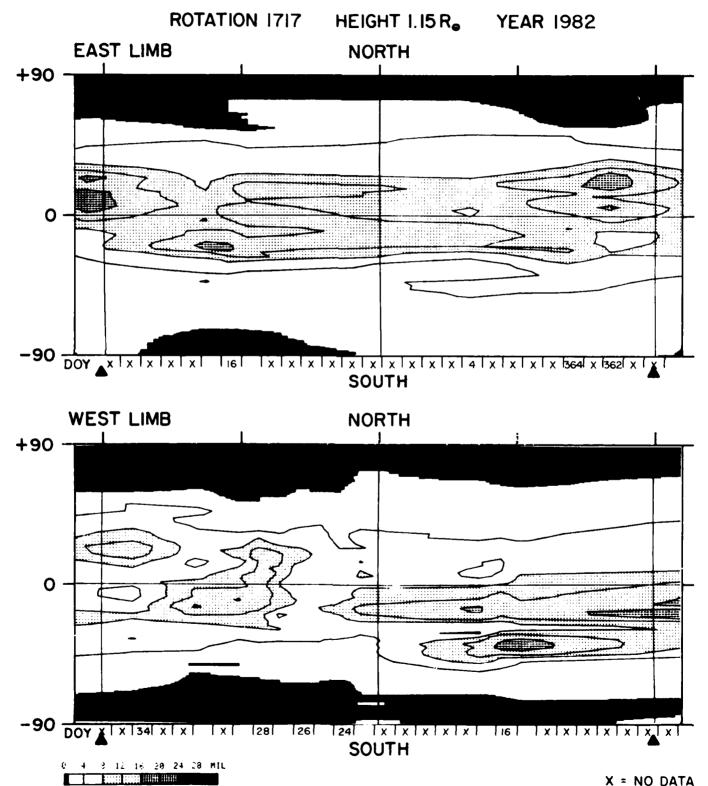


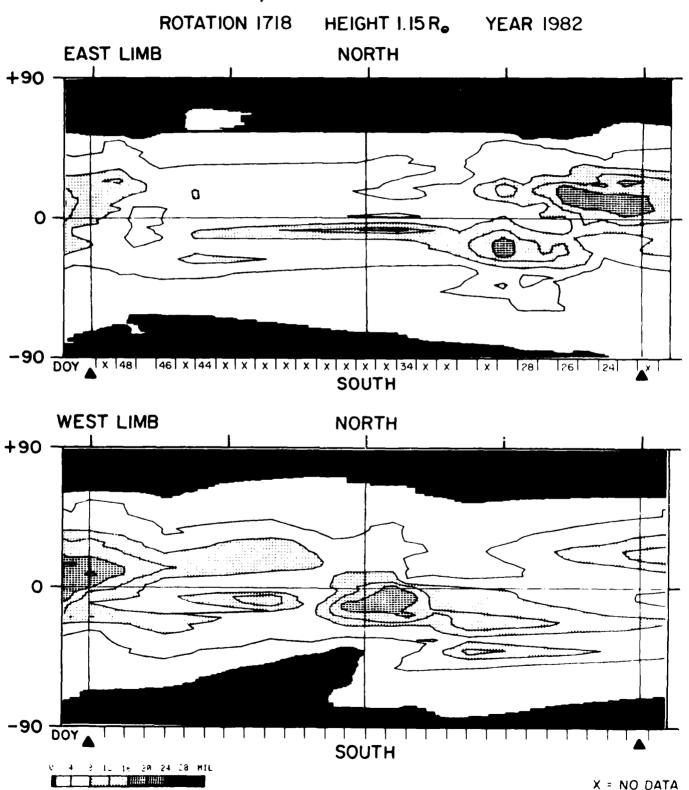


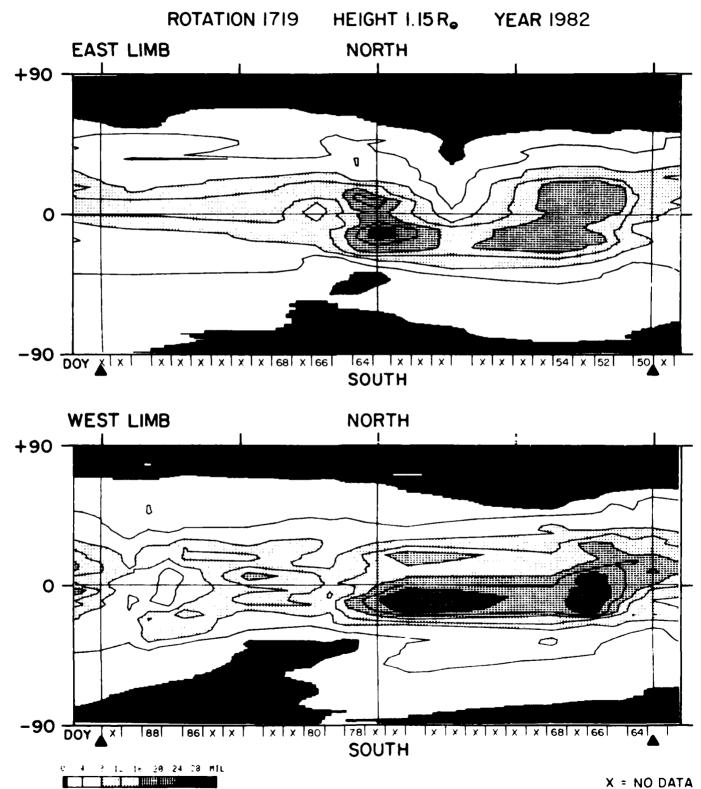


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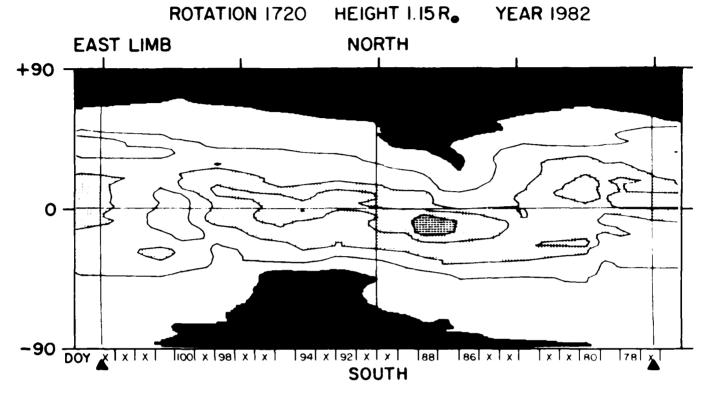


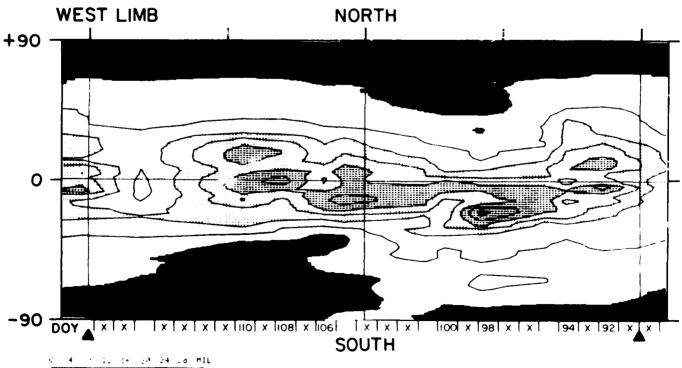


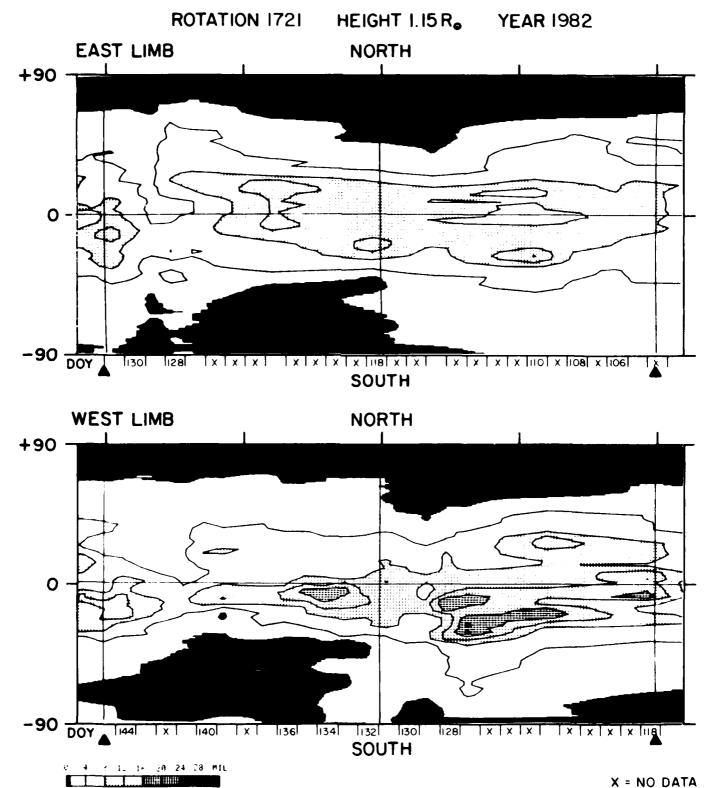




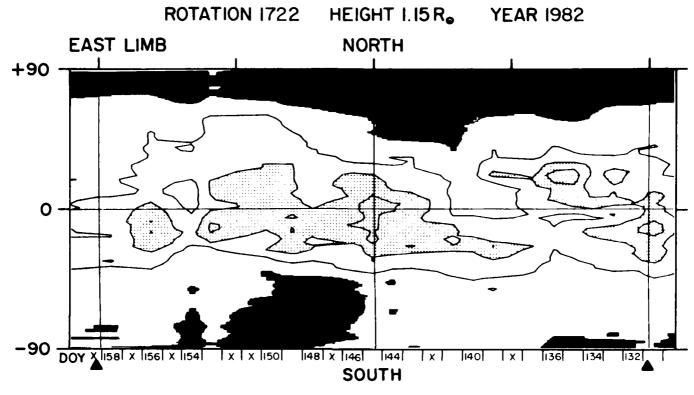
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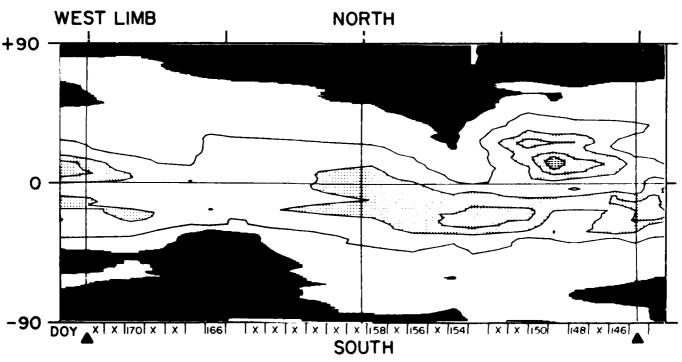






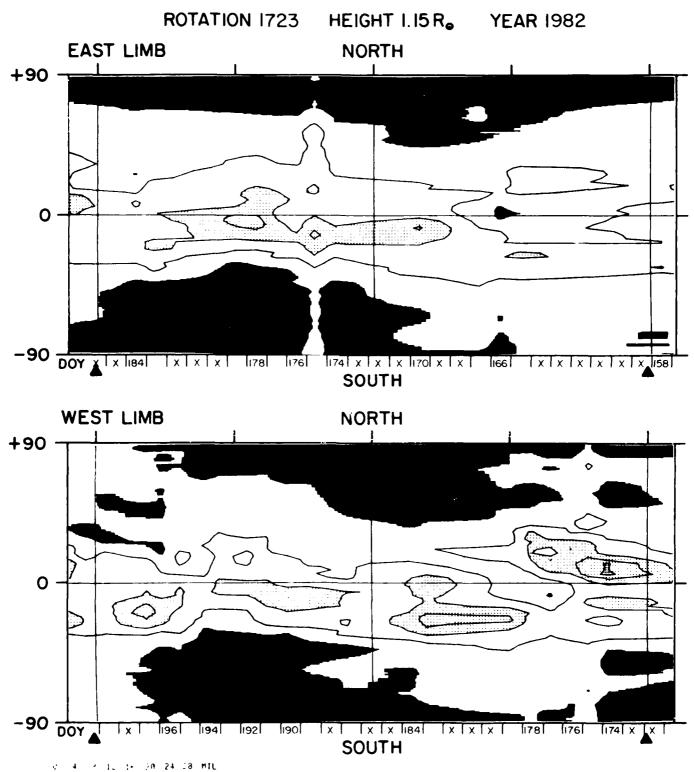
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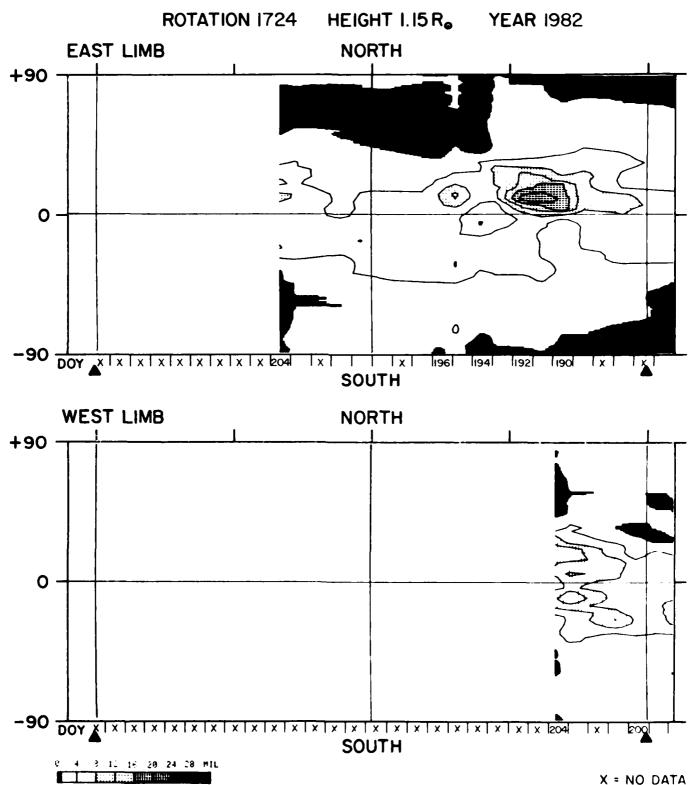




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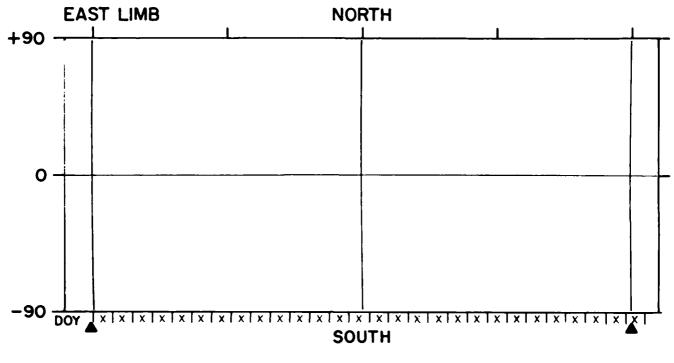
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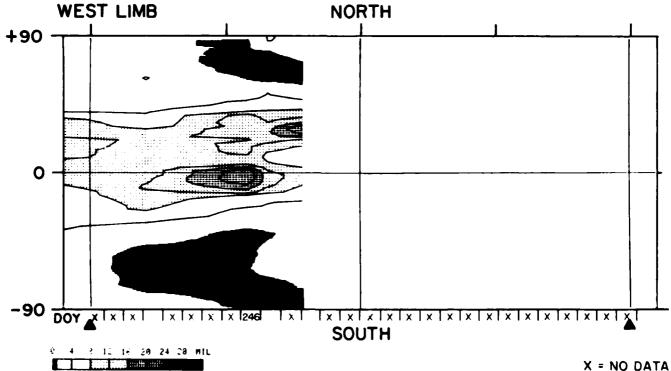




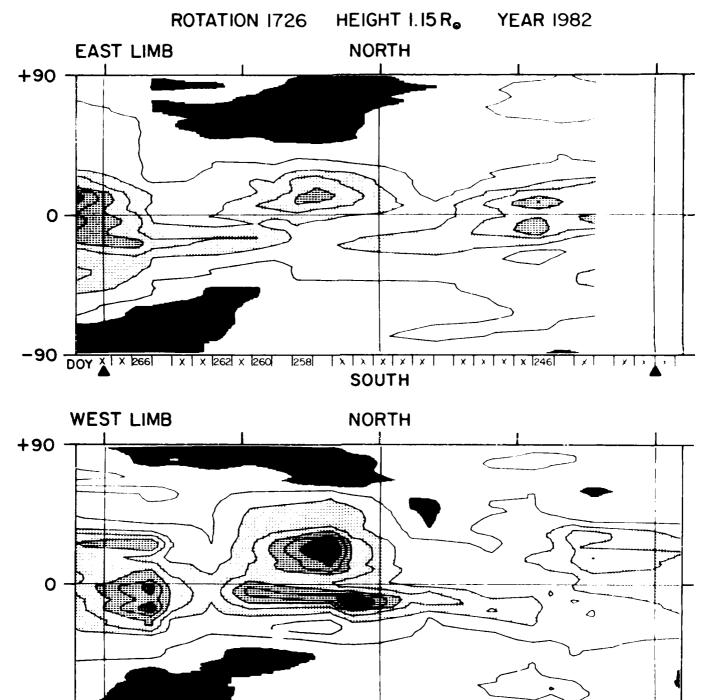
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ROTATION 1725 HEIGHT 1.15 R. YEAR 1982

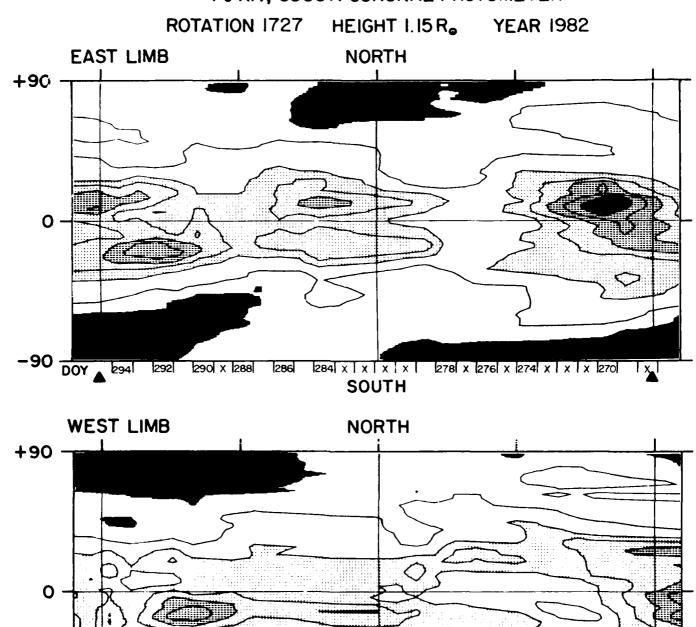




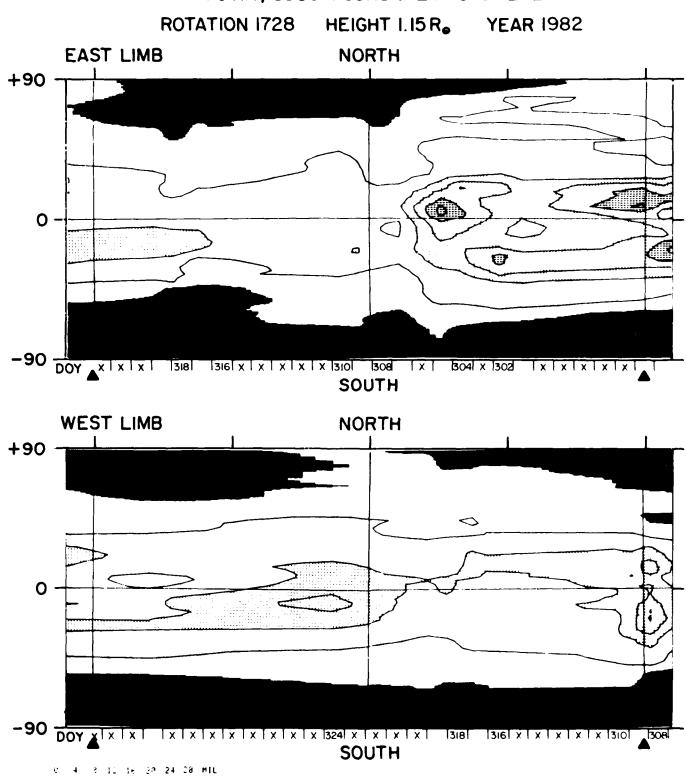
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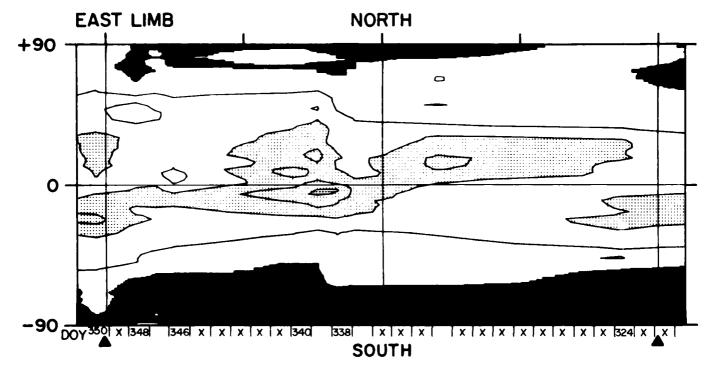


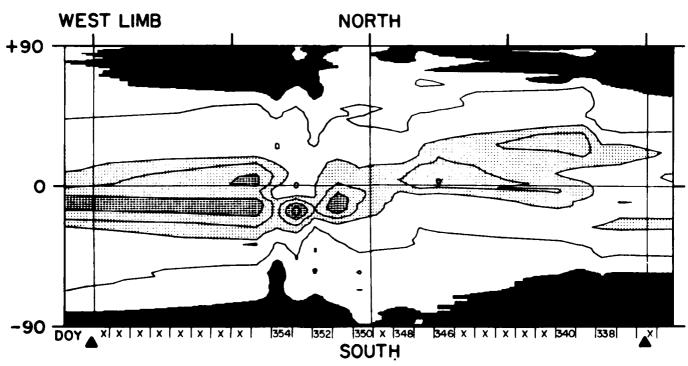
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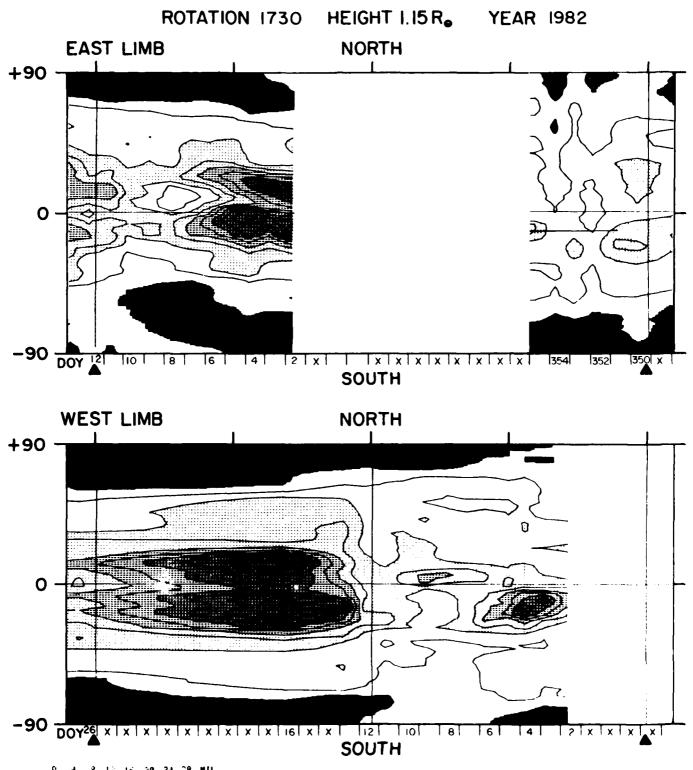
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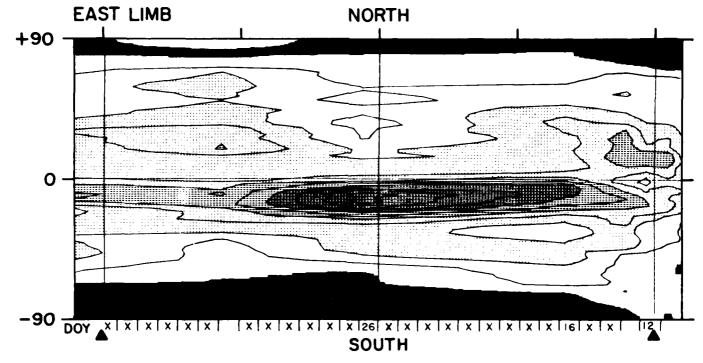
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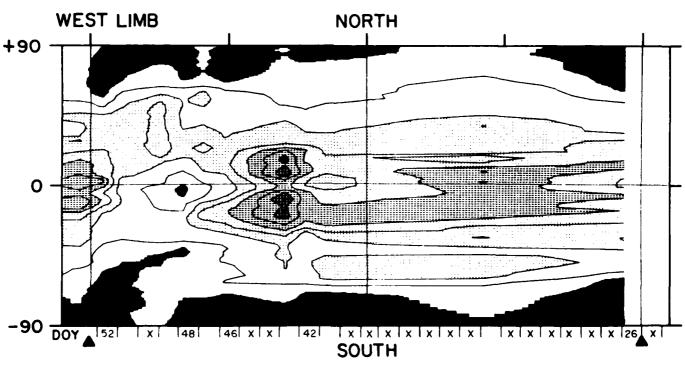
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Fe XIV, 5303 A CORONAL PHOTOMETER

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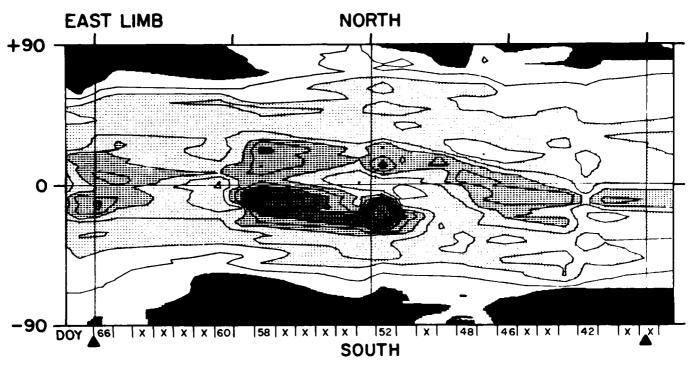


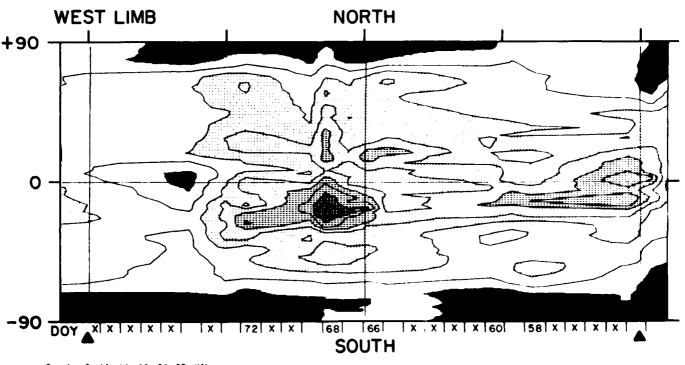


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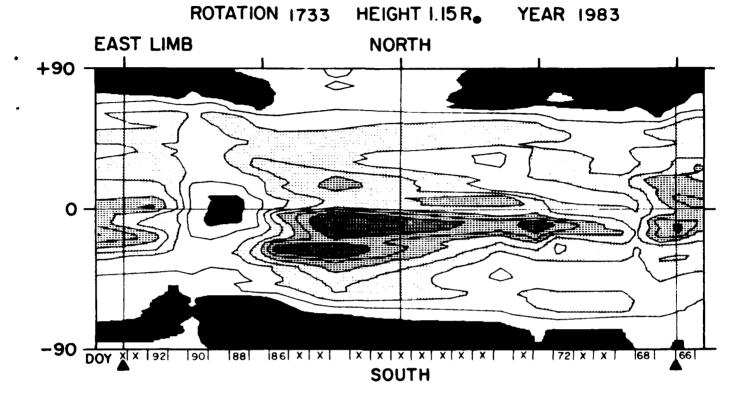
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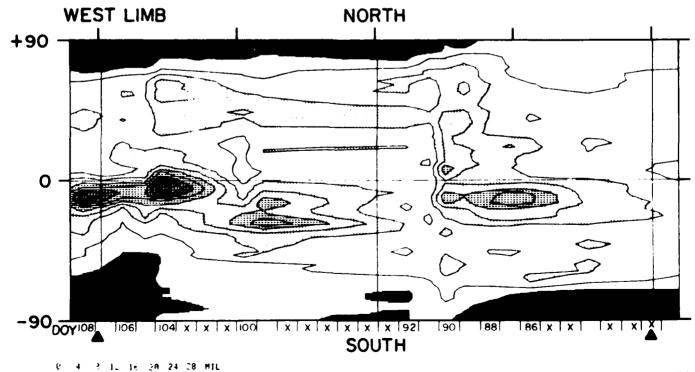




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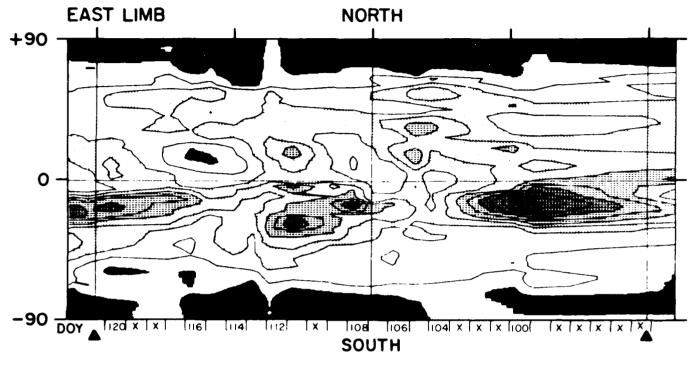
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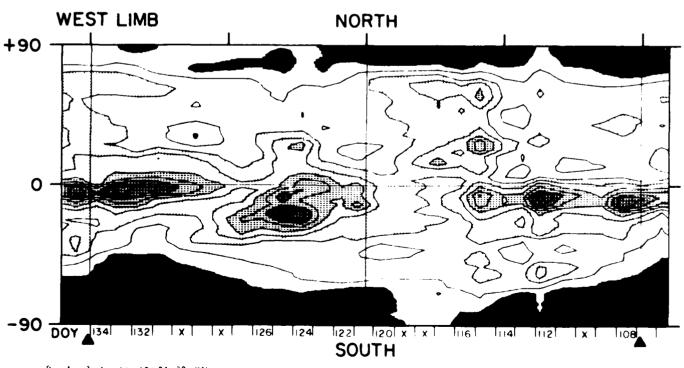




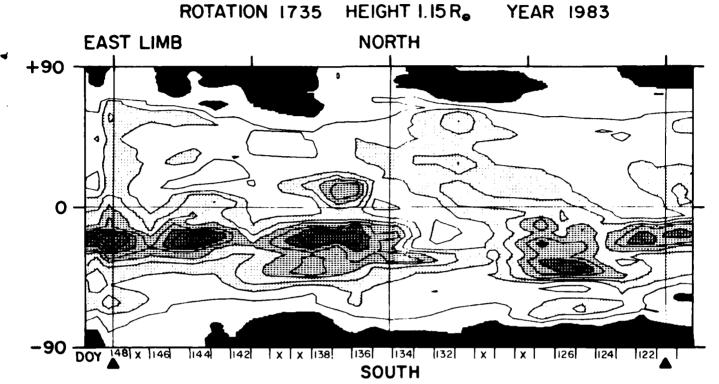
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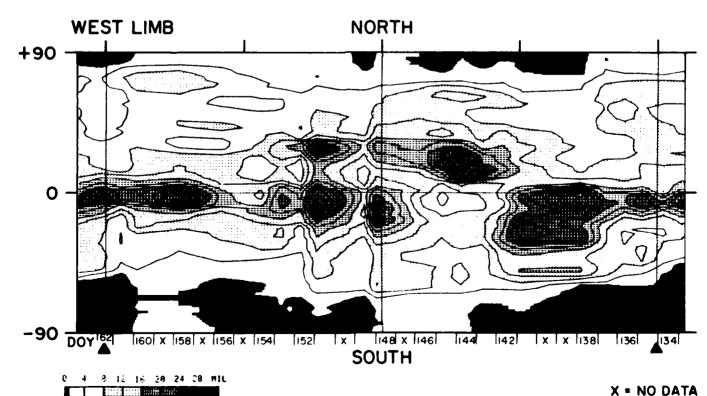
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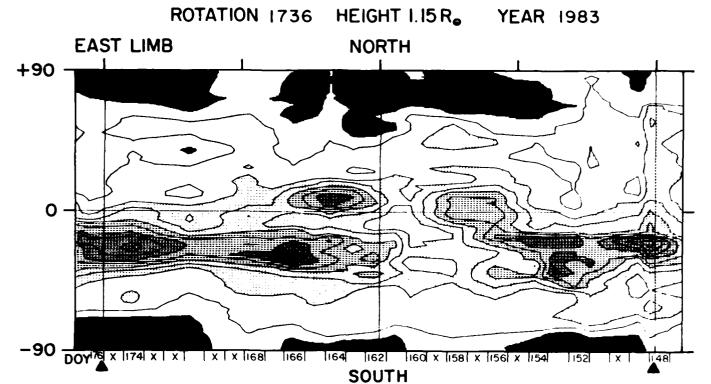


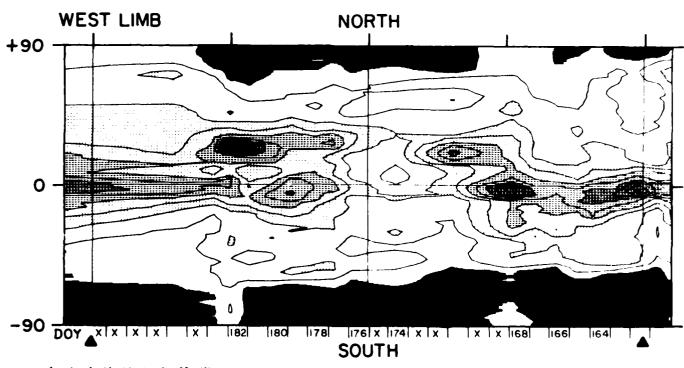
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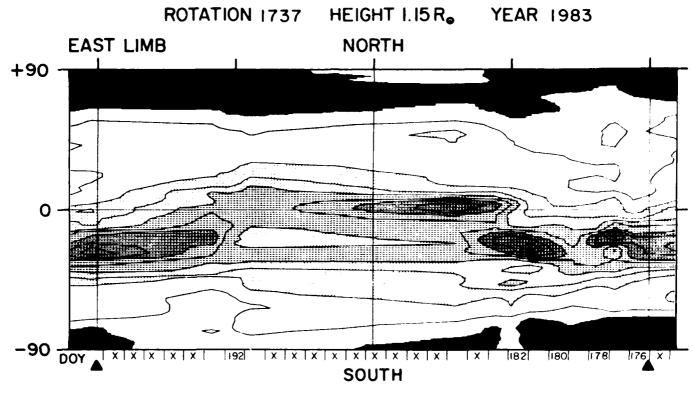
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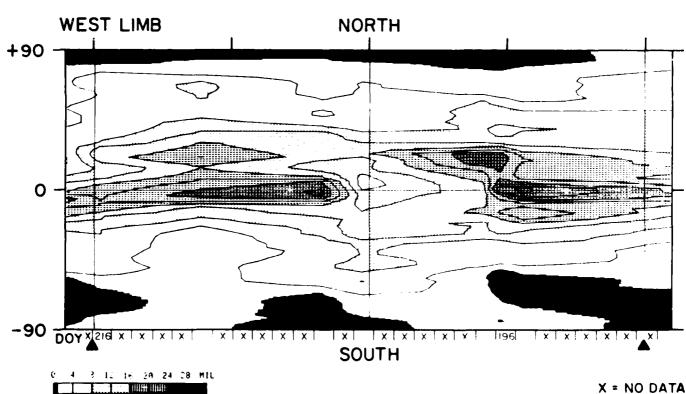




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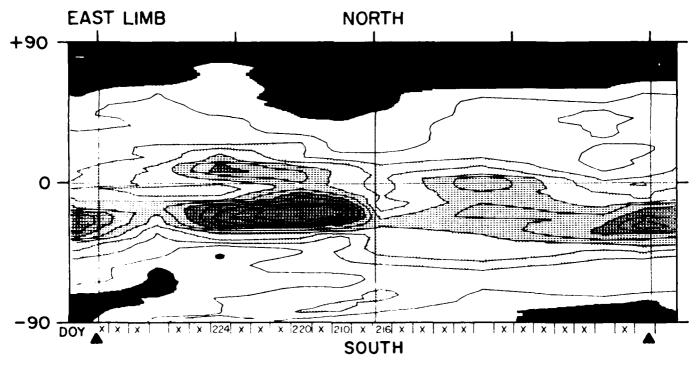
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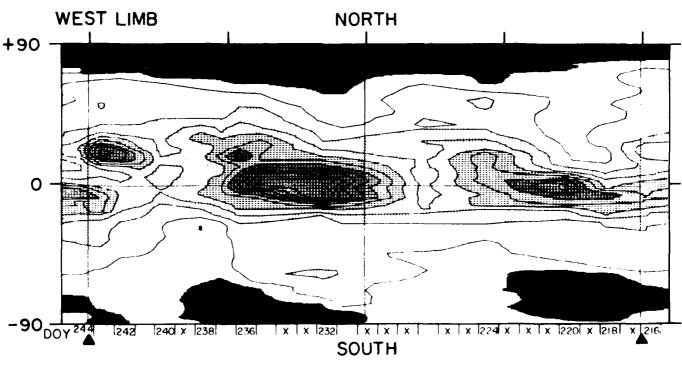




Fe XIV, 5303 Å CORONAL PHOTOMETER

ROTATION 1738 HEIGHT 1.15 R. YEAR 1983

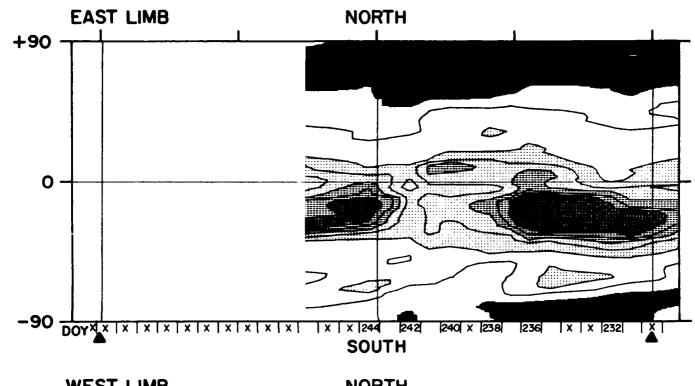


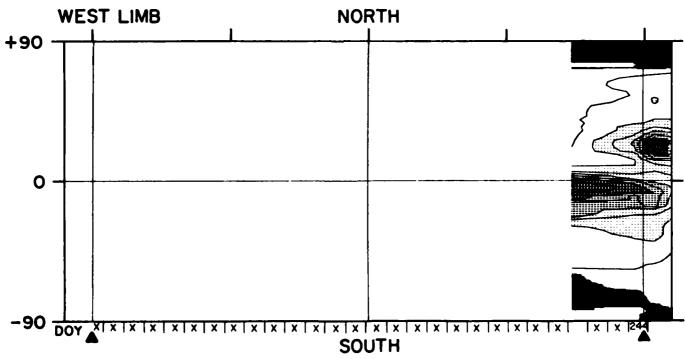


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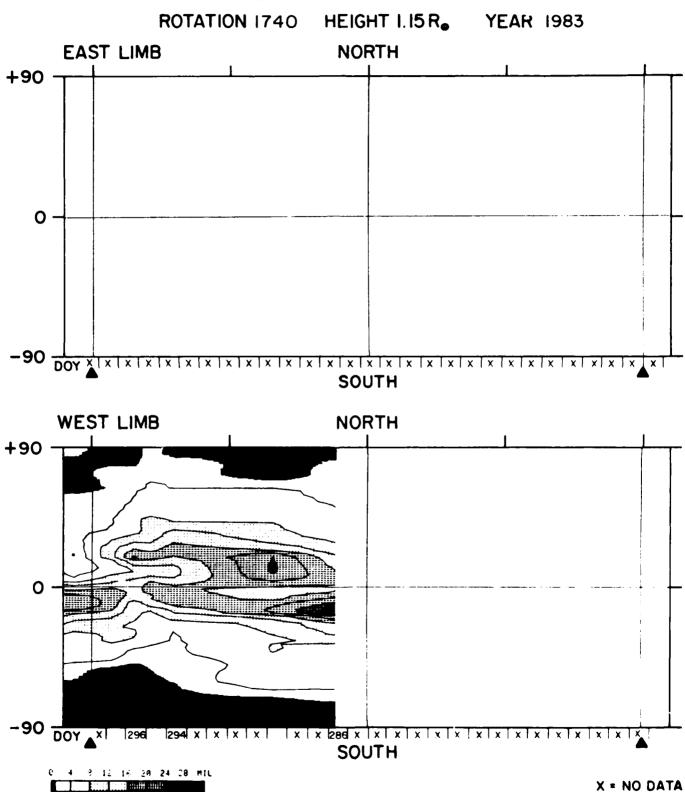
Fe XIV, 5303 & CORONAL PHOTOMETER

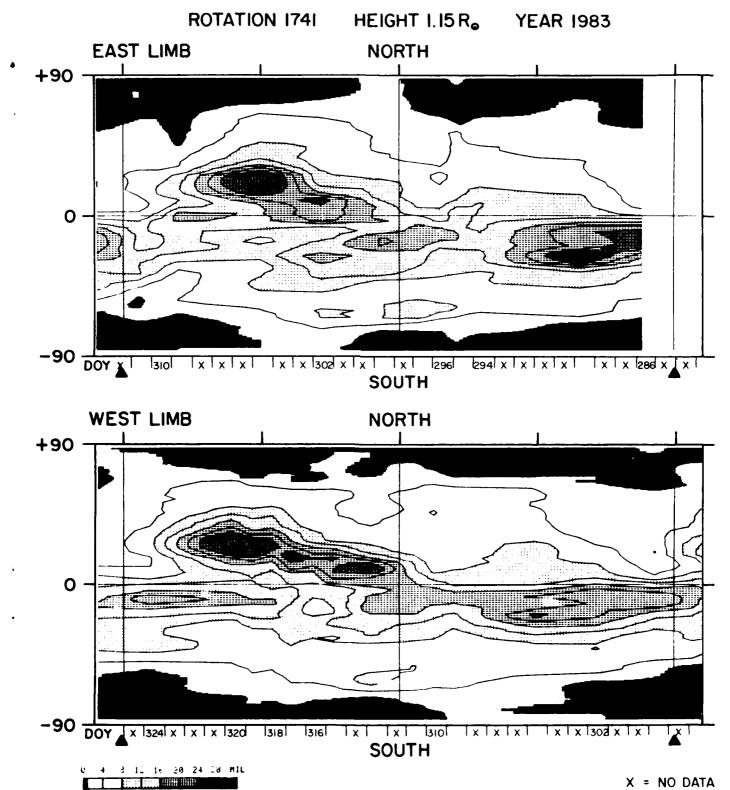
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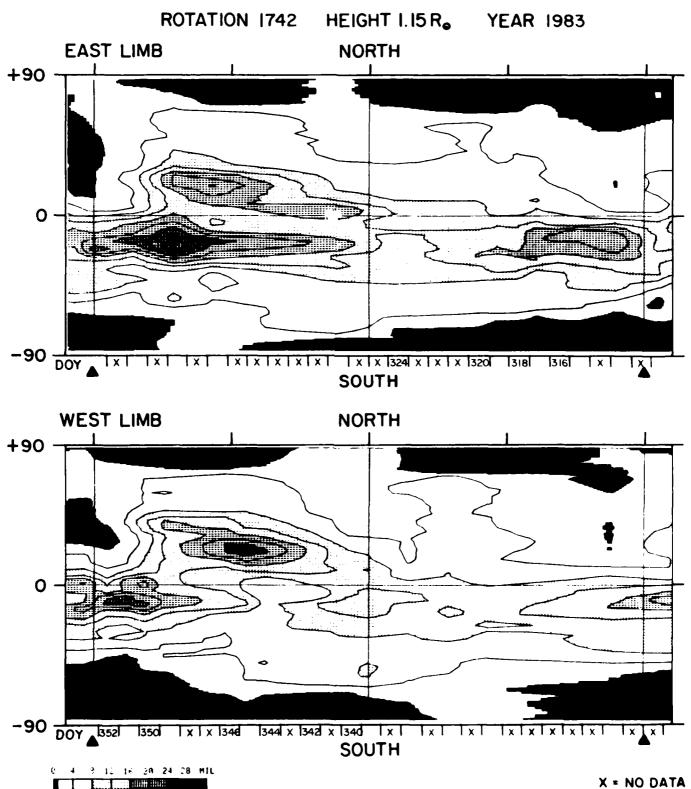




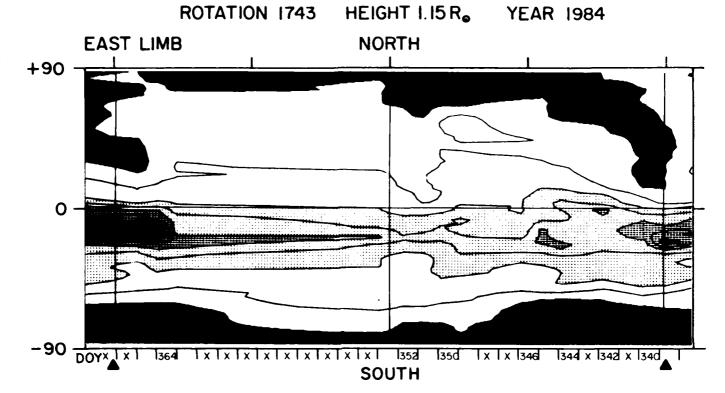
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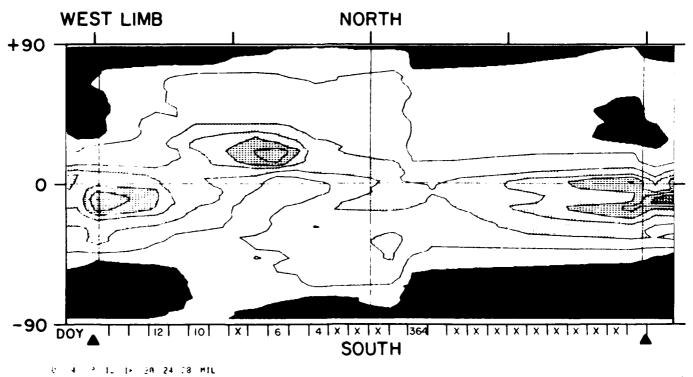




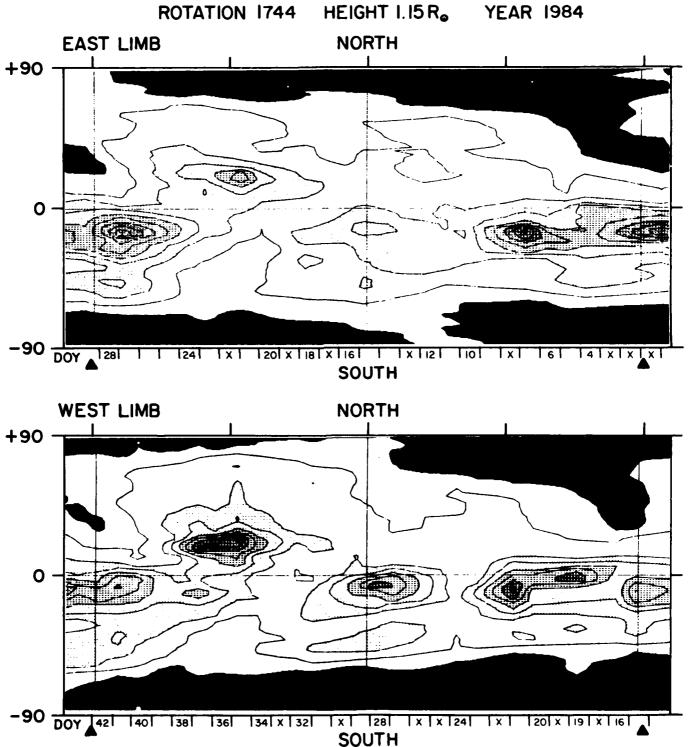


Fe XIV, 5303 & CORONAL PHOTOMETER





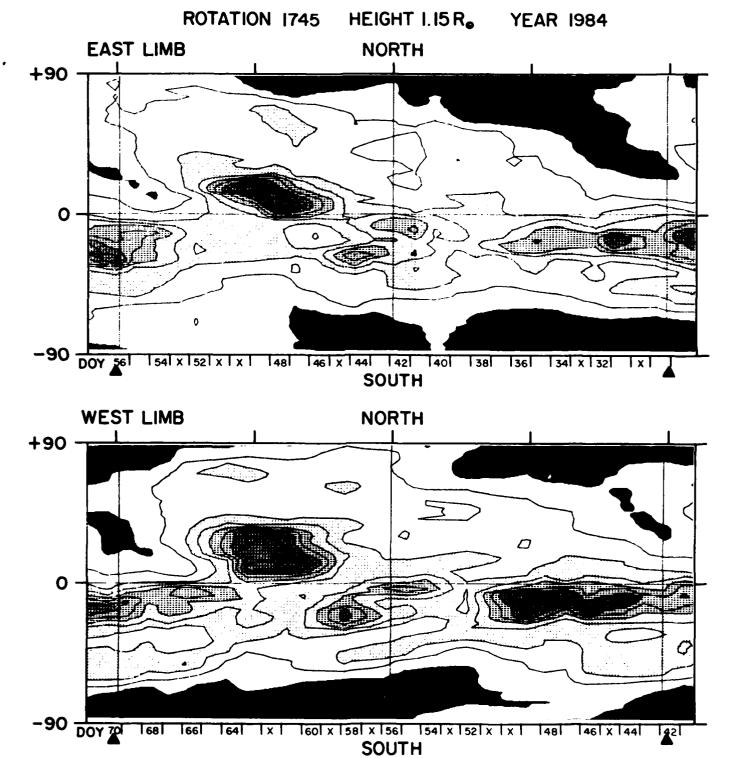
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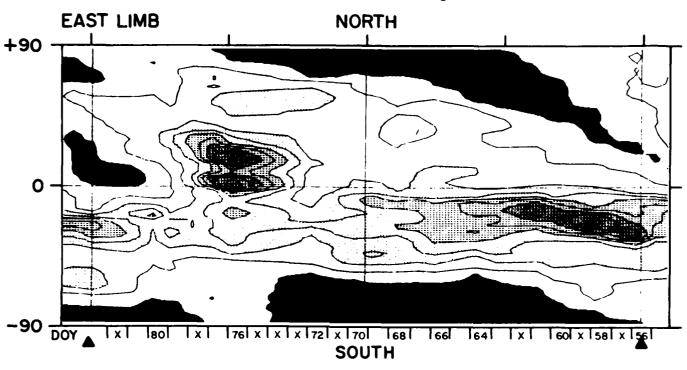
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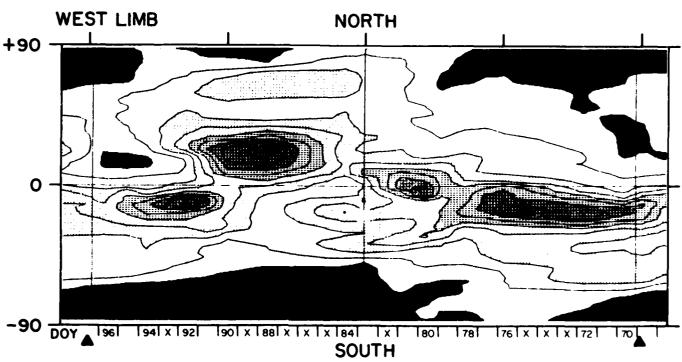
Fe XIV, 5303 A CORONAL PHOTOMETER



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Fe XIV, 5303 & CORONAL PHOTOMETER
ROTATION 1746 HEIGHT 1.15 R. YEAR 1984

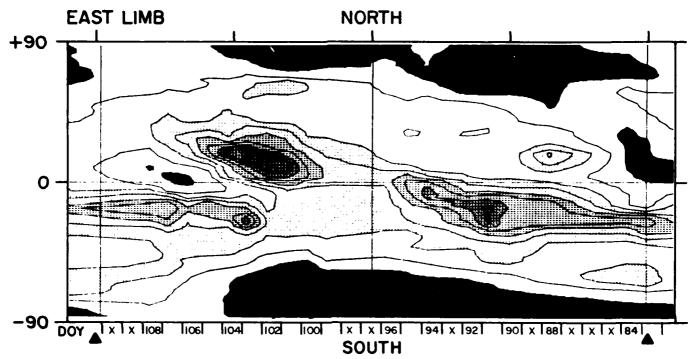


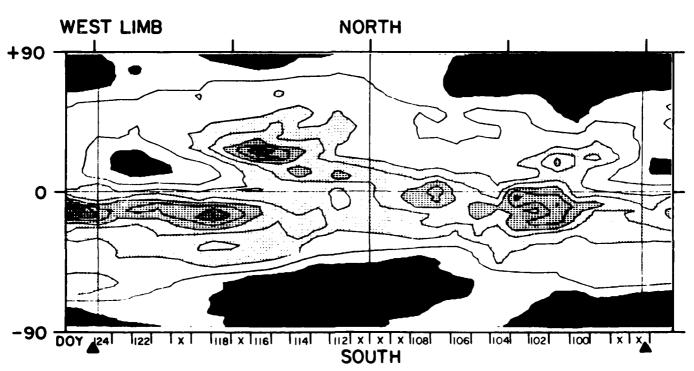


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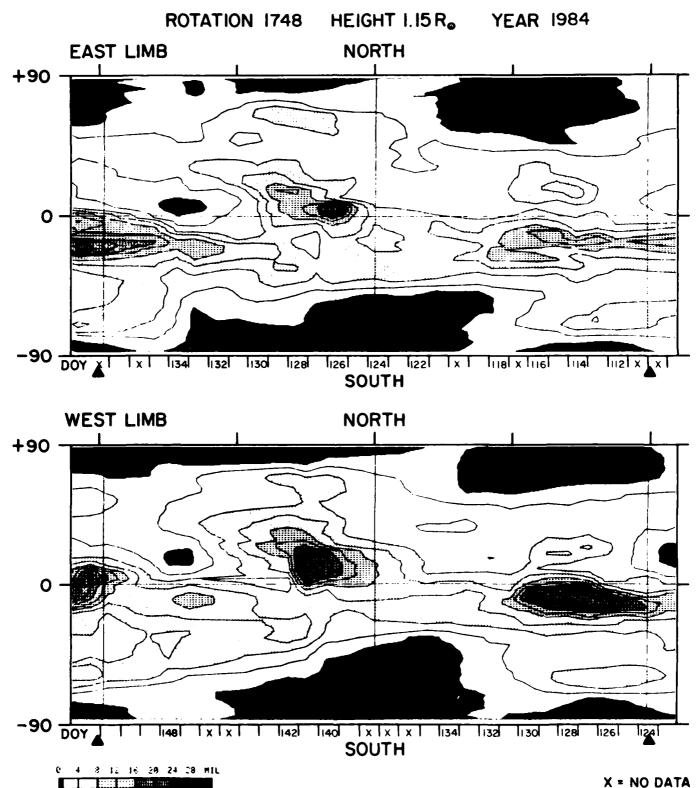
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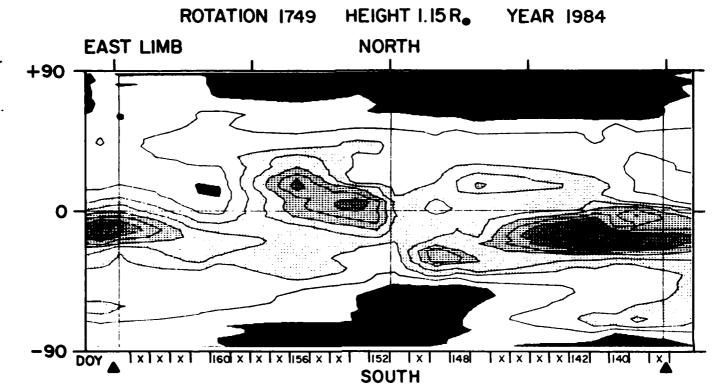


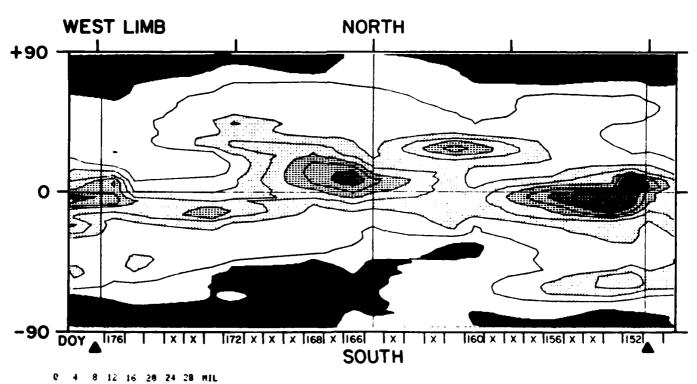


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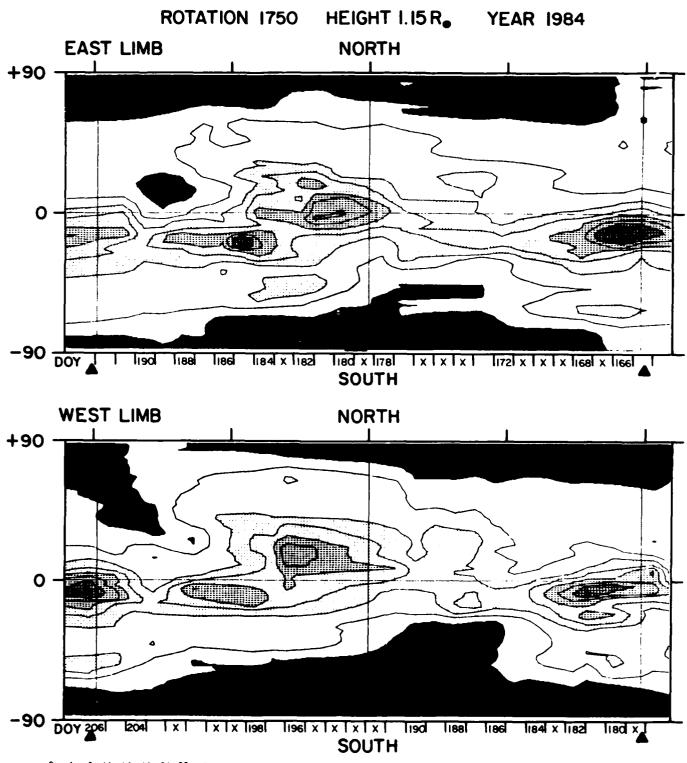


Fe XIV, 5303 Å CORONAL PHOTOMETER

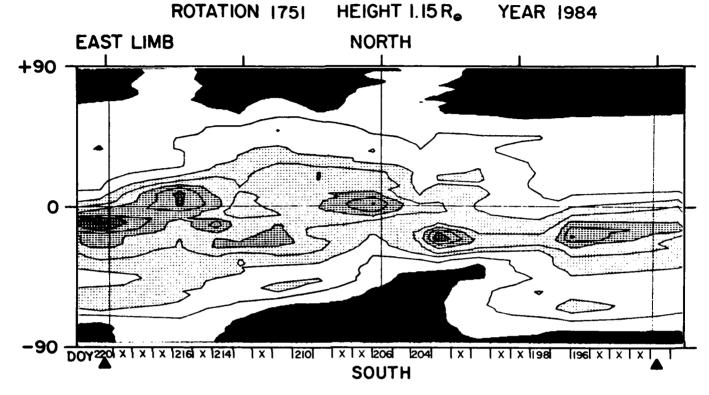


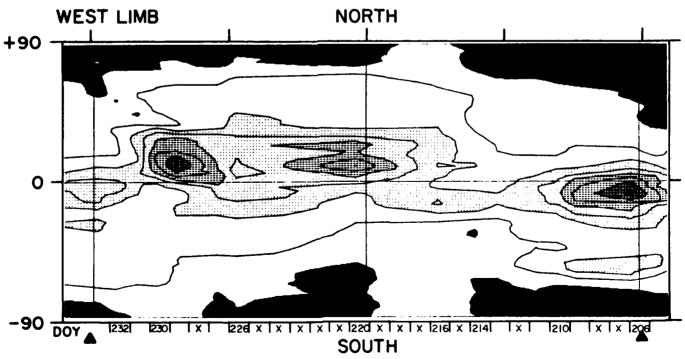


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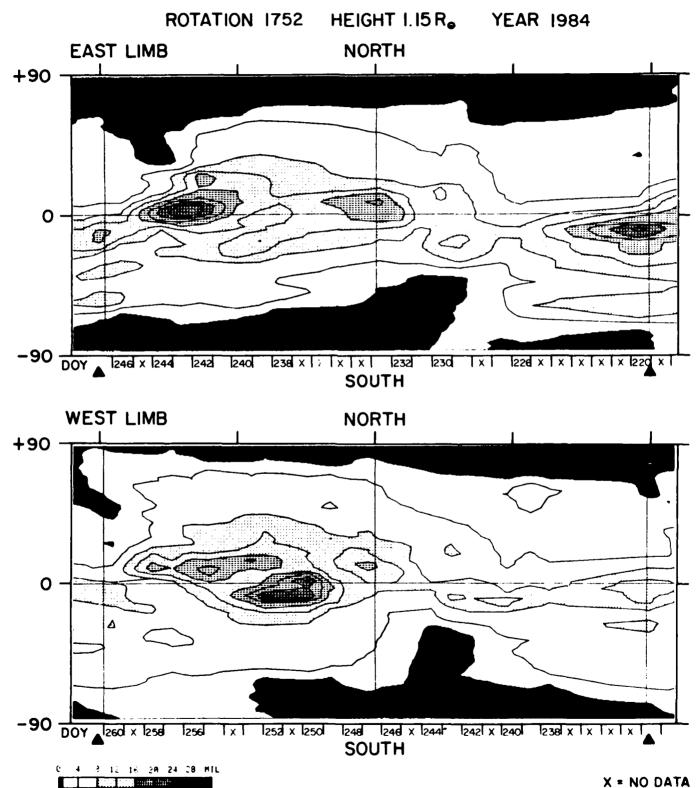
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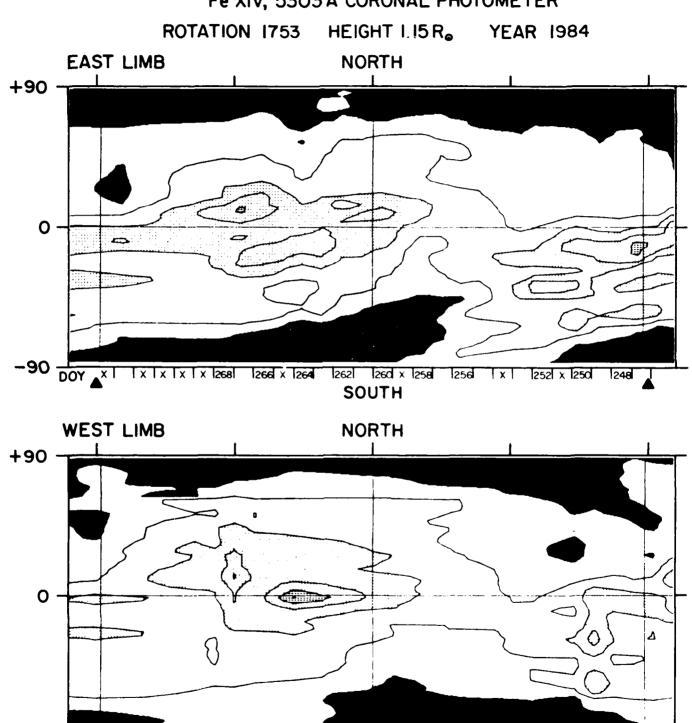


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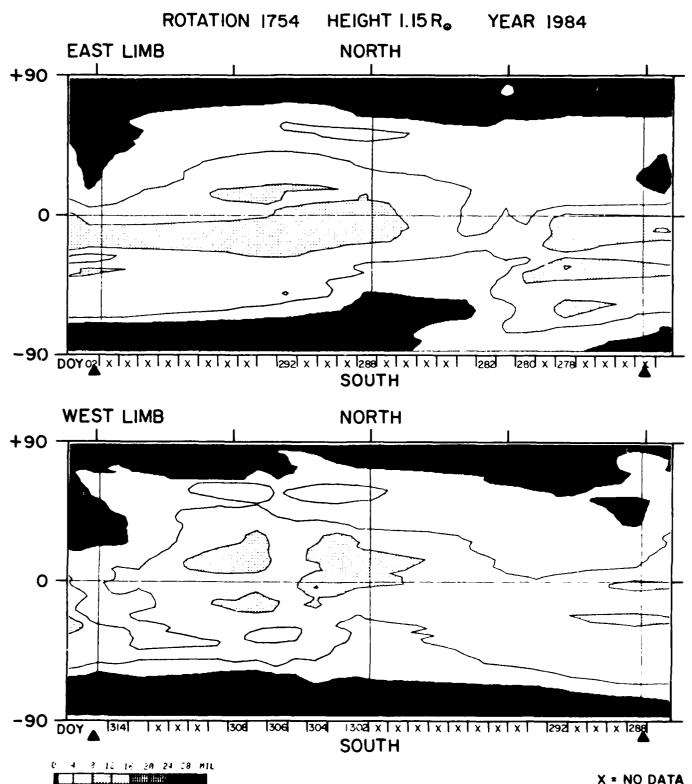


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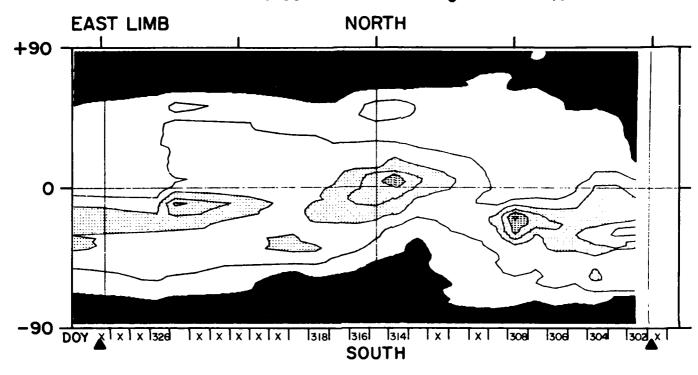


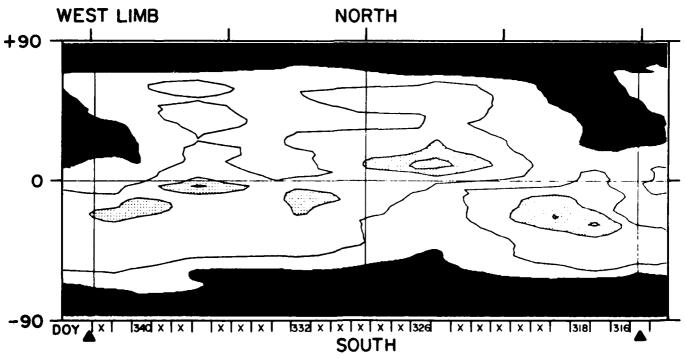
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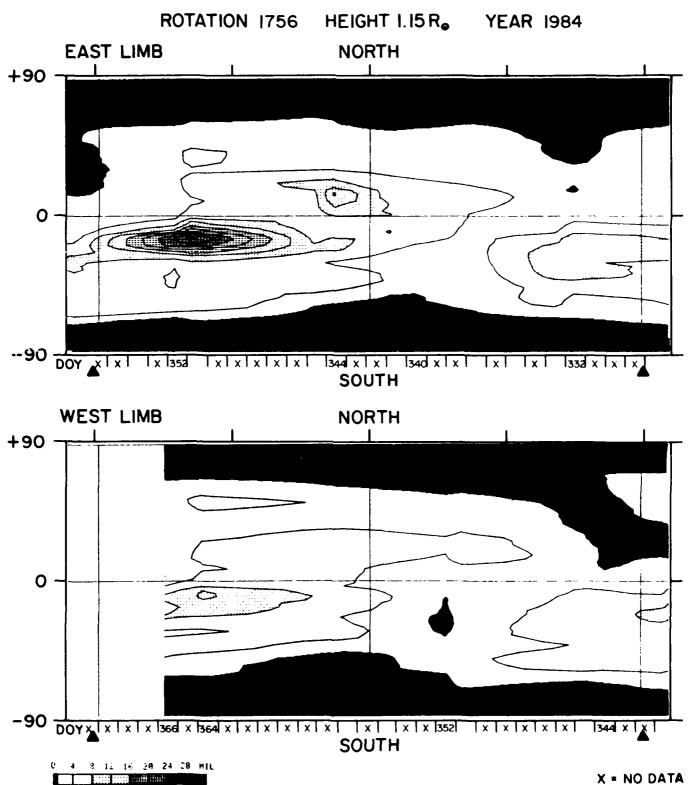
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ROTATION 1755 HEIGHT I.15 R. YEAR 1984



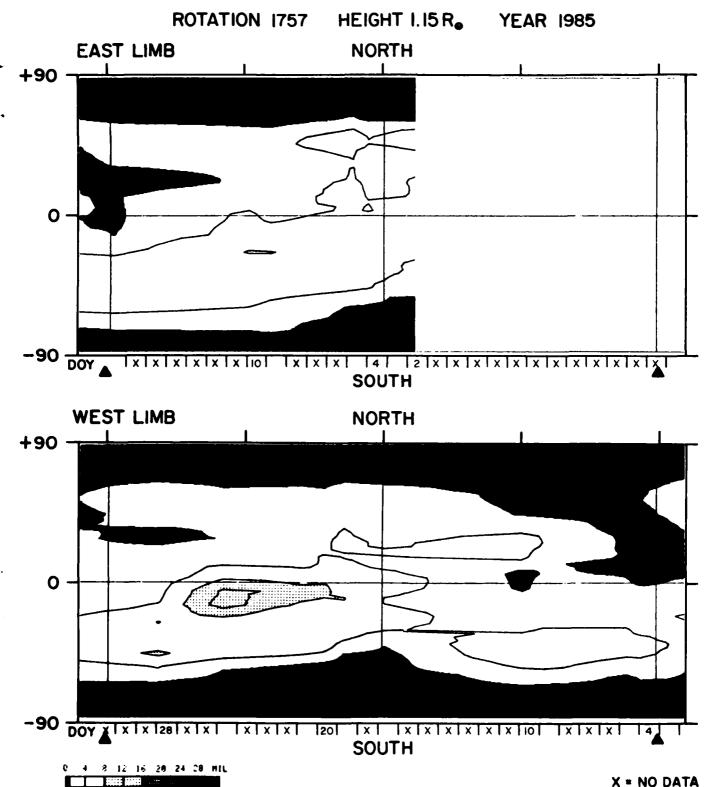


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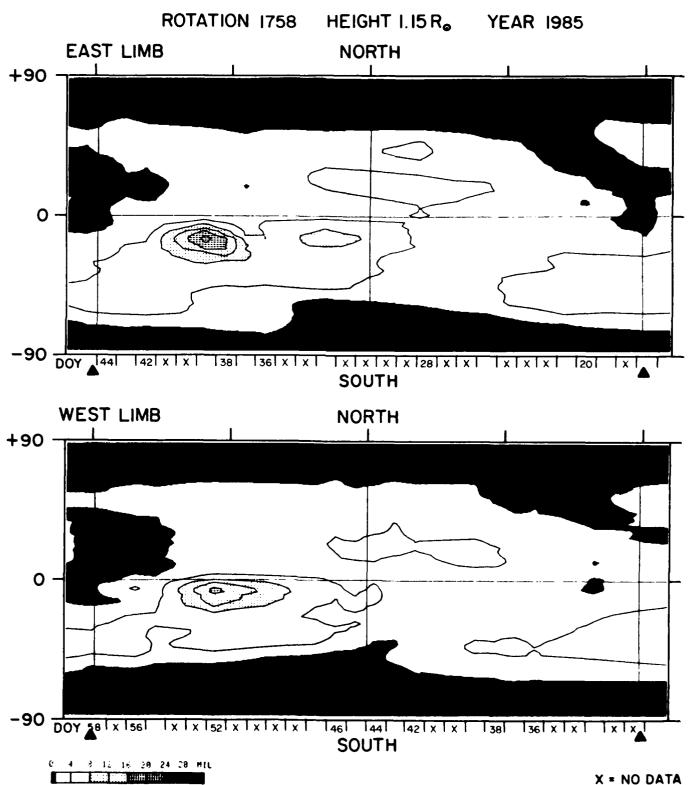
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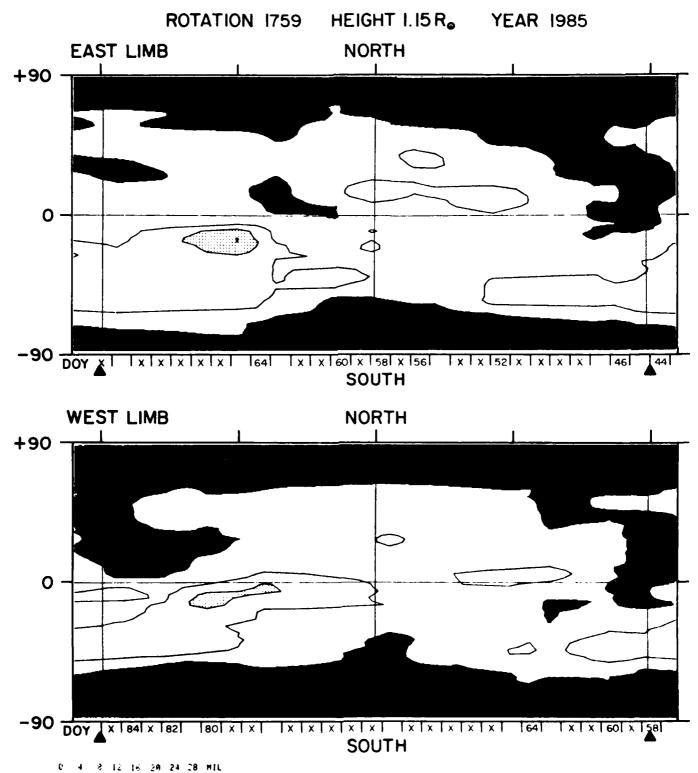
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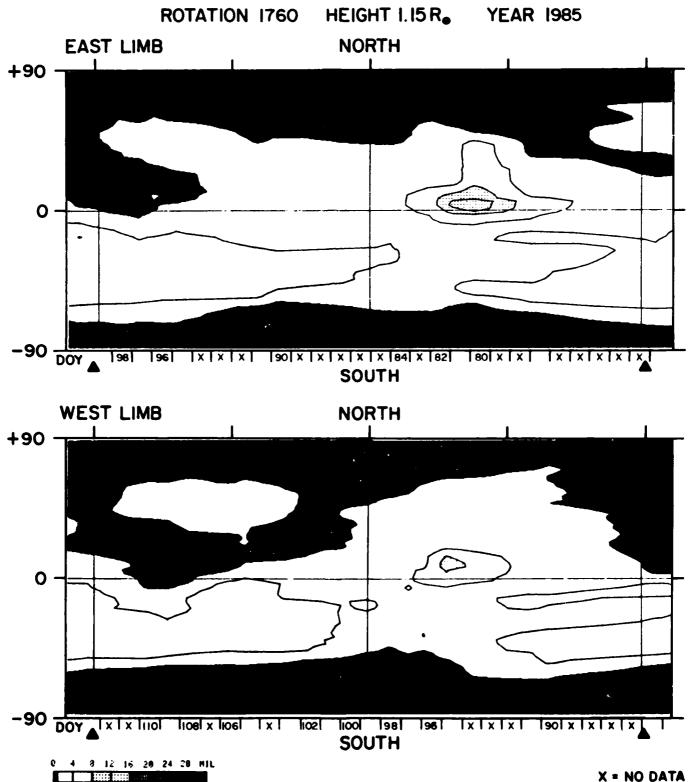
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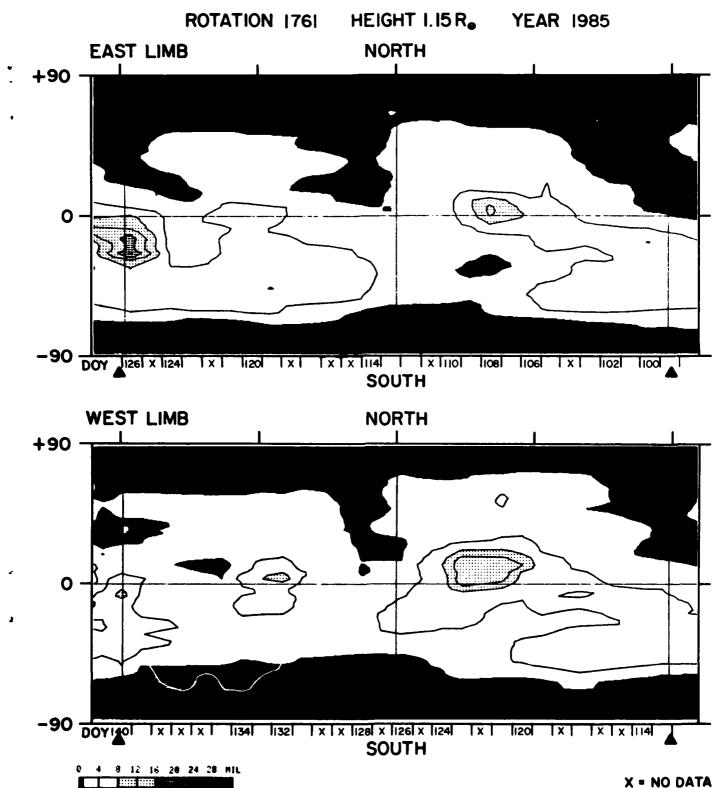
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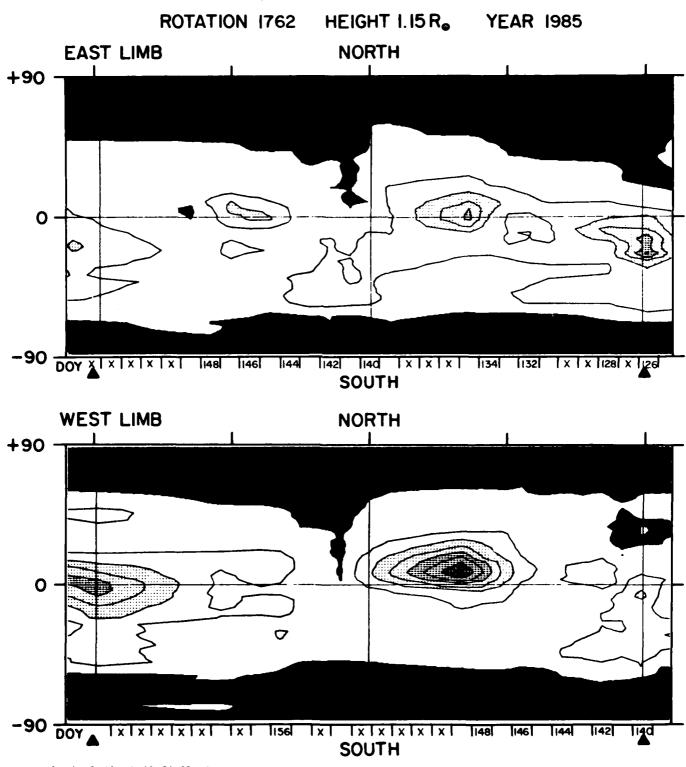
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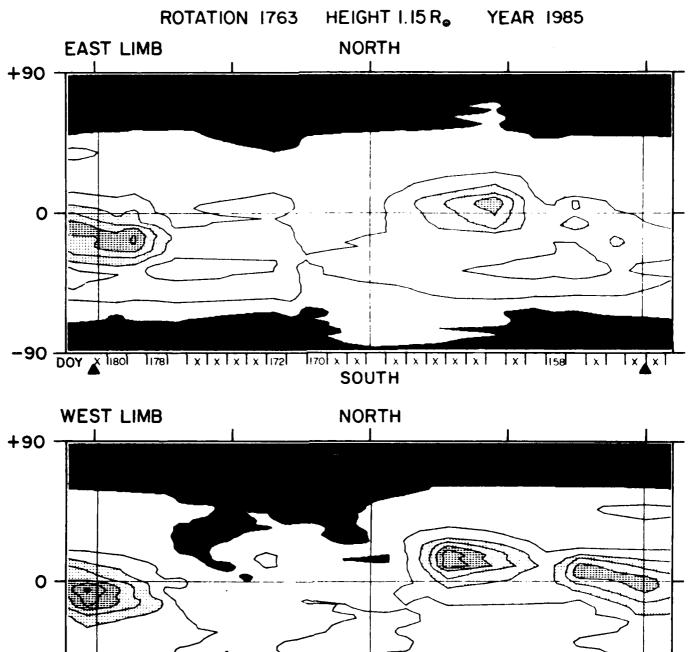
Fe XIV, 5303 A CORONAL PHOTOMETER



Fe XIV, 5303 Å CORONAL PHOTOMETER



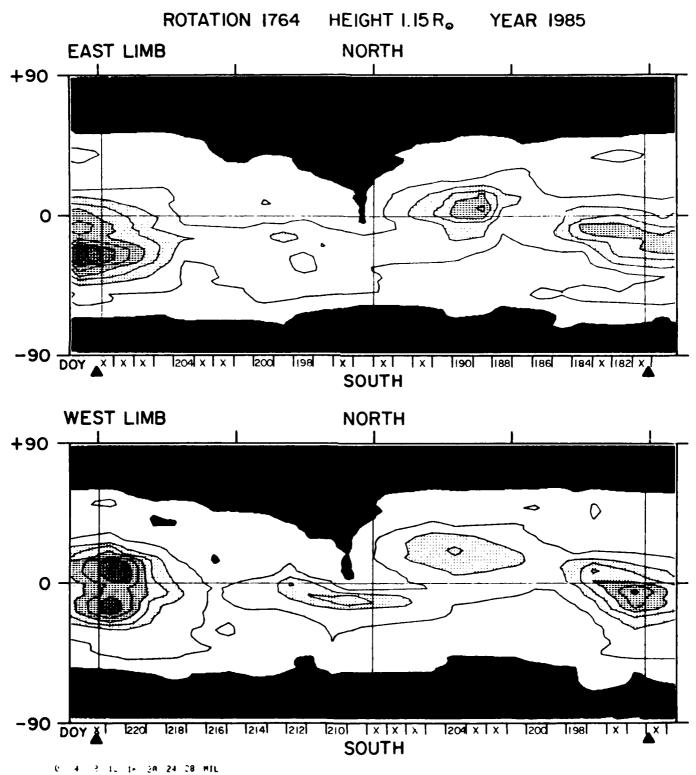
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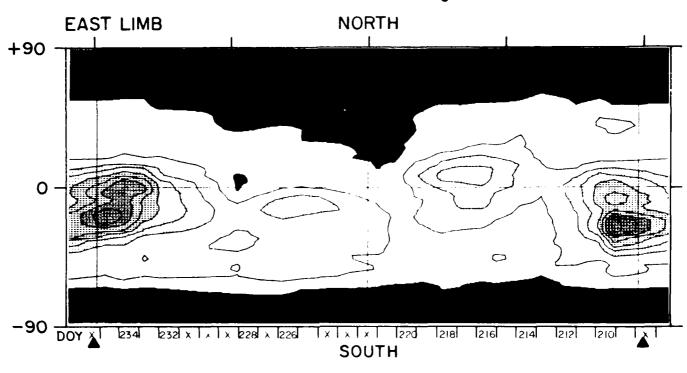
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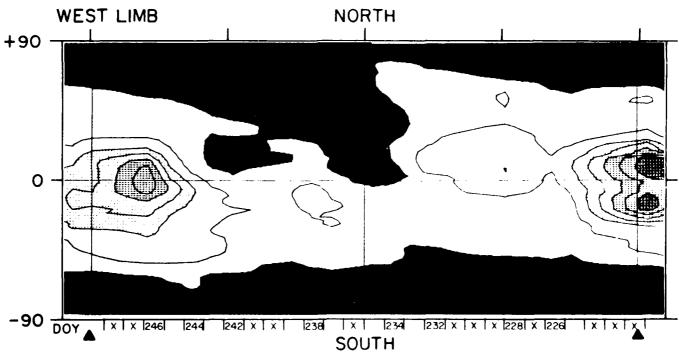
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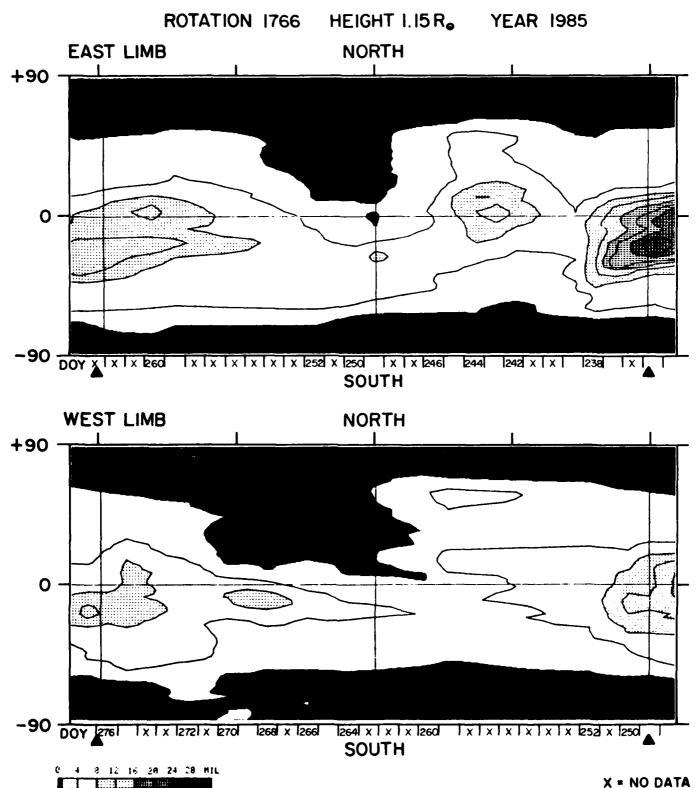
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ROTATION 1765 HEIGHT 1.15 R. YEAR 1985



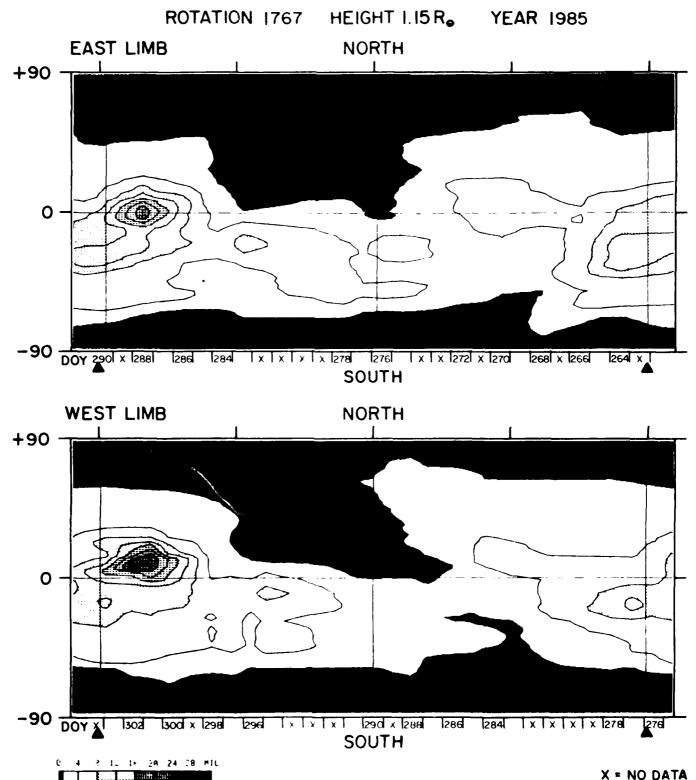


V 4 2 1 1 → 20 24 18 MIL

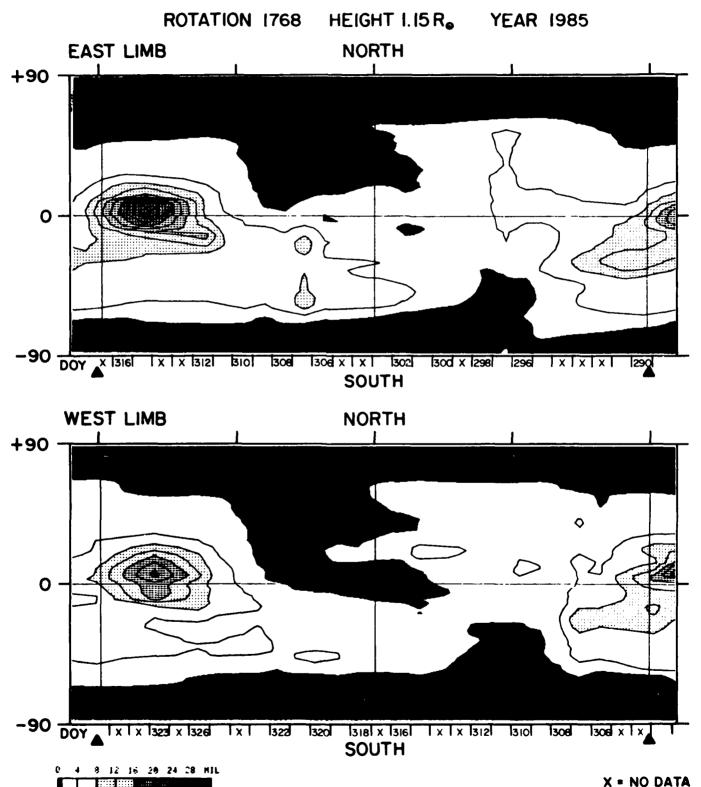
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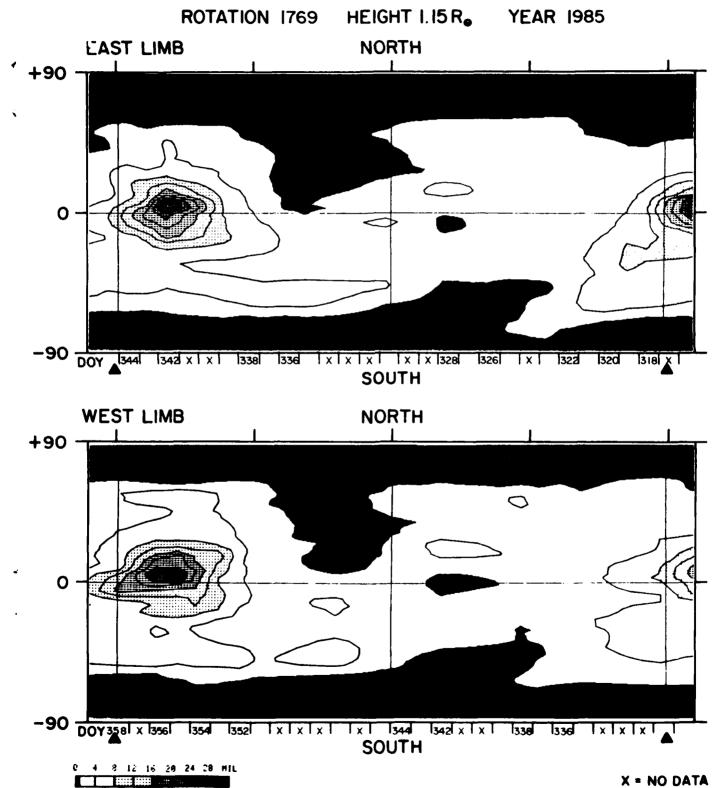
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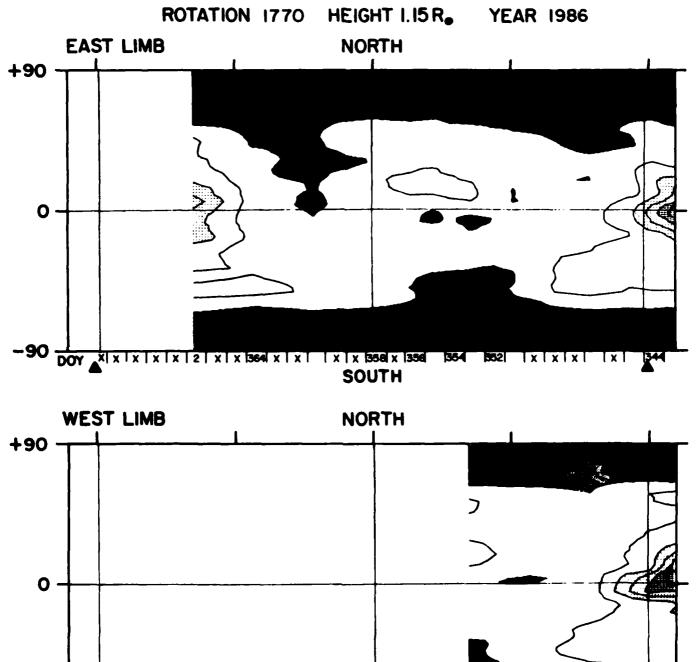
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Fe XIV, 5303 Å CORONAL PHOTOMETER



Fe XIV, 5303 Å CORONAL PHOTOMETER



SOUTH

0 4 8 12 16 20 24 28 MIL

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